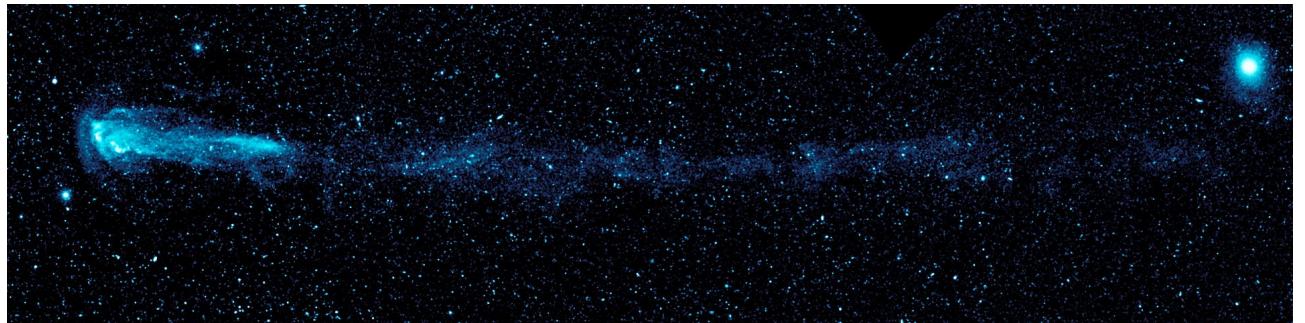


Spectroscopic Atlas of Mira (o Ceti)

3787Å-5750Å & 5829Å-8817Å



(source: www.nasa.gov/mission_pages/galex/20070815/a.html)

2014
David Reck
Hamburger Sternwarte

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A: Blue spectrum (3787Å-5750Å)

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1. Introduction:

This spectroscopic atlas about Mira (o Ceti) covers the wavelengths 3787Å-5750Å and 5829Å-8817Å.

The spectra of Mira were observed with a spectral resolution of approximately $R = \lambda/\Delta\lambda = 20000$ with the TIGRE Telescope in 2013.

TIGRE is operated by the Hamburger Sternwarte and the universities of Guanajuato and Liège at the La Luz observatory, near Guanajuato (Mexico).

1.1 The star Mira (o Ceti)

Mira, or Omicron Ceti, is a red giant star and forms a binary system with Mira B, or VZ Ceti.

The spectral type of Mira is M1e-M9e (Simbad), which is changing during its variability period.

Mira has a mean diameter of approximately 400 solar diameters (550 Mio. km).

The distance to Mira amounts to 92 ± 10 pc or 300 ± 33 ly.

Its radial velocity is $+63.8 \pm 0.9$ km/s.

Mira has a tail, in appearance similar to that of a comet.

However, Mira's tail extends over 13 light years in length.

1.2 Spectra used for the catalog

The chosen two dates are 2013-08-08 (dotted line) and 2013-10-19 (continuous line).

I chose these two dates because the first (2013-08-08) is the earliest available of Mira and the second (2013-10-19) is the latest spectrum where the noise in the blue spectral range is so low, that a comparison with the first spectrum still allows to study most of the spectral lines visible in the first spectrum.

1.3 Normalization of the spectra

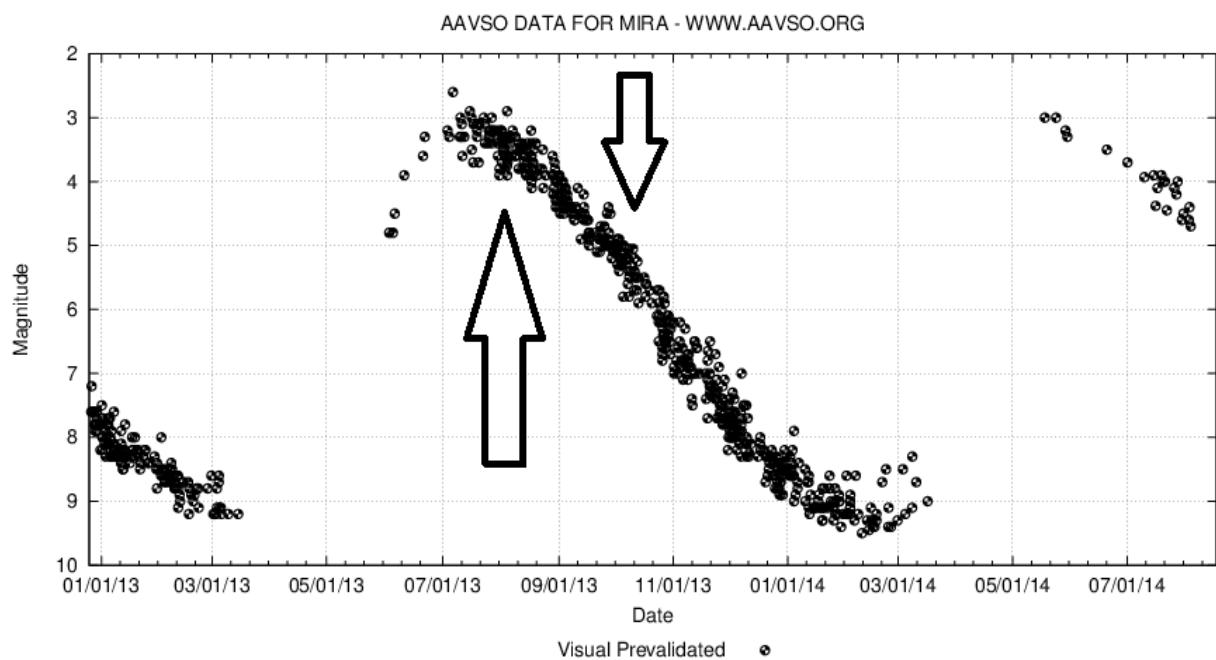
The spectra have been individually normalized on each plotted 60Å – wide

subinterval by multiplication with a constant.

The normalization was performed using the package onedspec of the program IRAF (version: 2.16).

This process was necessary to allow a common representation of both datasets in one graph.

2. Cycle of Mira (o Ceti):



(source: www.aavso.org/lcg)

The two arrows show the observing dates of the spectra in this atlas.

Mira changes its visual brightness between typically V=3 mag and V=9 mag with a period of approximately 332 days (Woodruff et al. 2008, ApJ 673, 418) (for more information see www.aavso.org/vsots-mira)

3. Corrections (barycentric and radial velocity):

The spectra are corrected for barycentric shift, but not for the radial velocity ($+63.8 \pm 0.9$ km/s.)

For example the shift at 7000Å, caused by the radial velocity, amounts to about 1.5Å in the red spectral range.

4. Pronounced changes in Mira's spectrum

The following is a list of specific lines, that markedly change their amplitude, between the two observing dates.

- 3797.514 Å Fe I
- 3852.572 Å Fe I
- 3829 Å Mg I ; 8332 Å Mg I ; 8338 Å Mg I
- 3905.523 Å Si I
- 3906.479 Å Fe I
- 3938.560 Å Mg I
- 4030.76 Å Mn I
- 4033.07 Å Mn I
- 4216.183 Å Fe I
- 4512 Å ?
- 4573 Å ?
- 4839.544 Å Fe I
- 4969 Å ?
- 8413 Å ?
- 8420 Å ?
- 8429 Å ?
- 8436 Å ?

4.1 H-Balmer Series

Most of the visible H-Balmer lines have a huge amplitude which is not covered by the given scaling of the y-axis, which was chosen in order to make the weaker lines well visible. The following list gives the amplitudes of the strongest Balmer lines relative to the surrounding quasi-continuum. They all appear in emission.

- H10 amplitude: 27.9
- H9 amplitude: 25.1
- H8 amplitude: 35.7
- H ϵ amplitude: 4.2
- H δ amplitude : 46.3
- H γ amplitude : 46.4
- H β visible in the plot
- H α amplitude: 2.4

5. Telluric lines

Below each 60Å interval I show a telluric spectrum for comparision, if the interval contains significant telluric lines.

The telluric spectrum is a synthetic spectrum computed using the LBRTM package (Clough et al., 1992, *S Geophysical Research*).

A good example of telluric contamination can be seen in the wavelength range between 8100Å and 8360Å.

6. P-Cygni lines

Only two of the Ca II infrared triplet lines show a pronounced P-Cygni shape. These are: 8664Å and 8544Å

7. References

Spectroscopic Journal of Mira CETI, 1943-1952 ApJS, vol 1 p.39 (1954), A. Joy. In the atlas listed by the abbreviation (Joy).

NIST http://physics.nist.gov/PhysRefData/ASD/lines_form.html

Atomic Spectra Database (version 5.1)

National Institut of Standards and Technology, Gaithersburg, MD.

In the atlas listed by the abbreviation (NIST).

VALD <http://vald.inasan.ru/~vald3/php/vald.php>

This work has made use of the VALD database, operated at Uppsala University, the Institute of Astronomy RAS in Moscow and the University of Vienna.

In the atlas listed by the abbreviation (VALD).

ApJ 498, 851-862, 1998, J. A. Valenti, N. Piskunov and C. M. Johns-Krull

In the atlas listed by the abbreviation (Valenti).

D. Walker, A spectroscopic Guide to Astronomical Objects and Terrestrial Light Sources, Version 5.0, 2014 http://www.ursusmajor.ch/downloads/spectroscopic-atlas-5_0-english.pdf

In the atlas listed by the abbreviation (Walker).

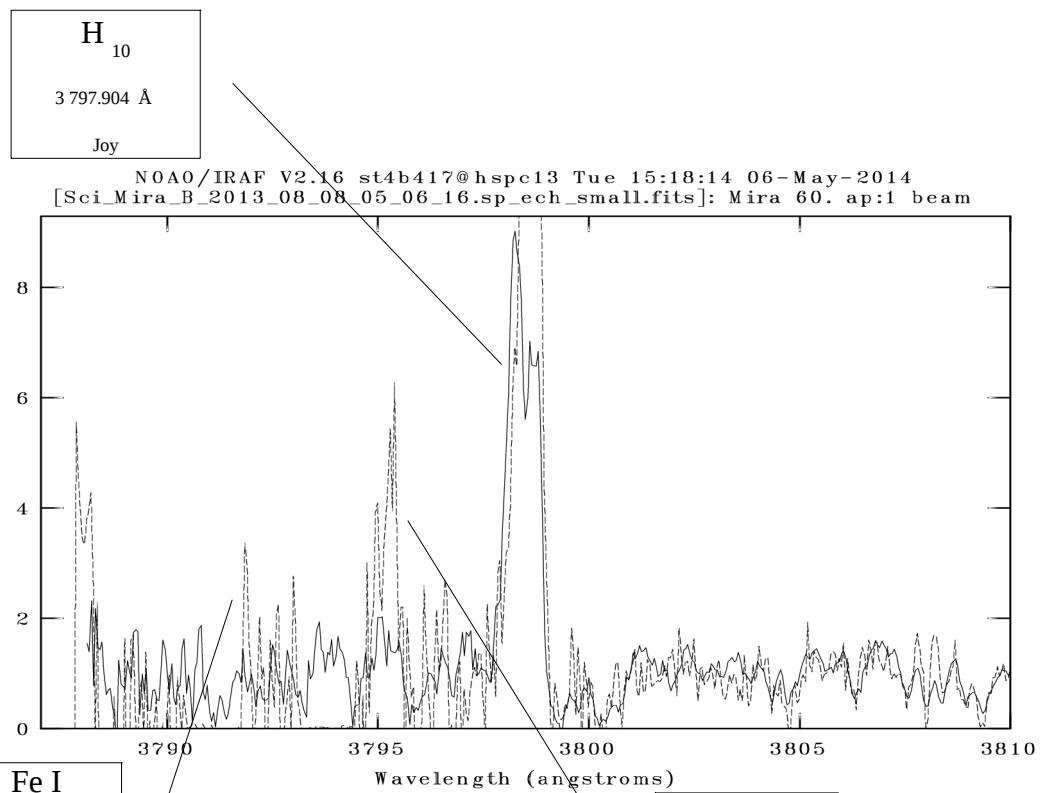
Wikipedia:

- [de.wikipedia.org/wiki/Mira_\(Stern\)](http://de.wikipedia.org/wiki/Mira_(Stern))
- en.wikipedia.org/wiki/Mira

9. Acknowledgements

I thank Dr. Uwe Wolter and Dr. Dieter Engels for supporting me during this work.

Appendix
A: Blue spectrum (3787Å-5750Å)



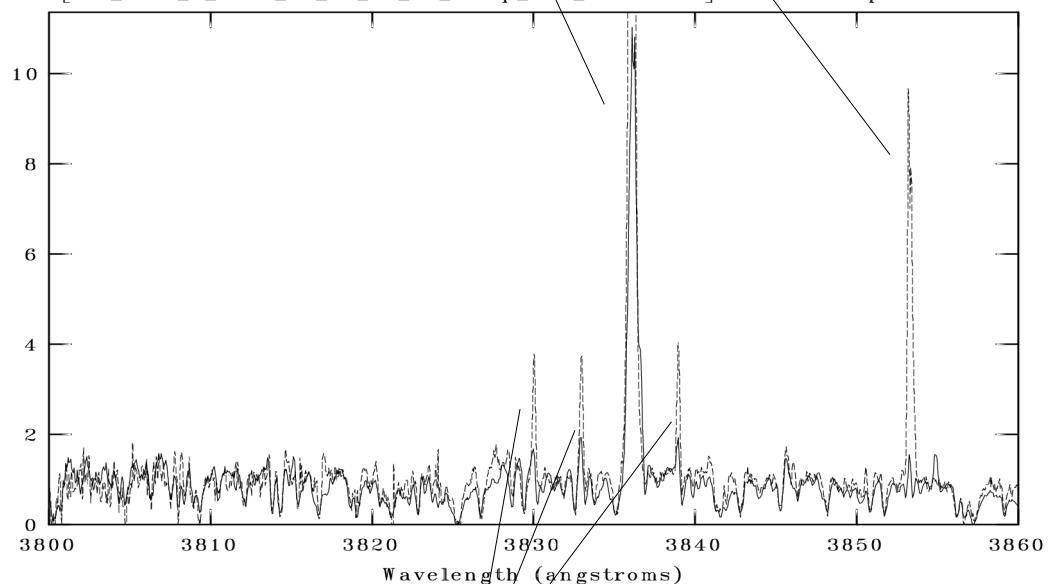
Fe I
 3 794.339 Å
 NIST

Fe I
 3 797.514 Å
 NIST

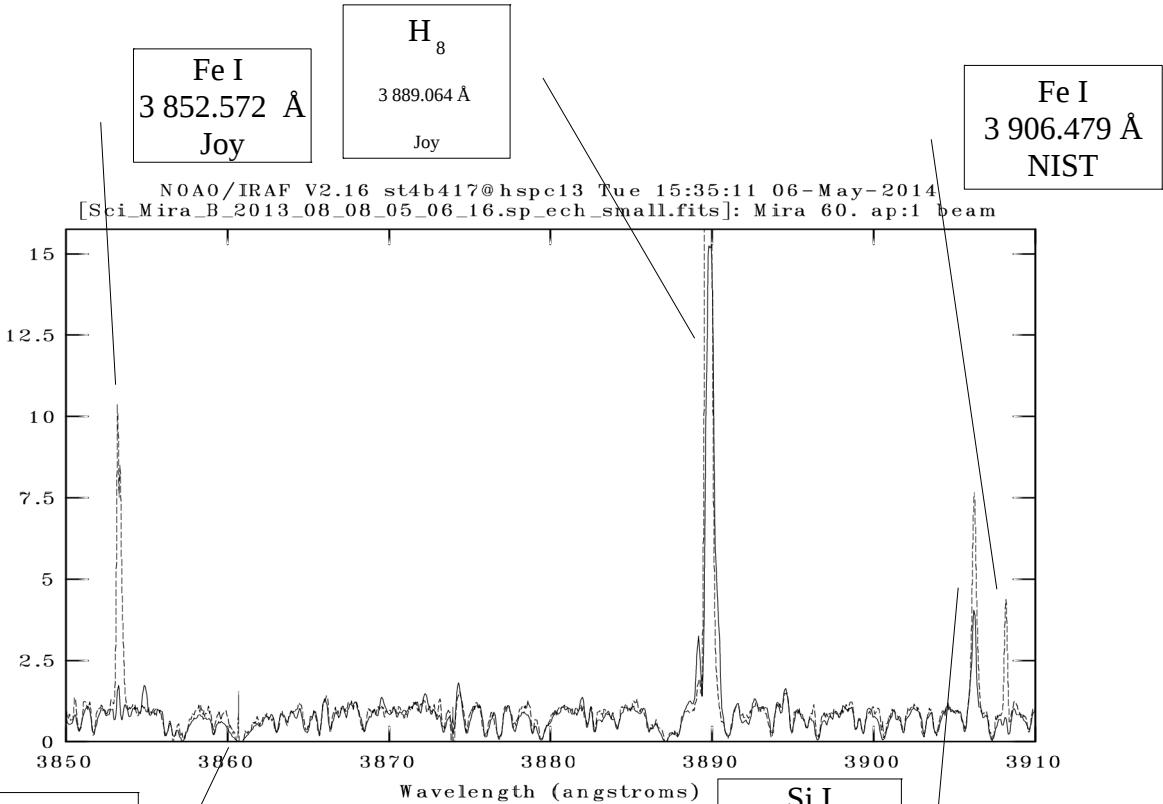
H₉
 3 835.390 Å
 NIST

Fe I
 3 852.572 Å
 Joy

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Mg I
 3829,32,38 Å
 Joy



Reduction artefact
Flux goes down to -30

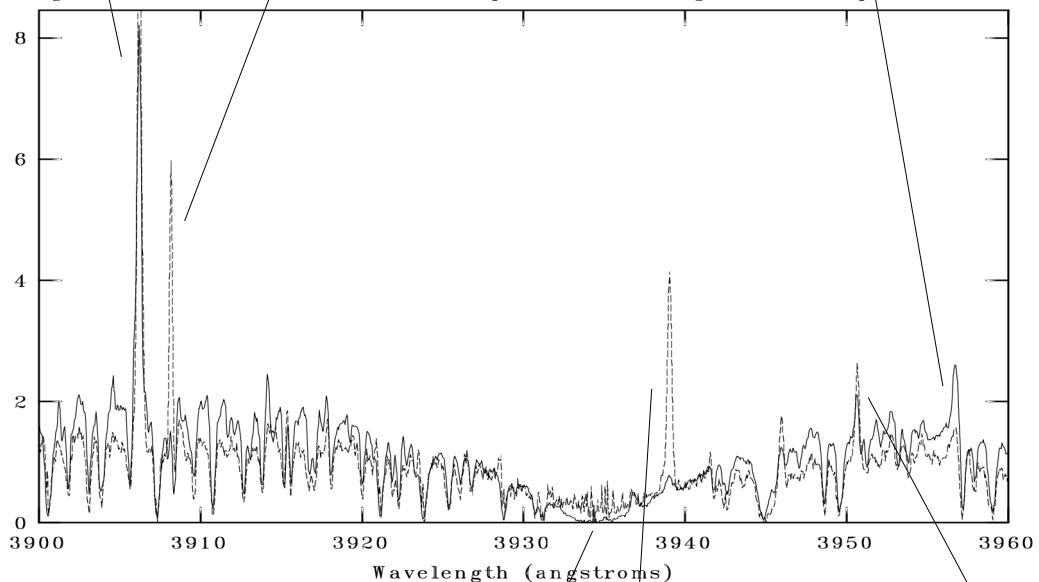
Si I
3905.523 Å
Joy

Si I
3905.523 Å
Joy

?

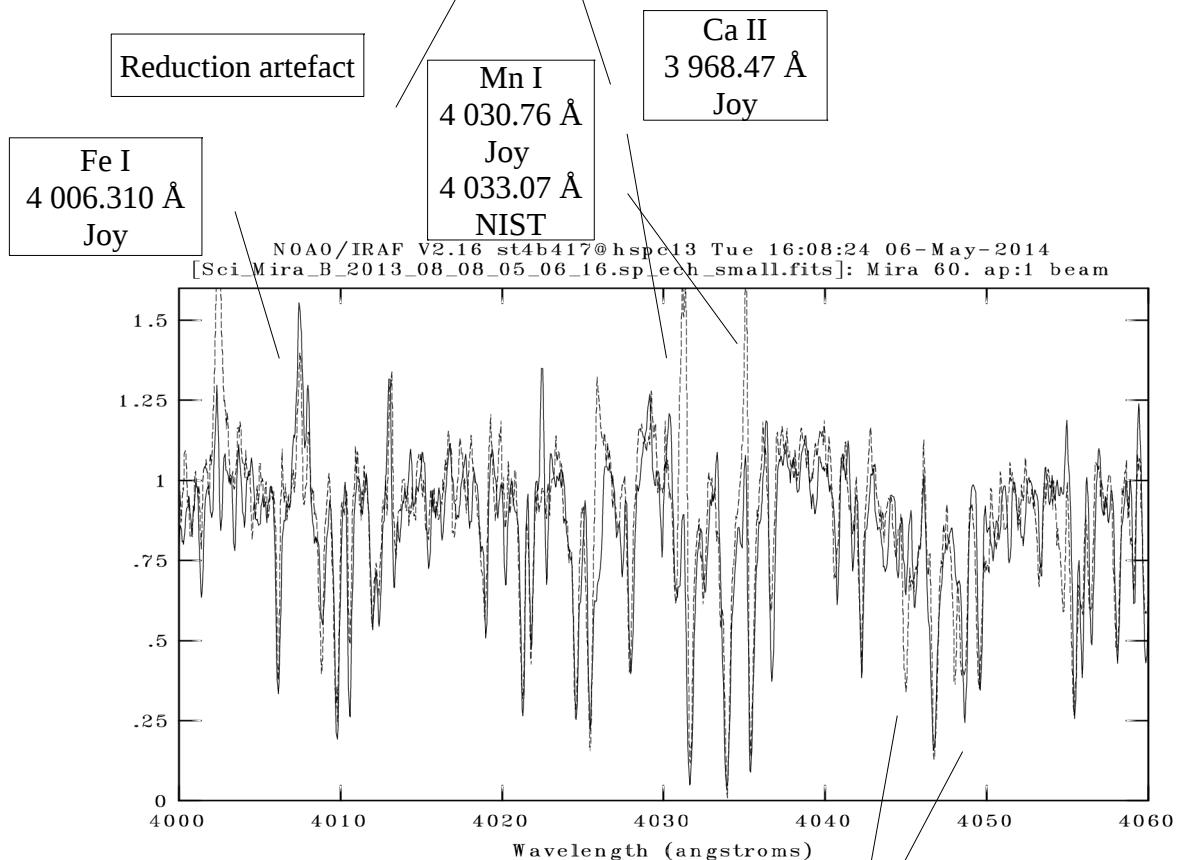
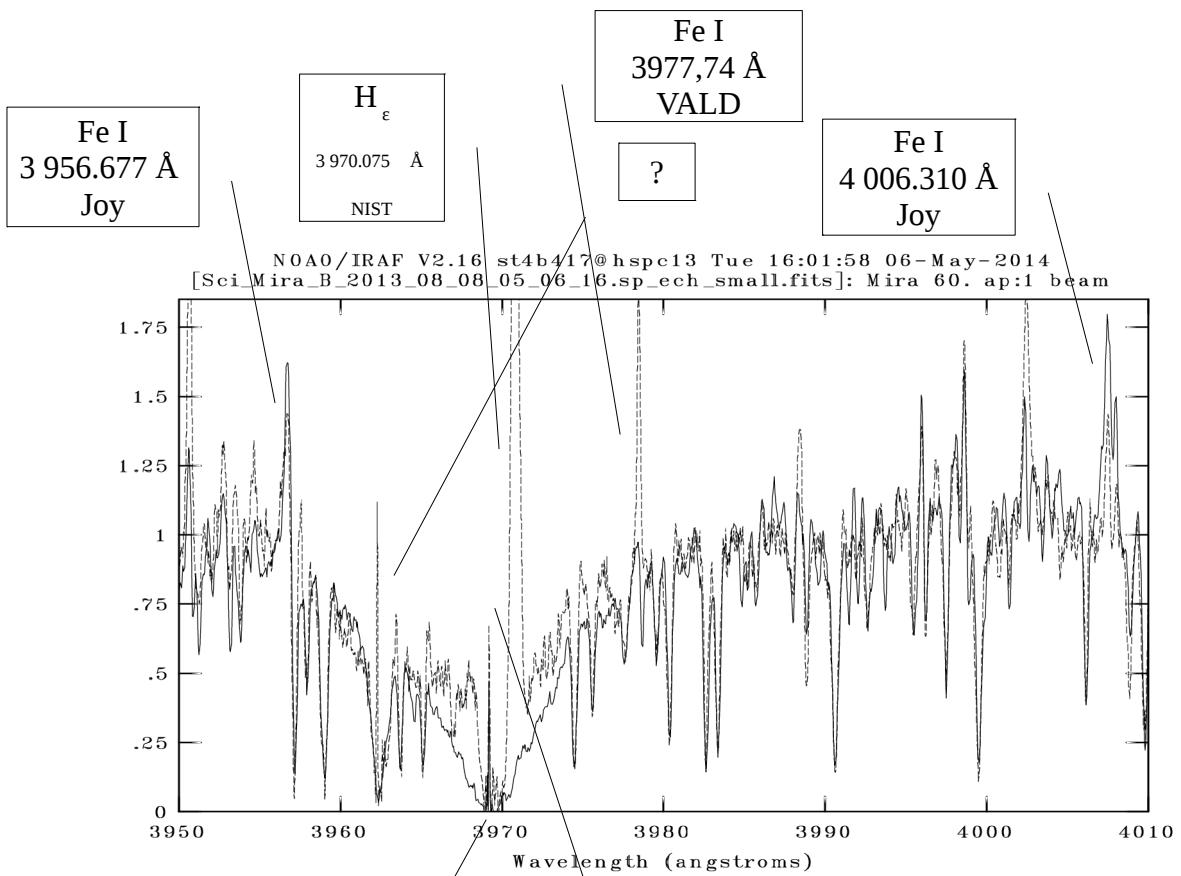
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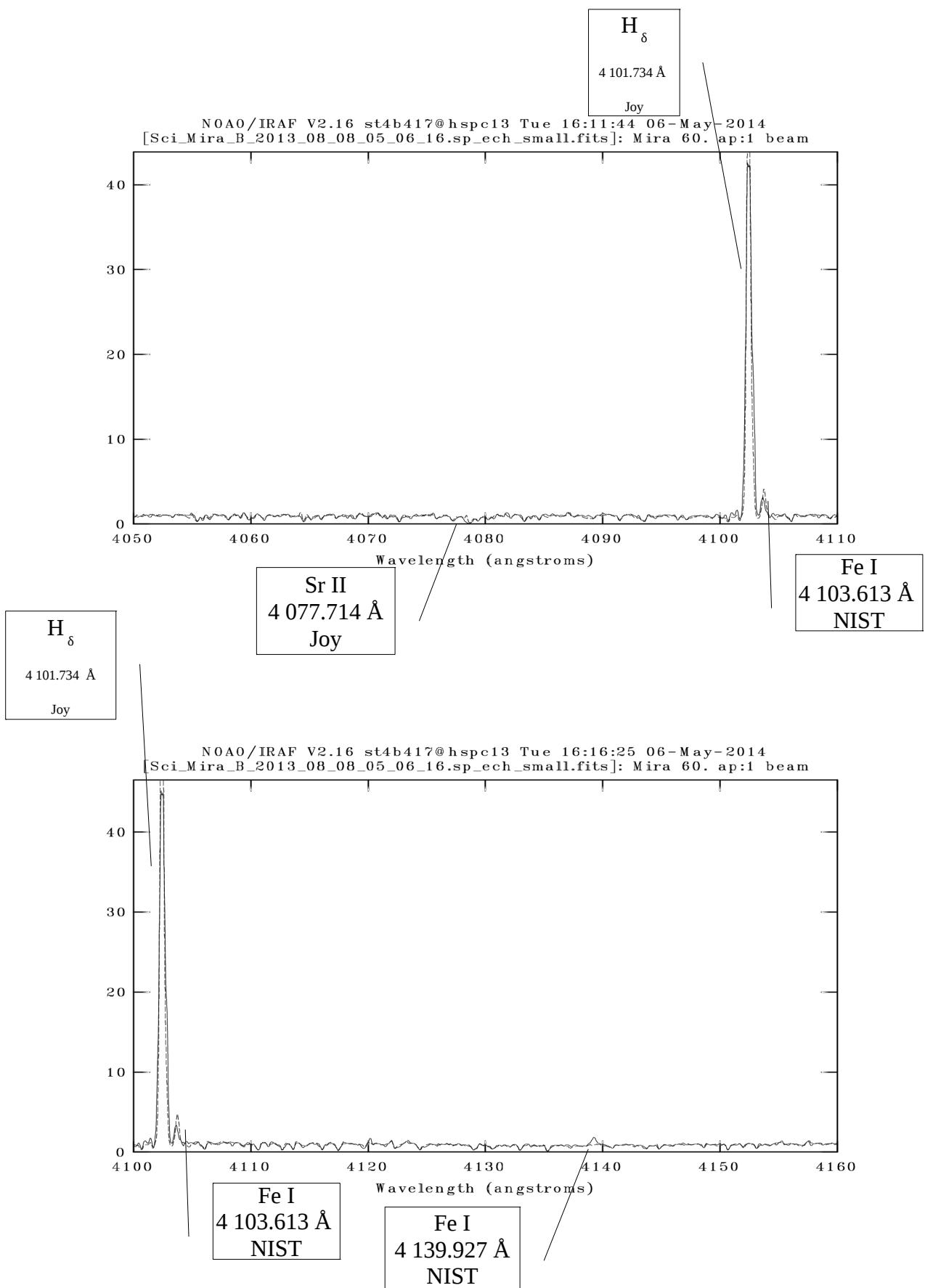


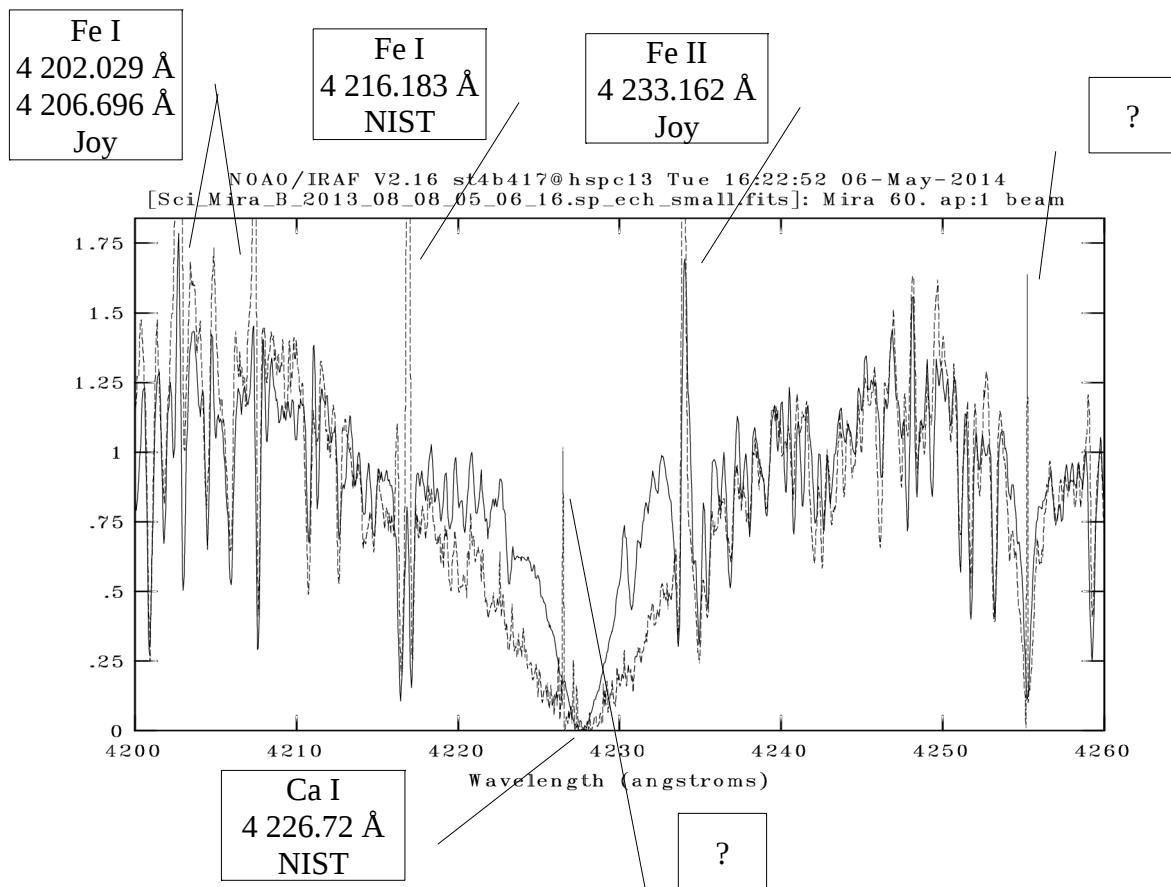
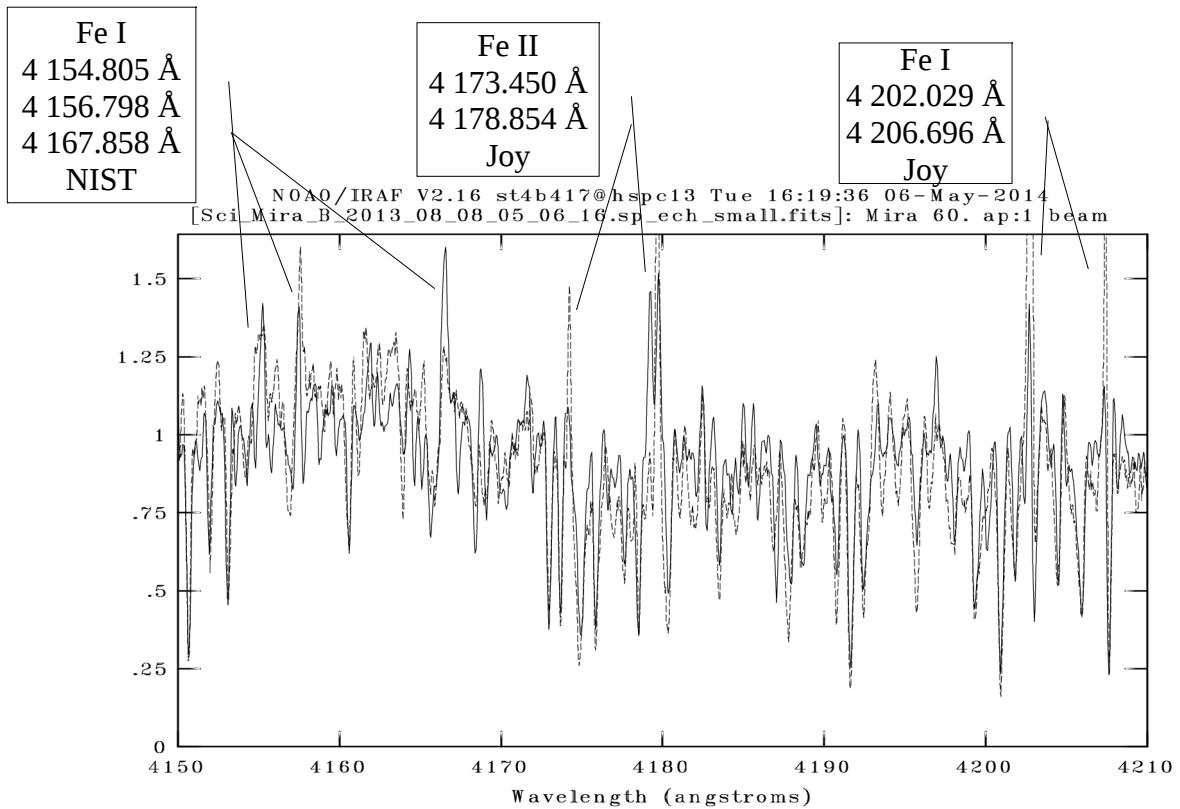
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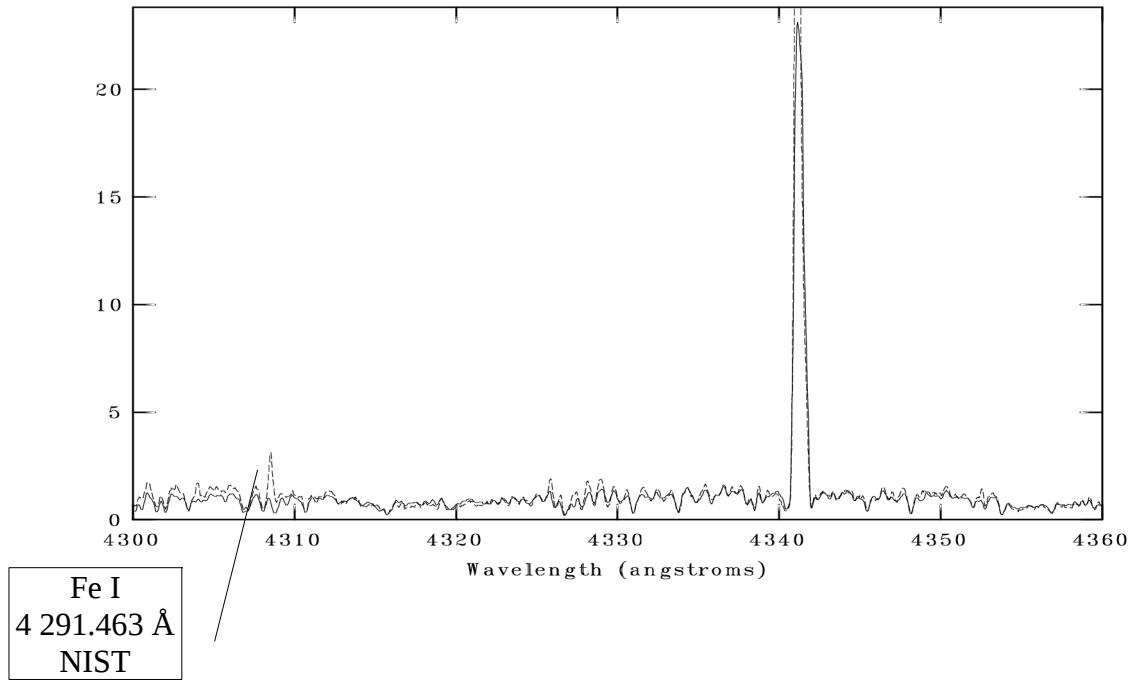
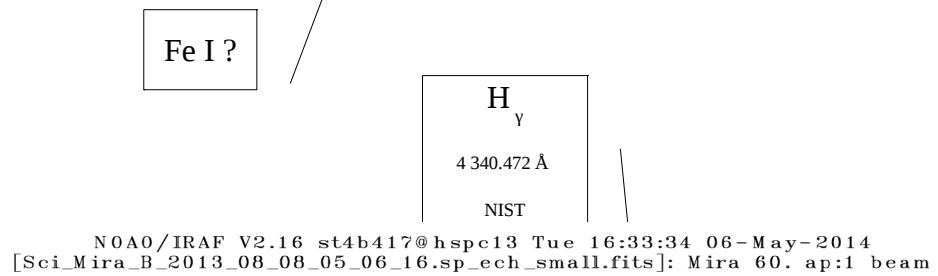
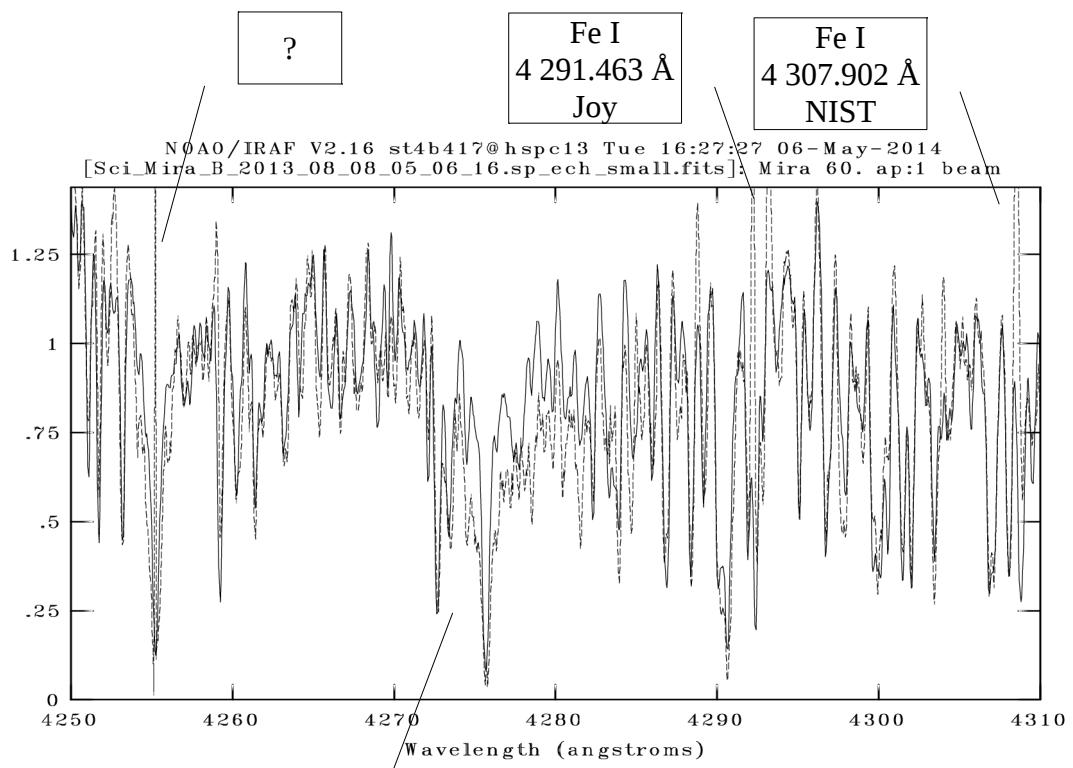
Fe I
3949.952 Å
Joy

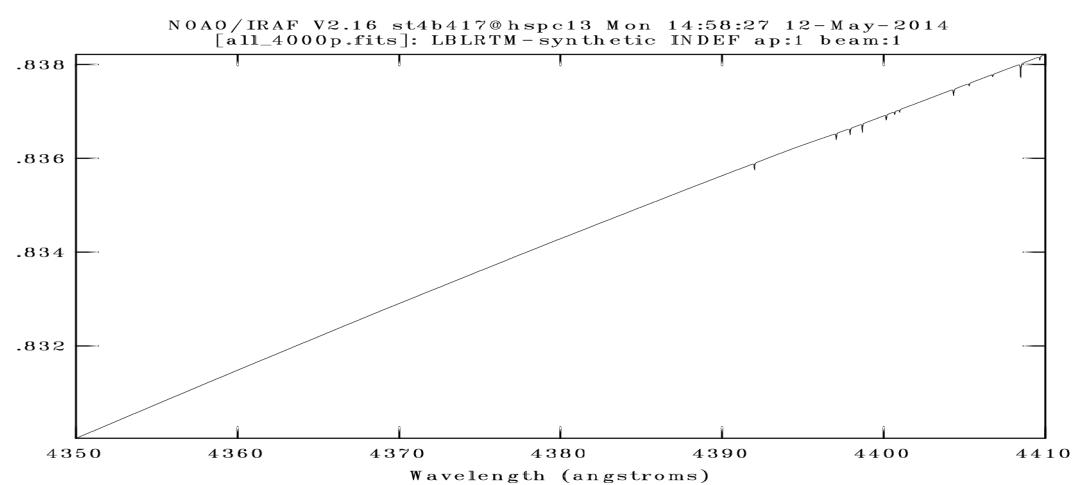
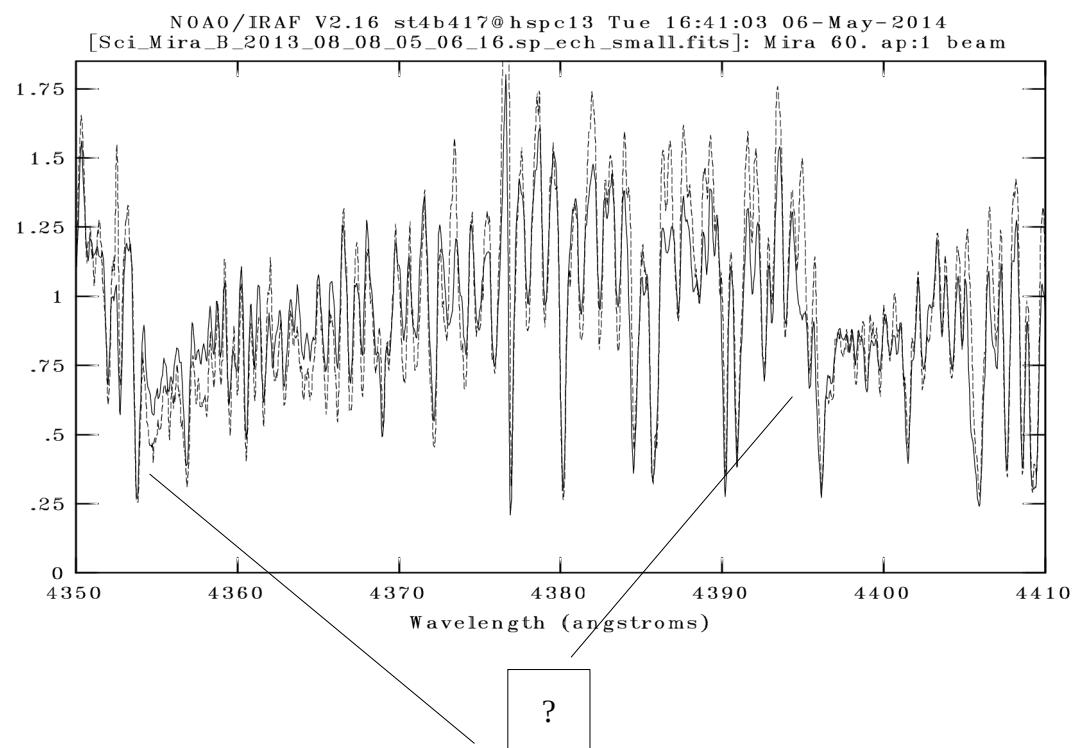


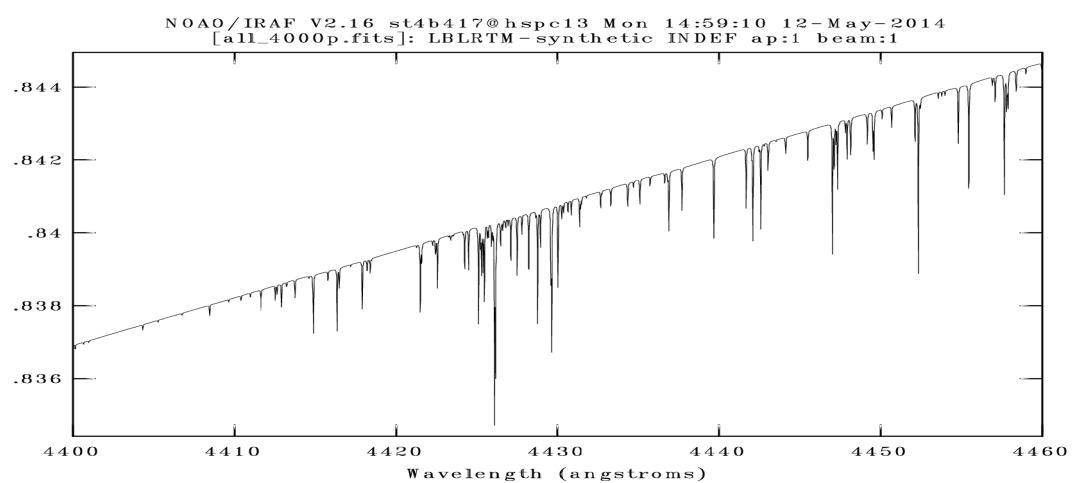
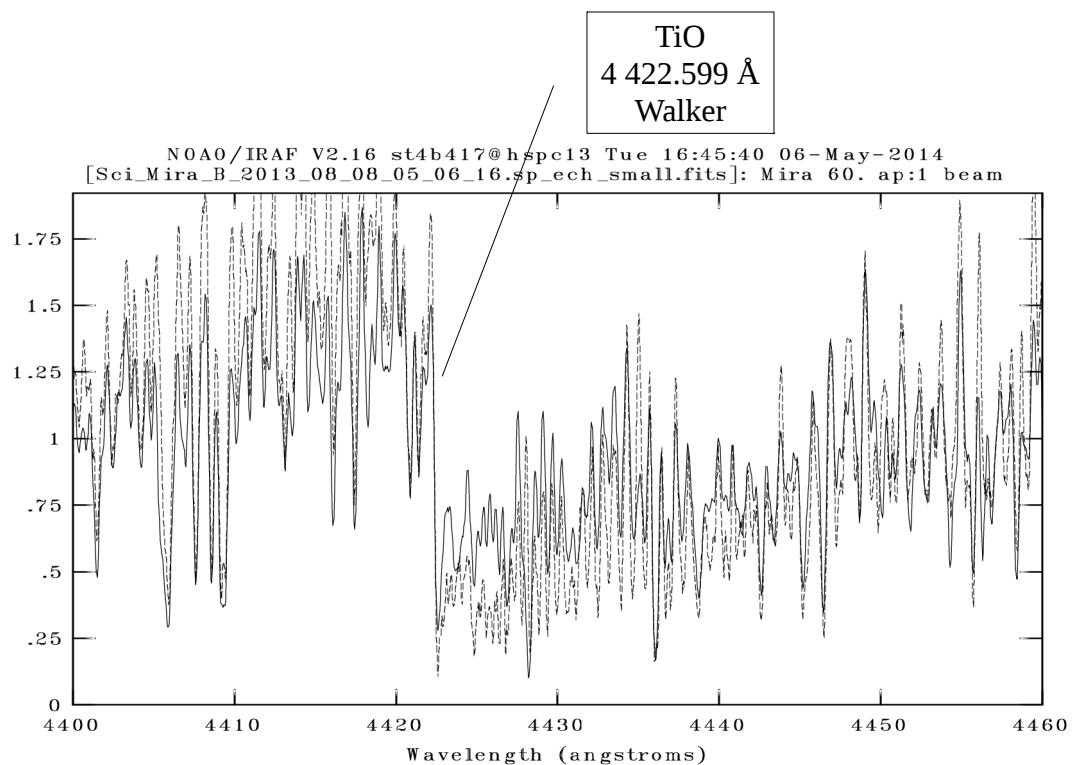
K I
4 044.136 Å
4 047.208 Å
Joy



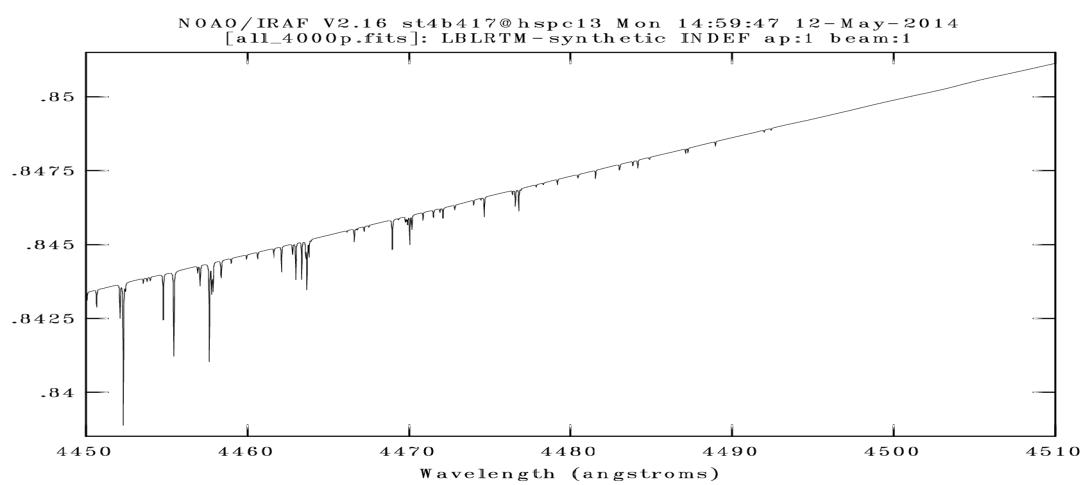
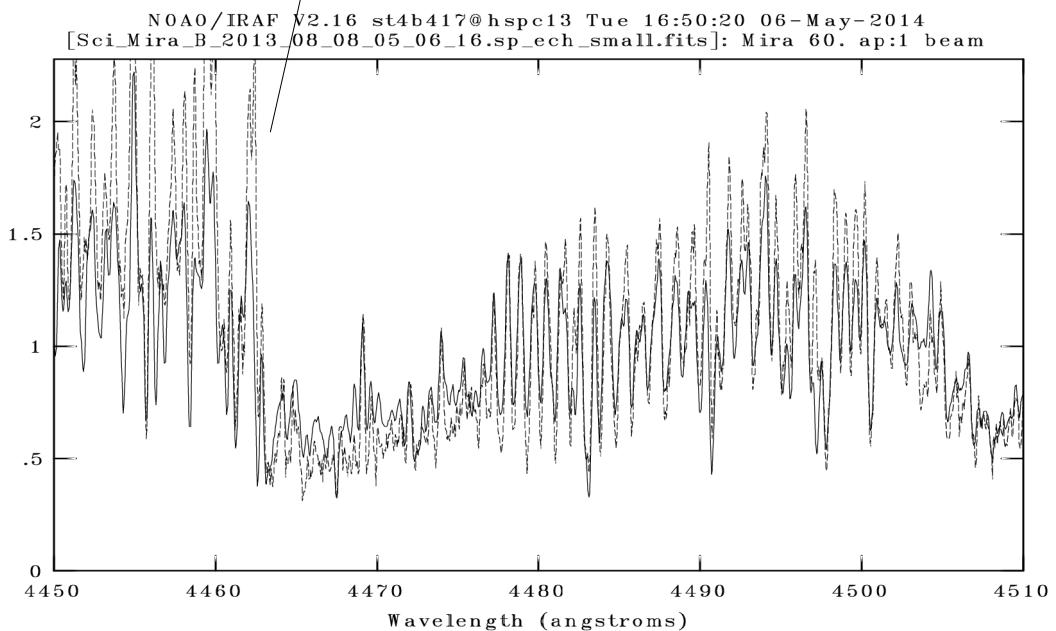


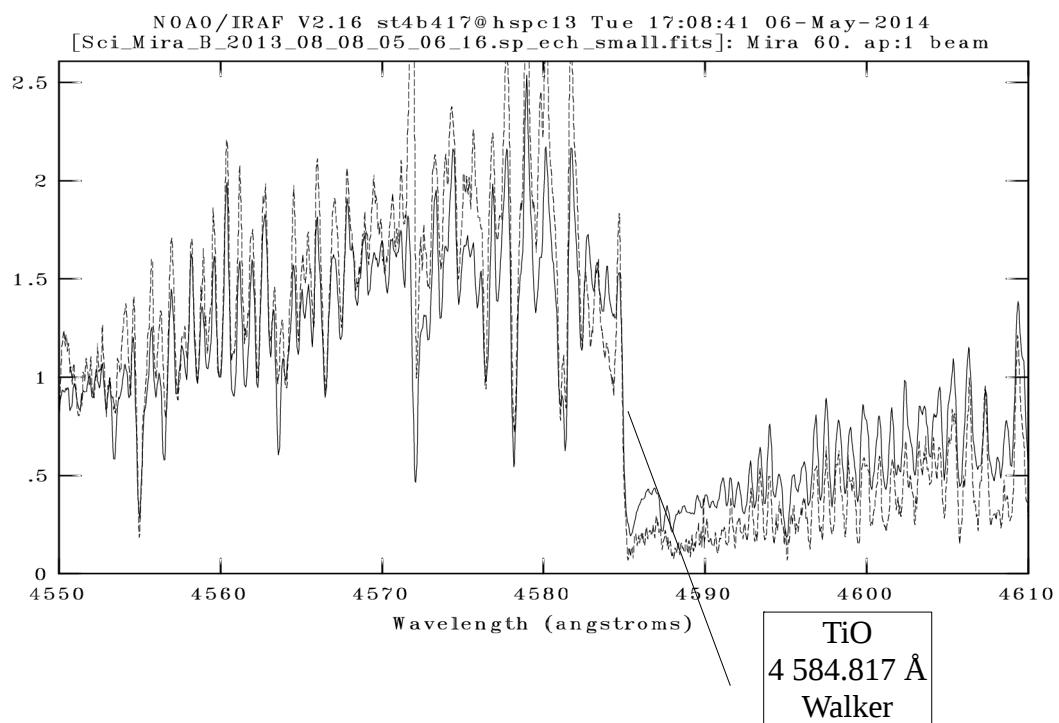
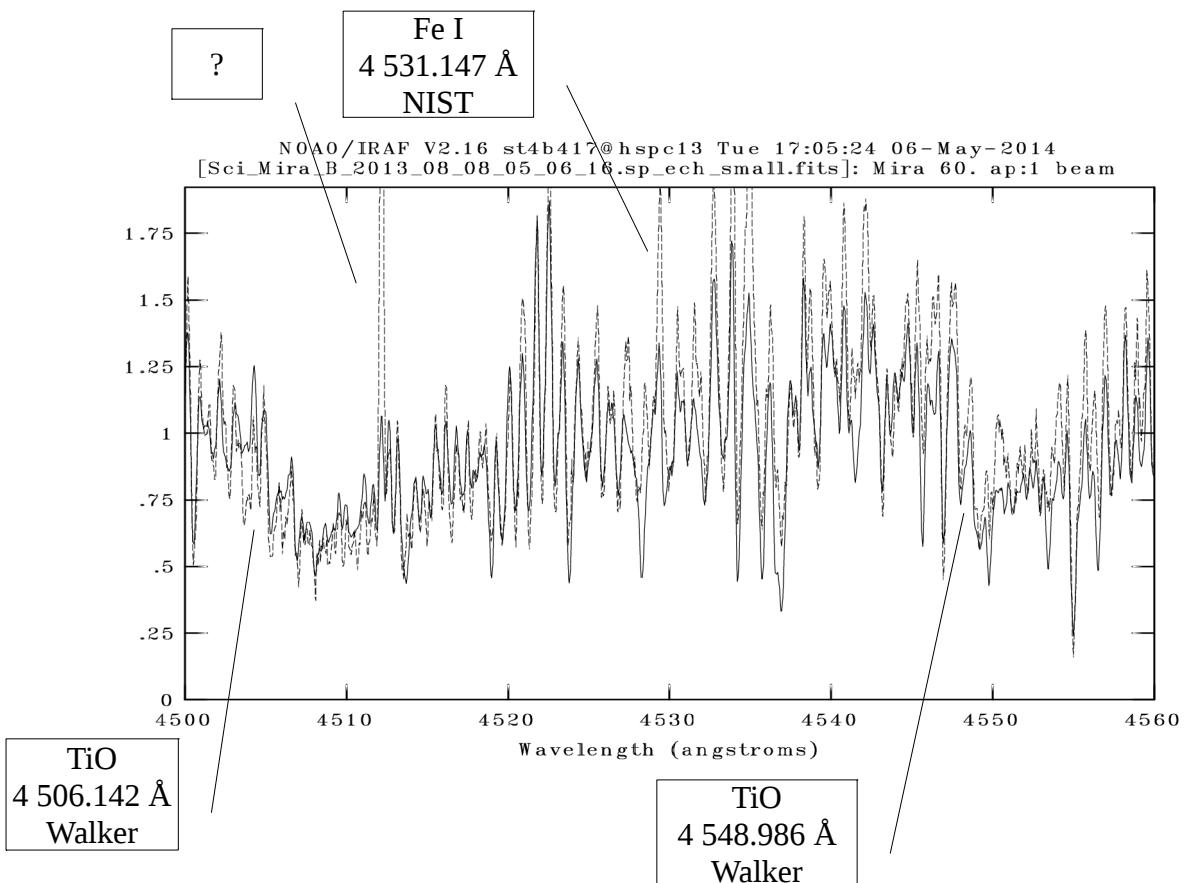


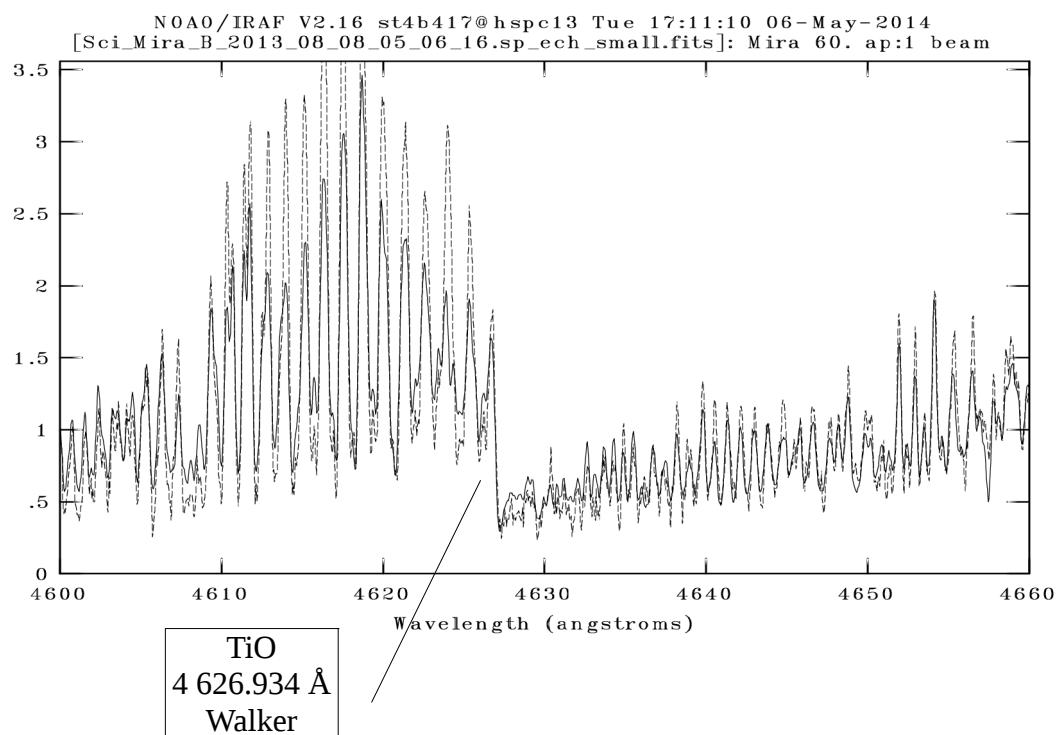


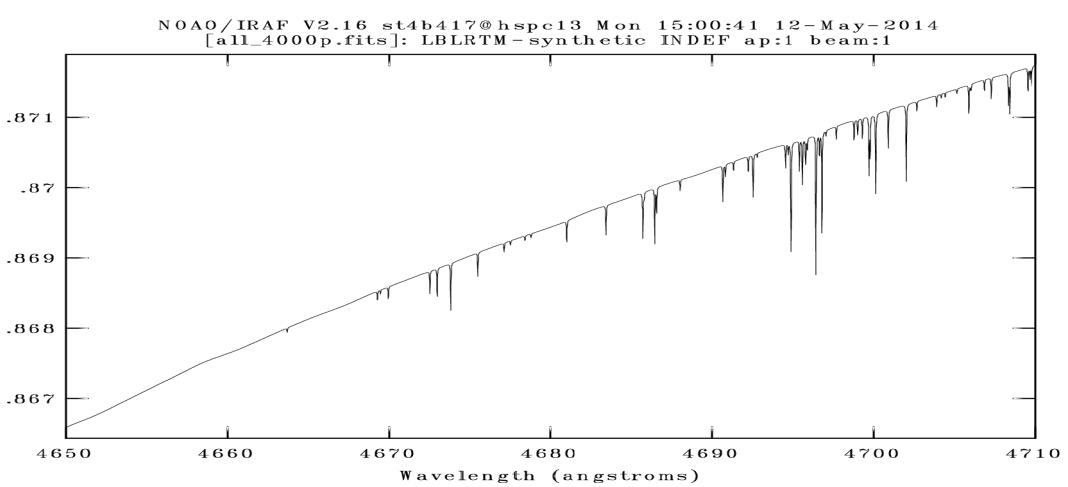
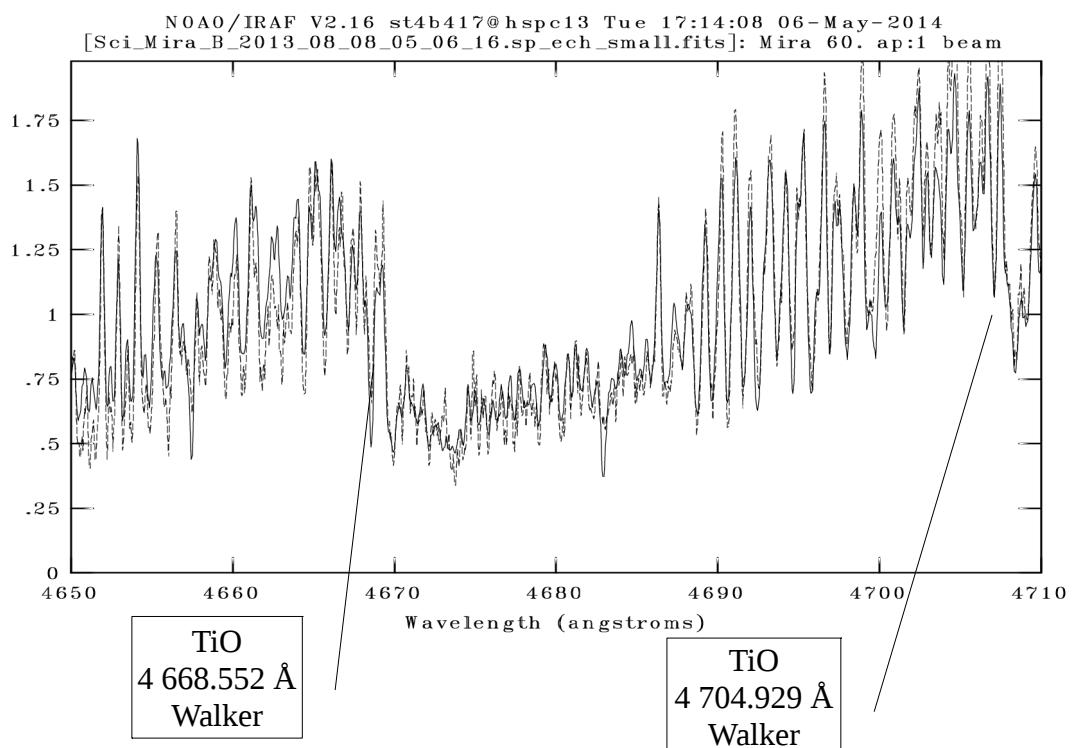


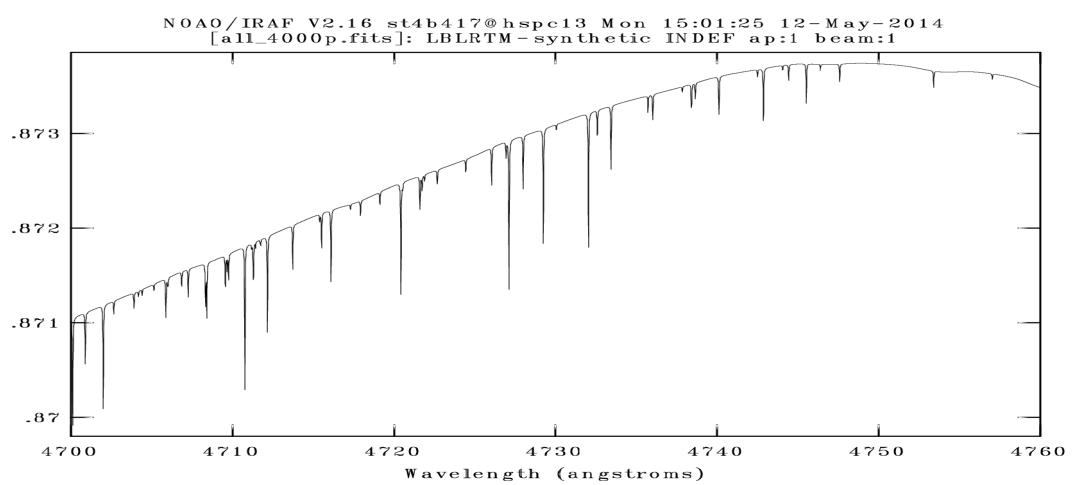
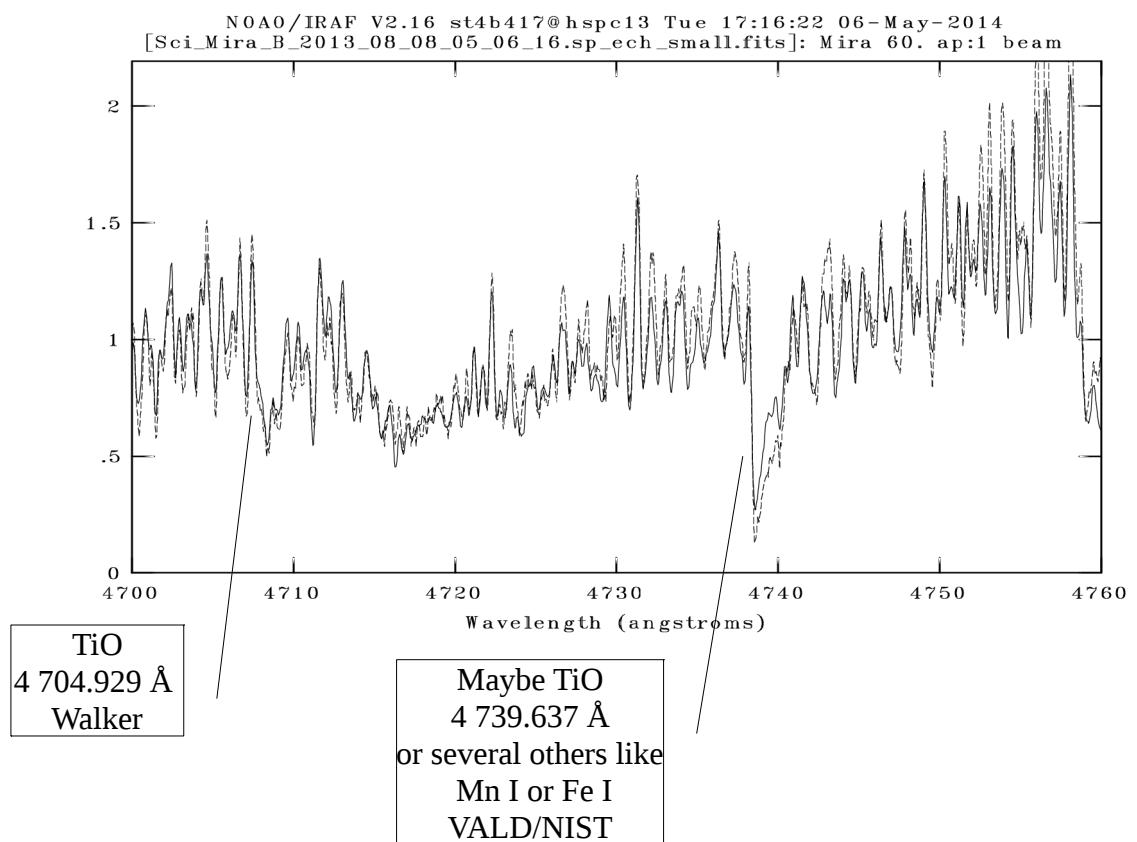
TiO
4 462.025 Å
Walker



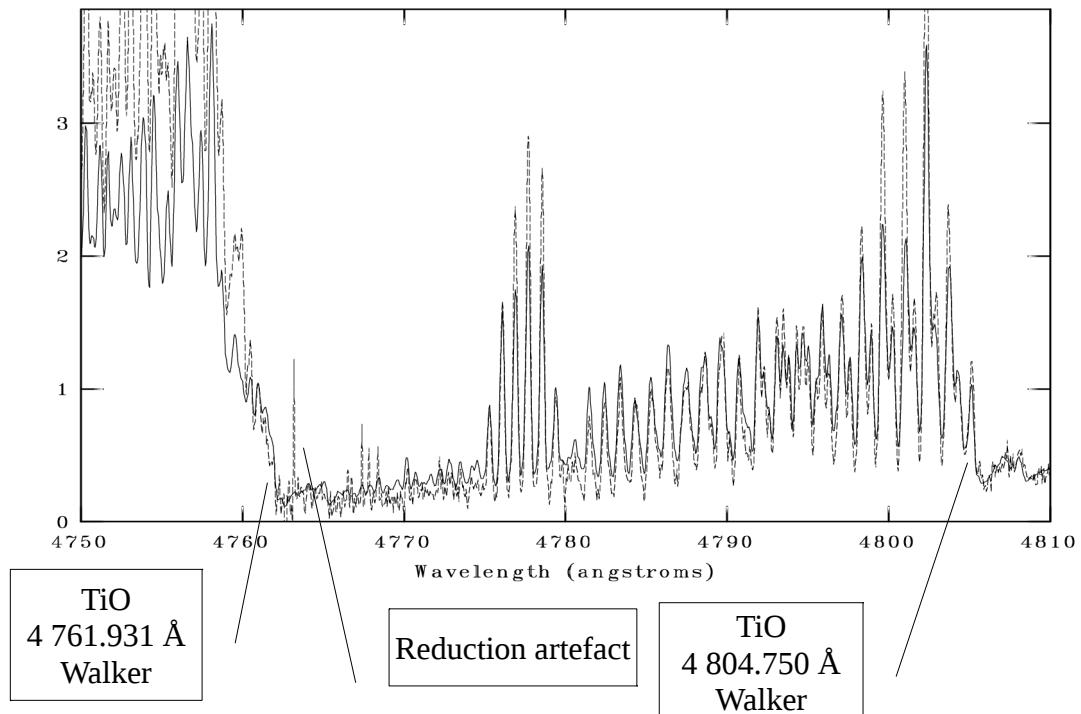






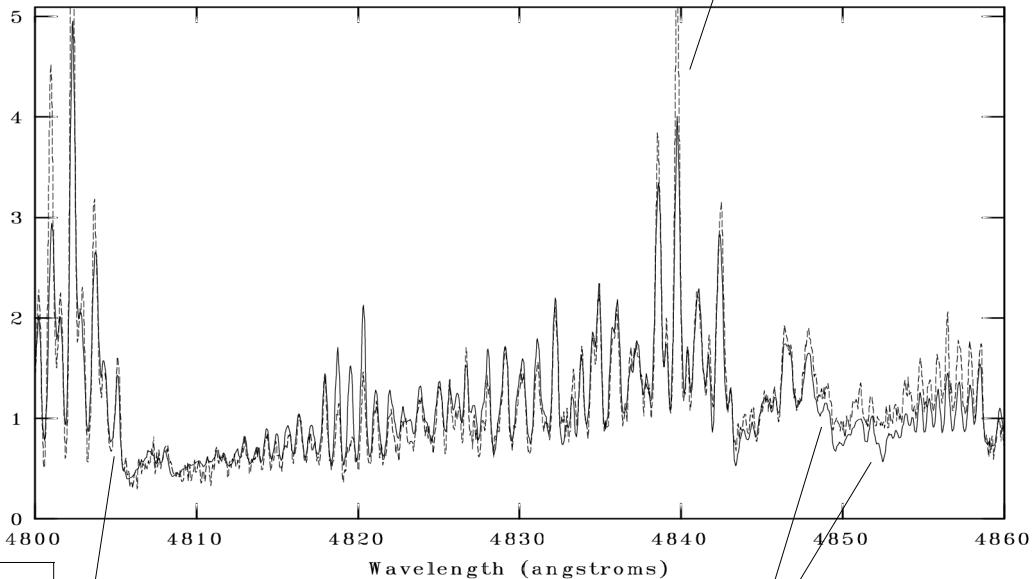


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Fe I
4 839.544 Å
NIST

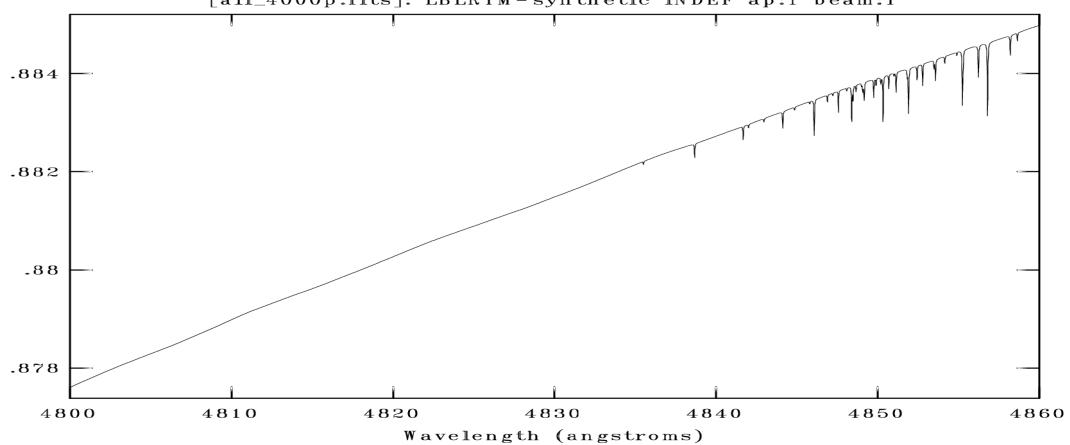
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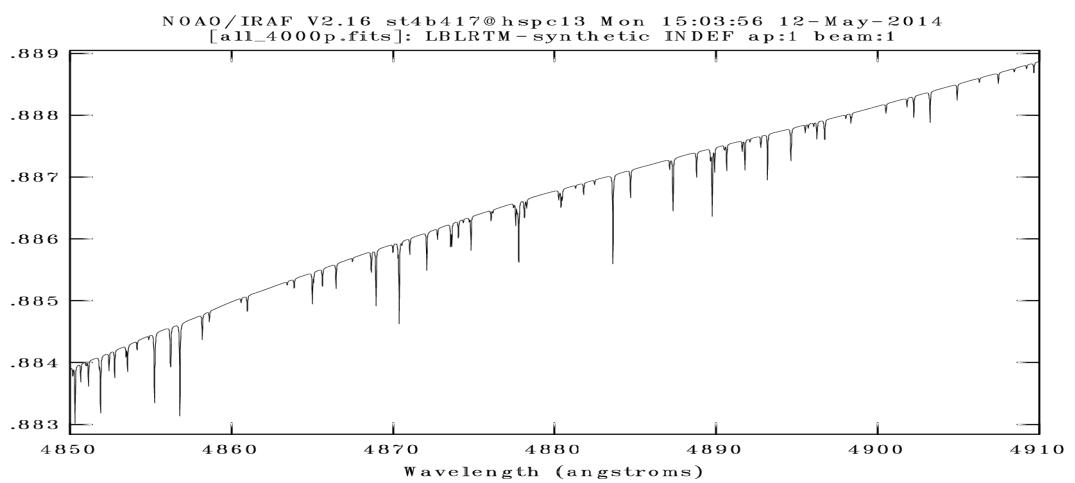
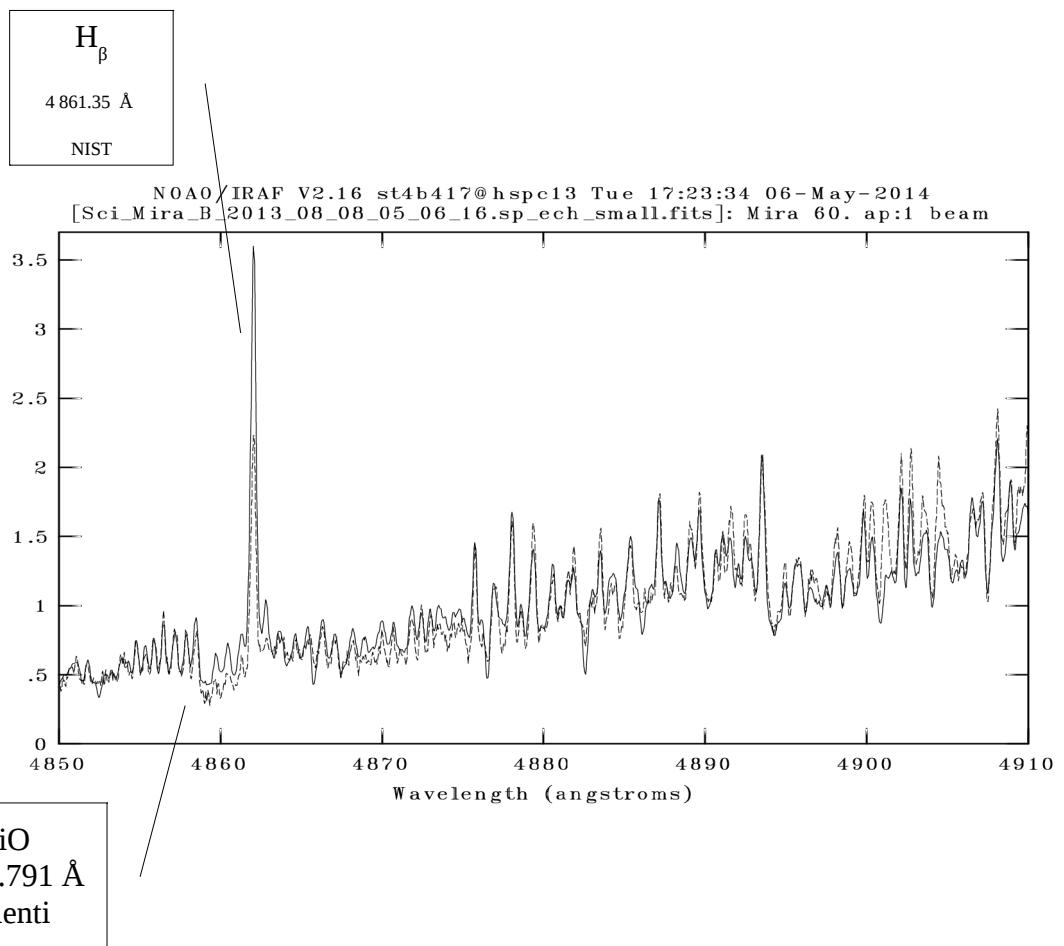


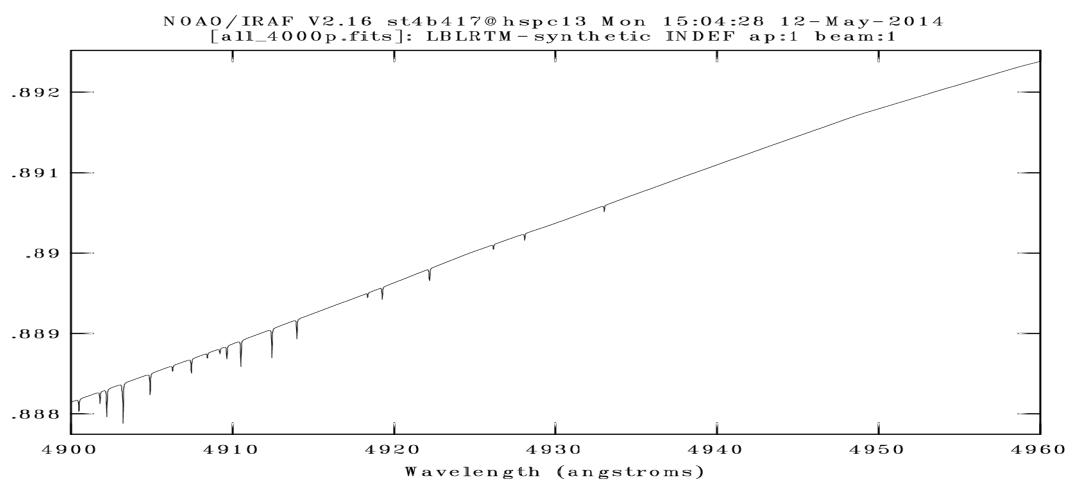
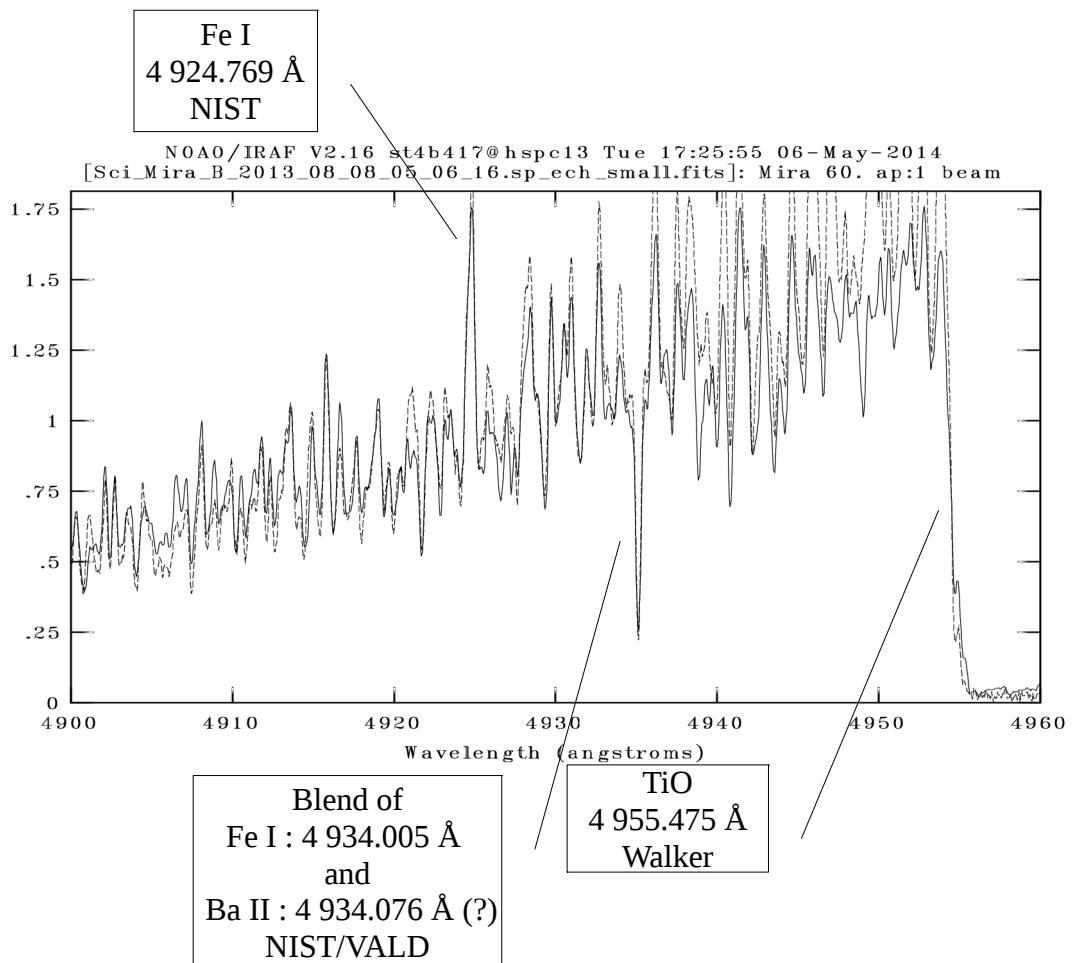
TiO
4 804.750 Å
Walker

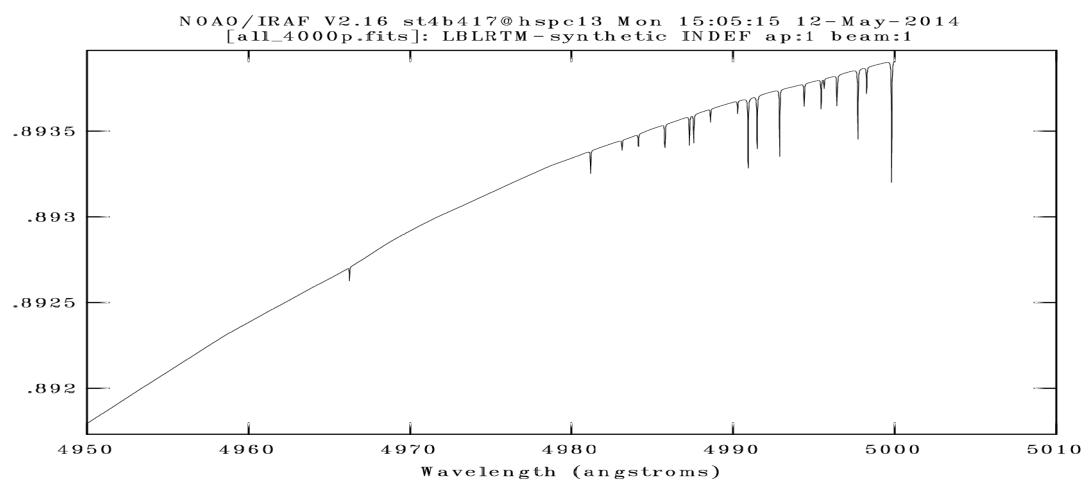
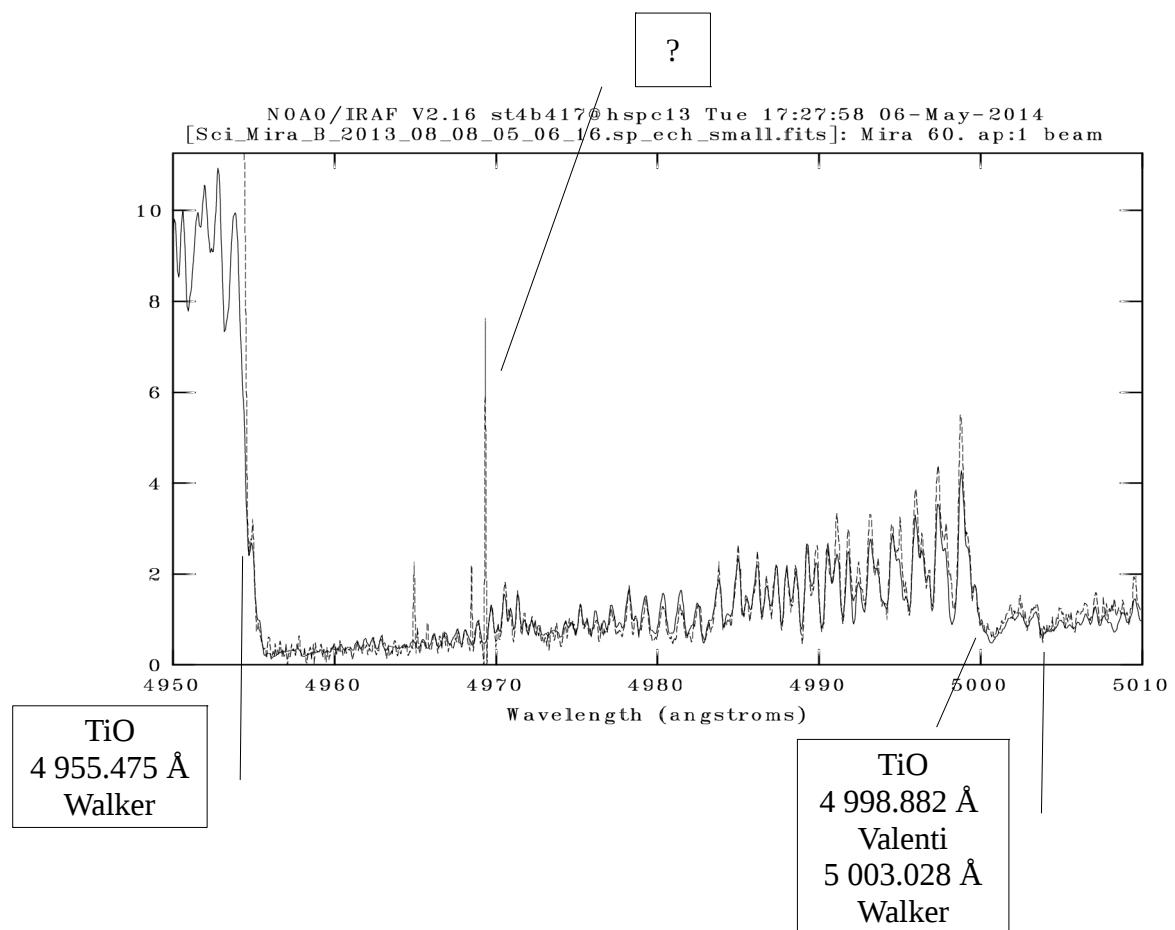
TiO
4 847.574 Å
Walker
4 851.949 Å
Valenti

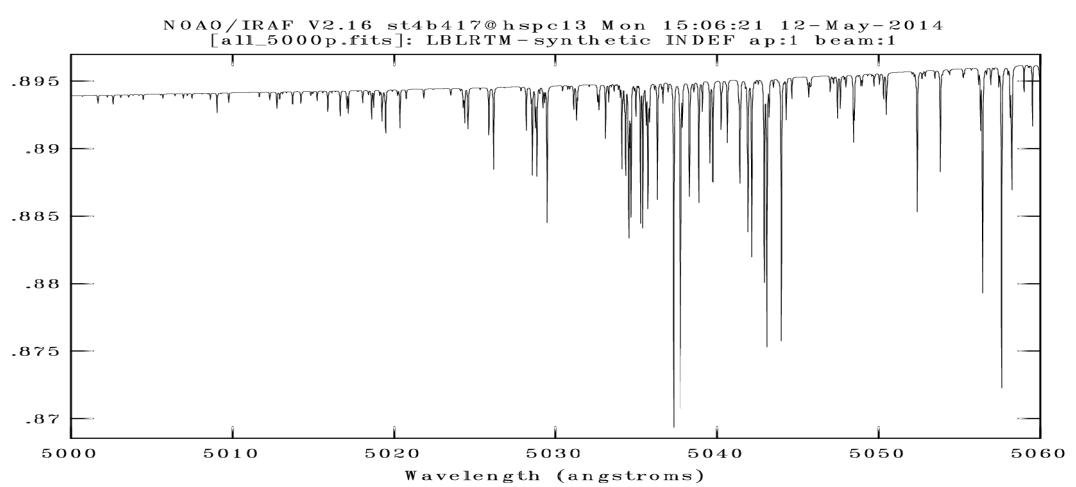
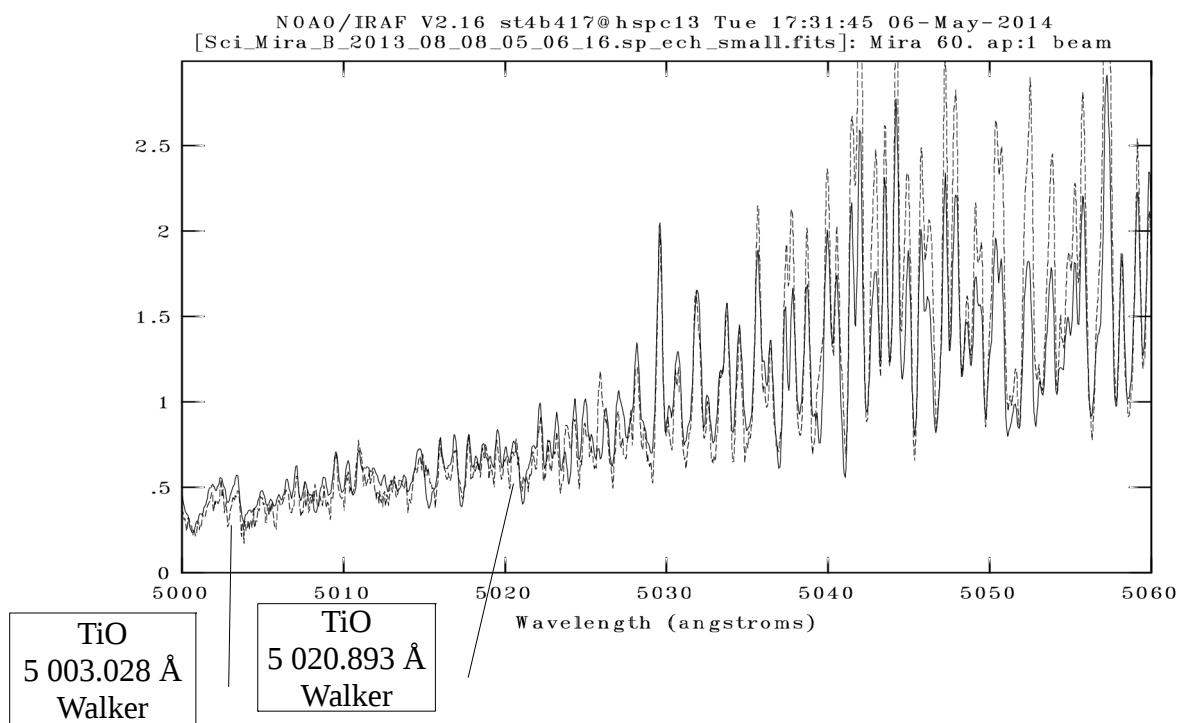
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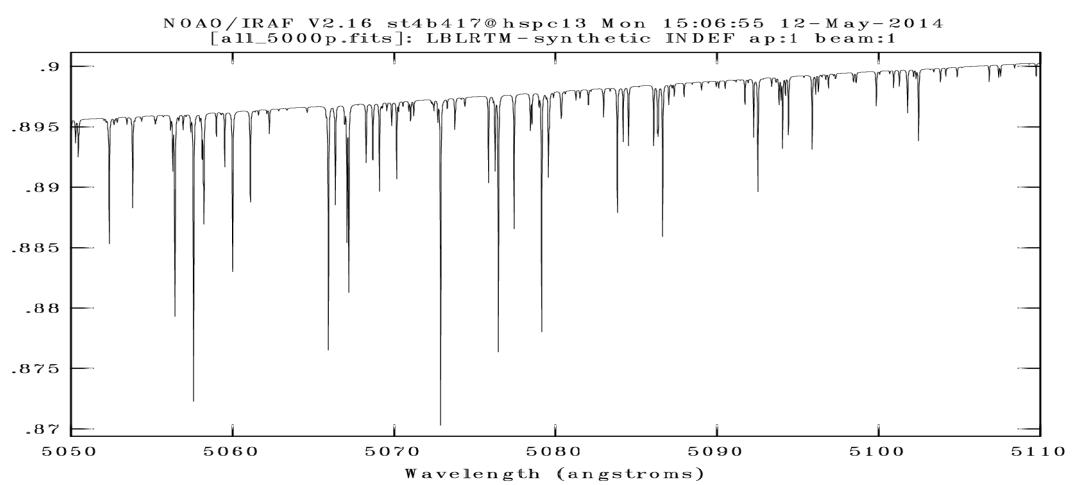
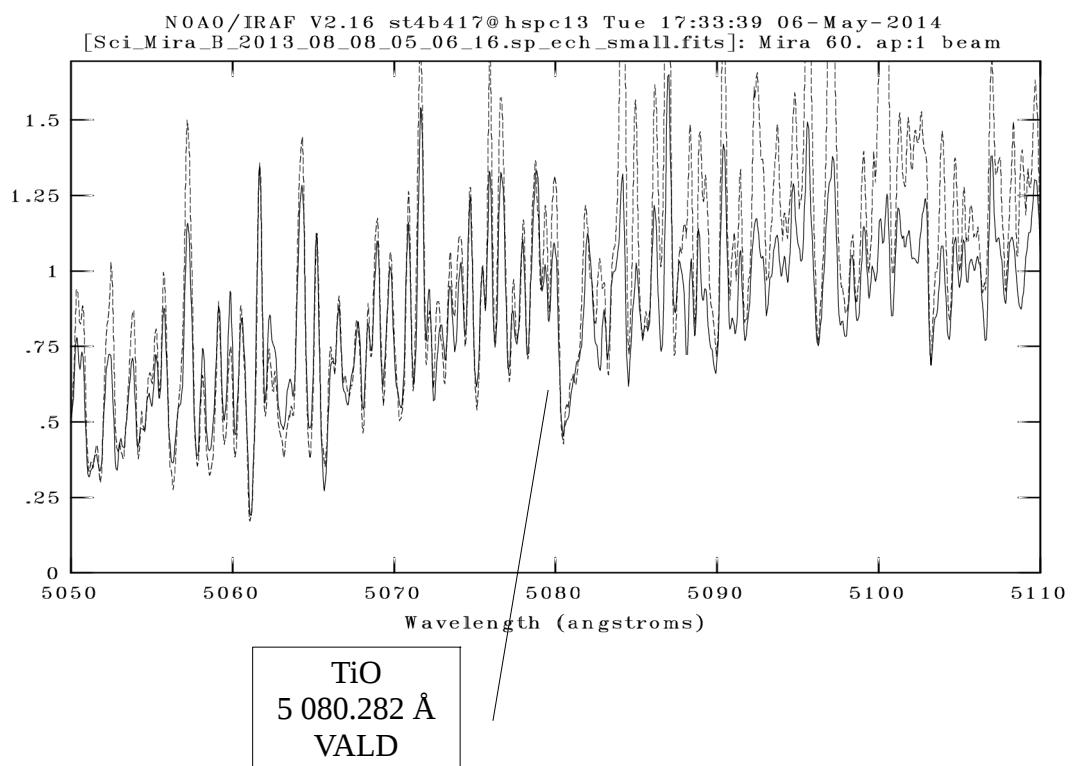


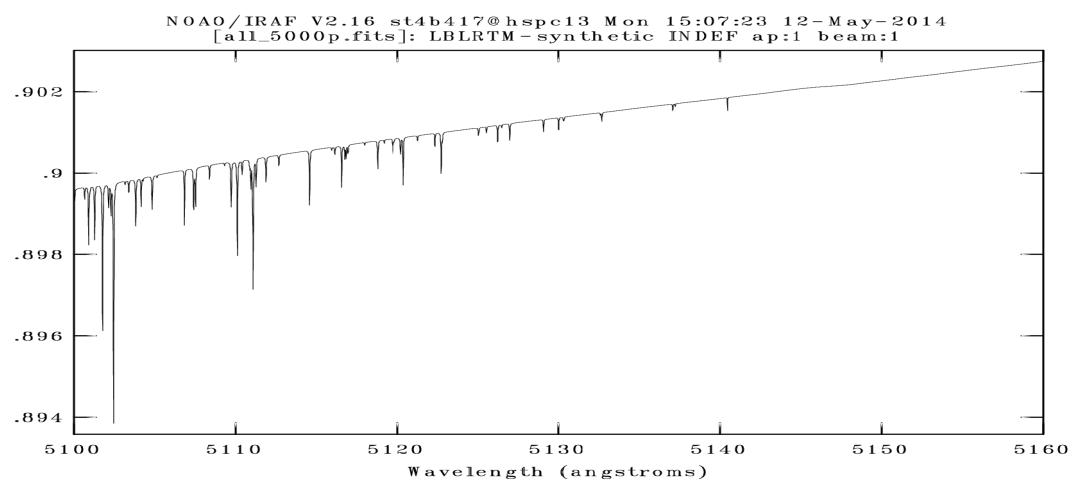
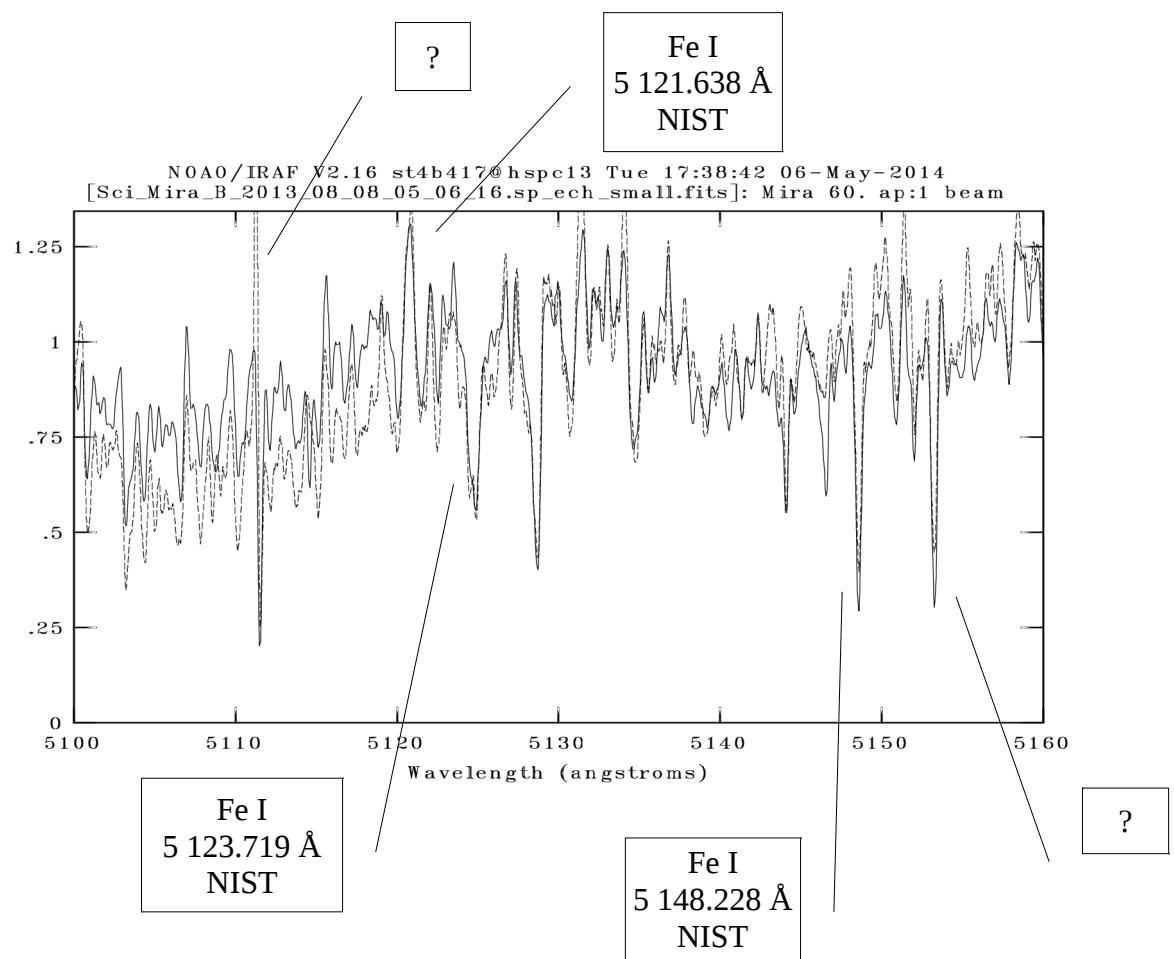


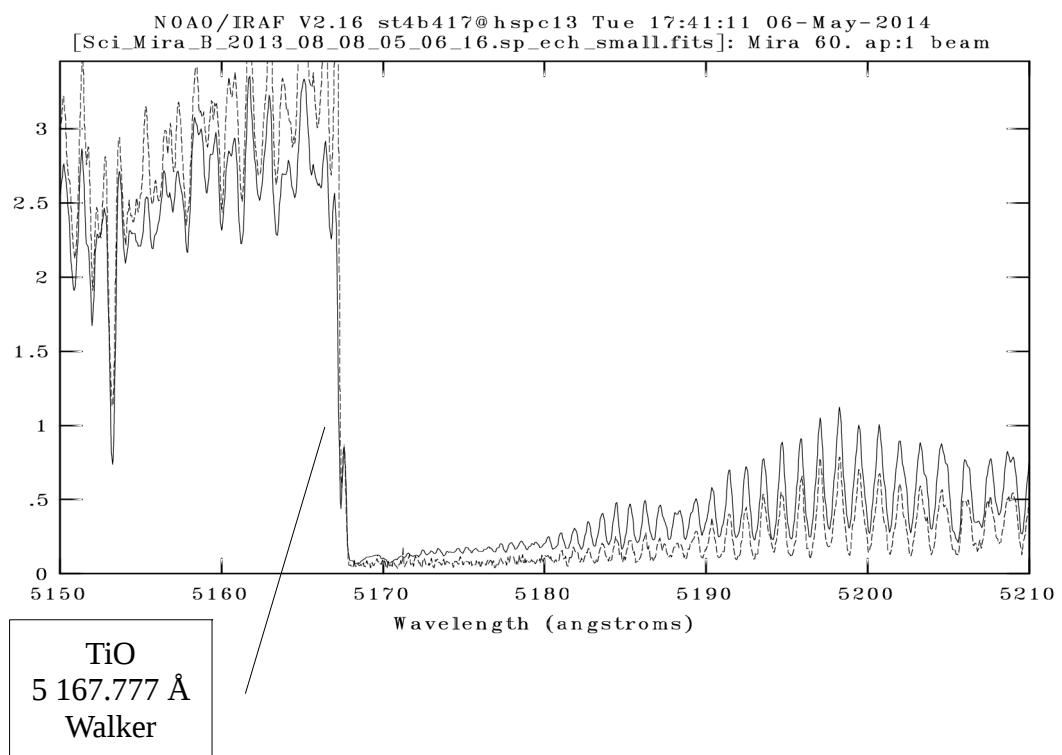


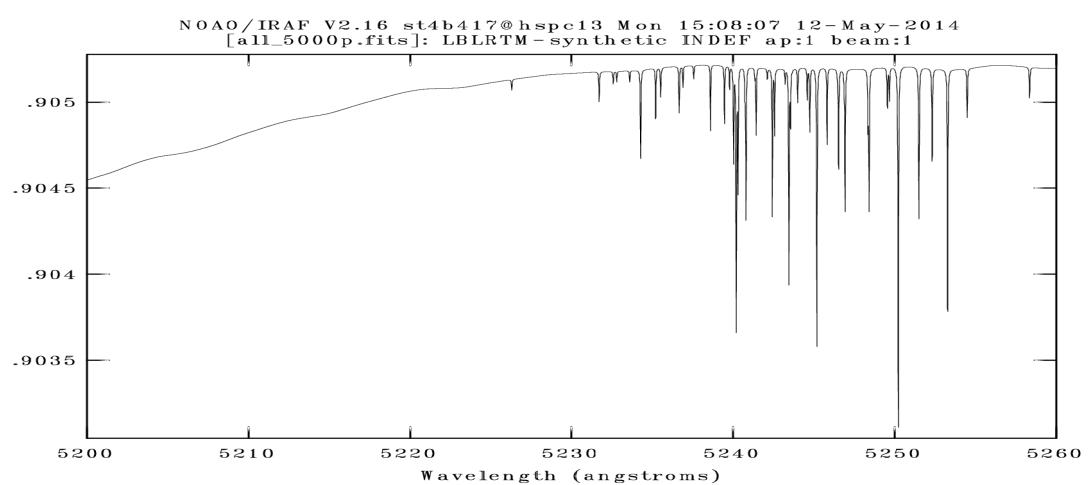
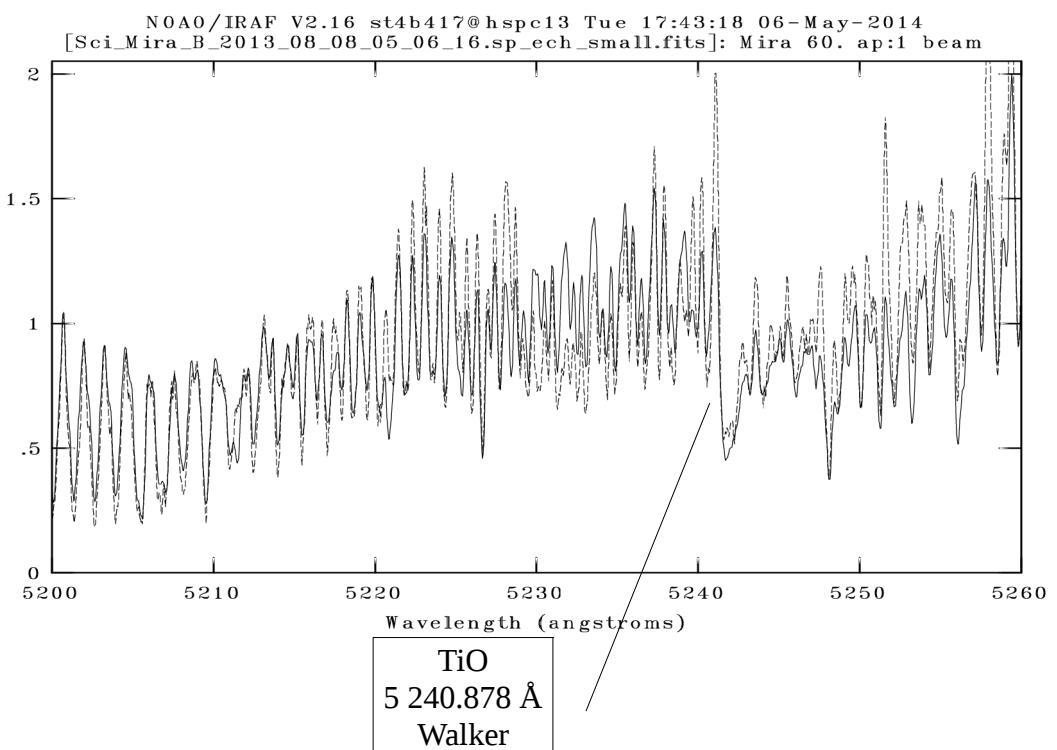


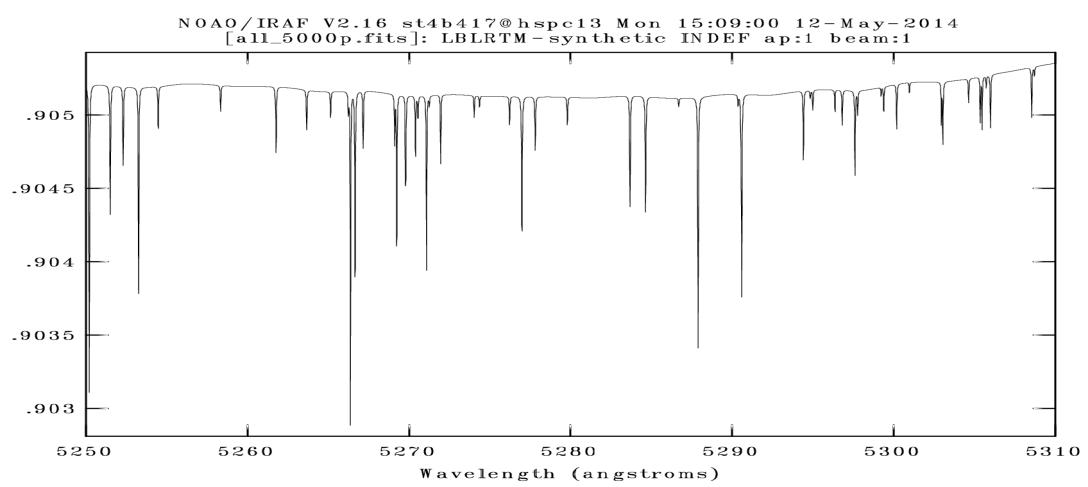
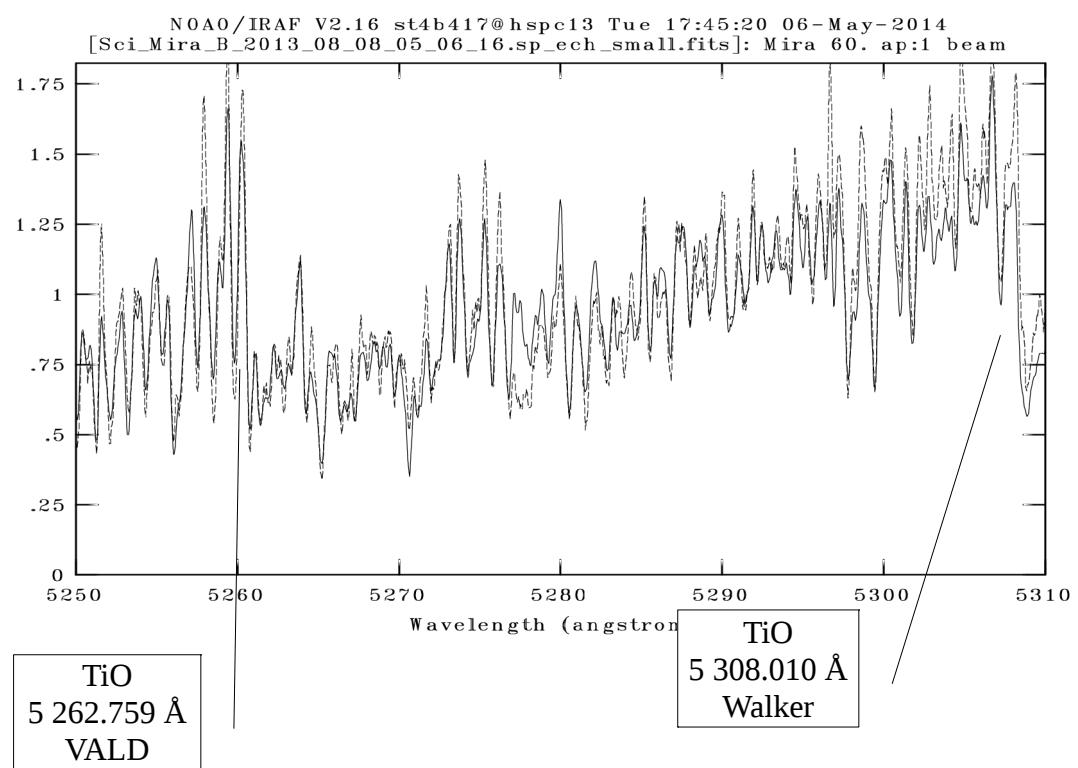


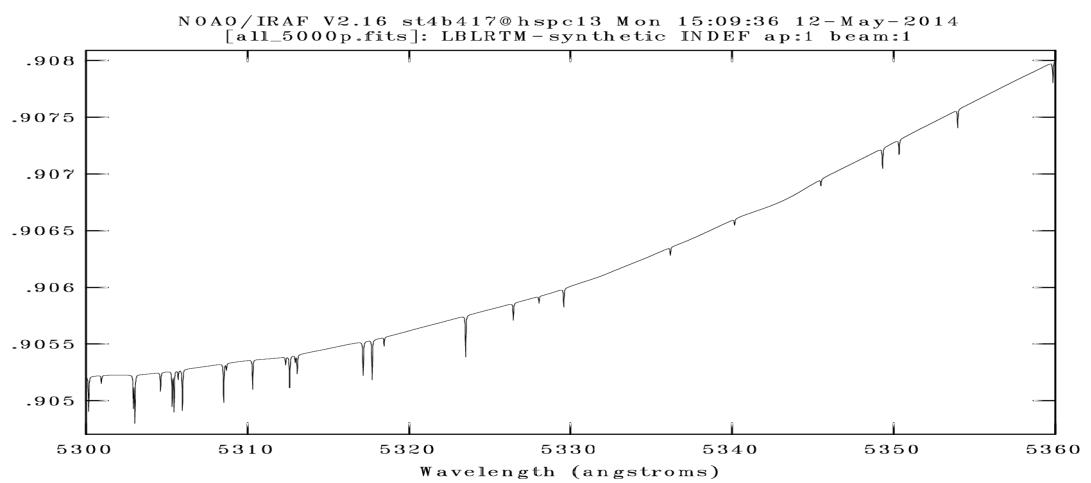
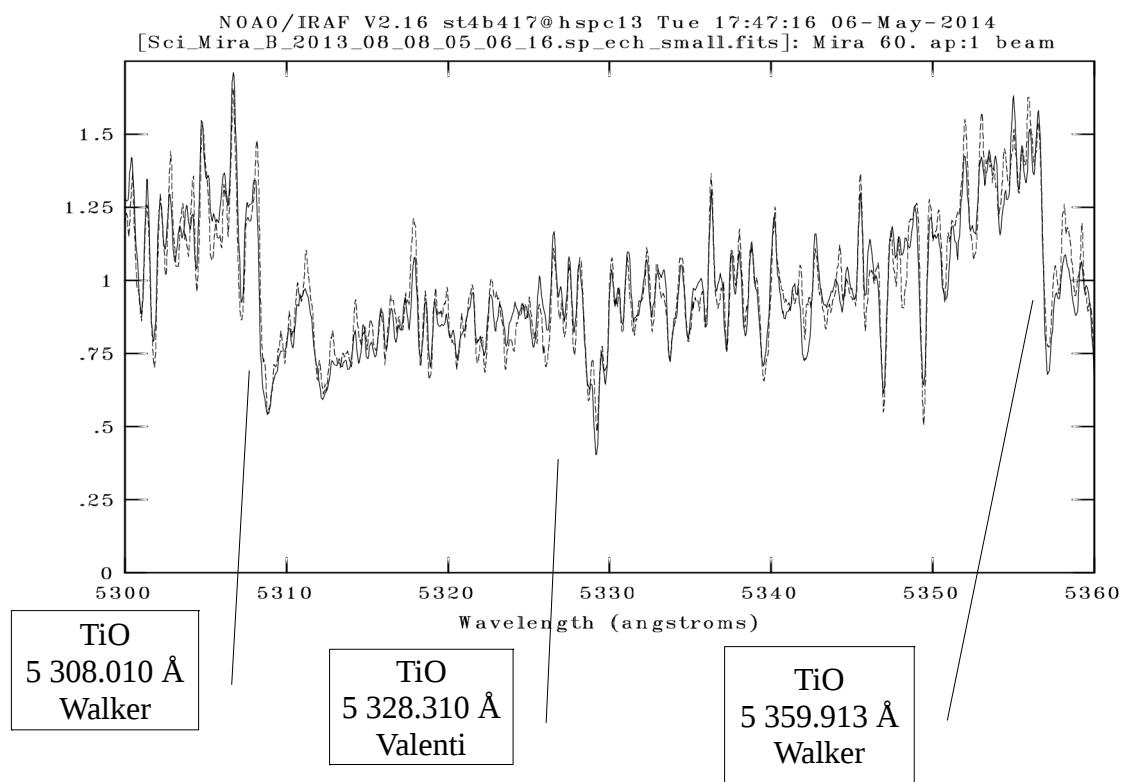


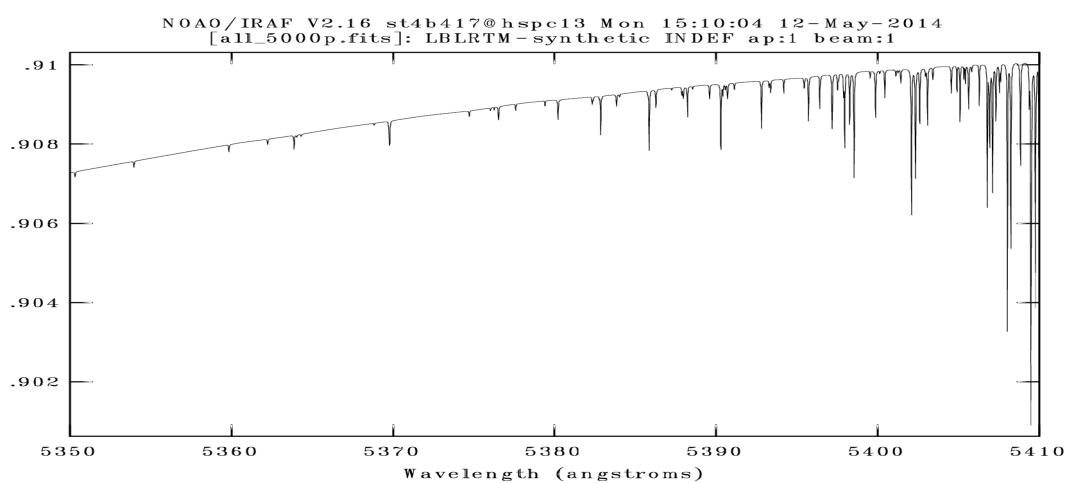
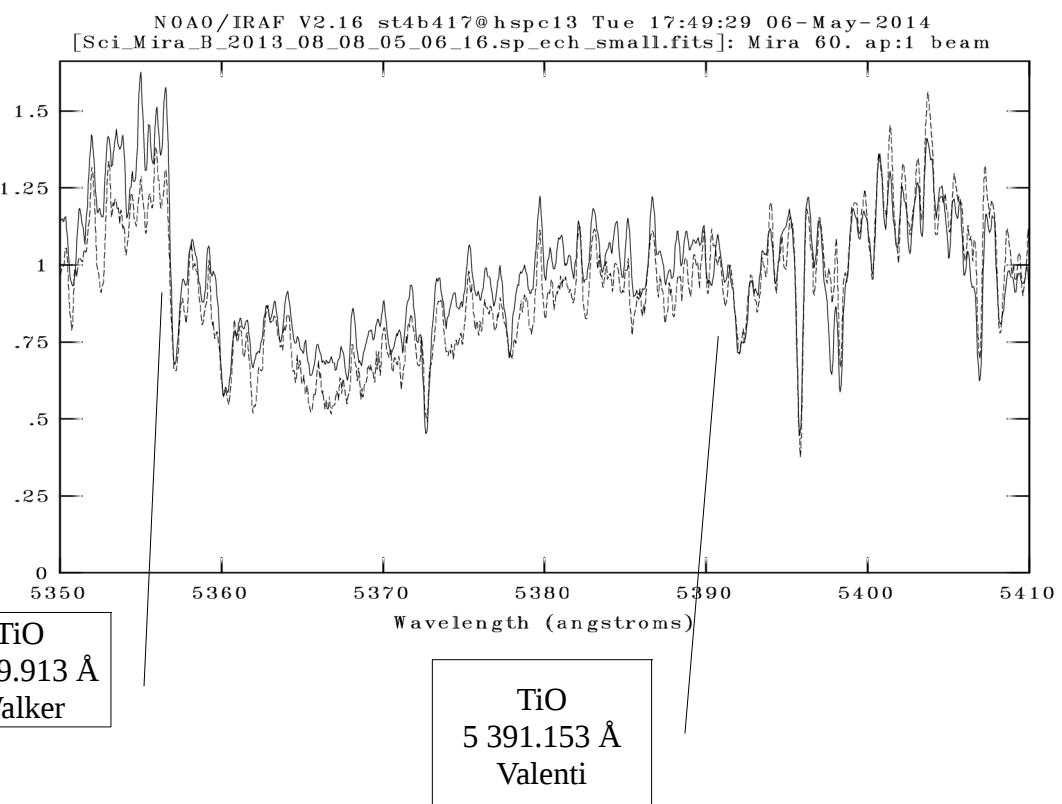


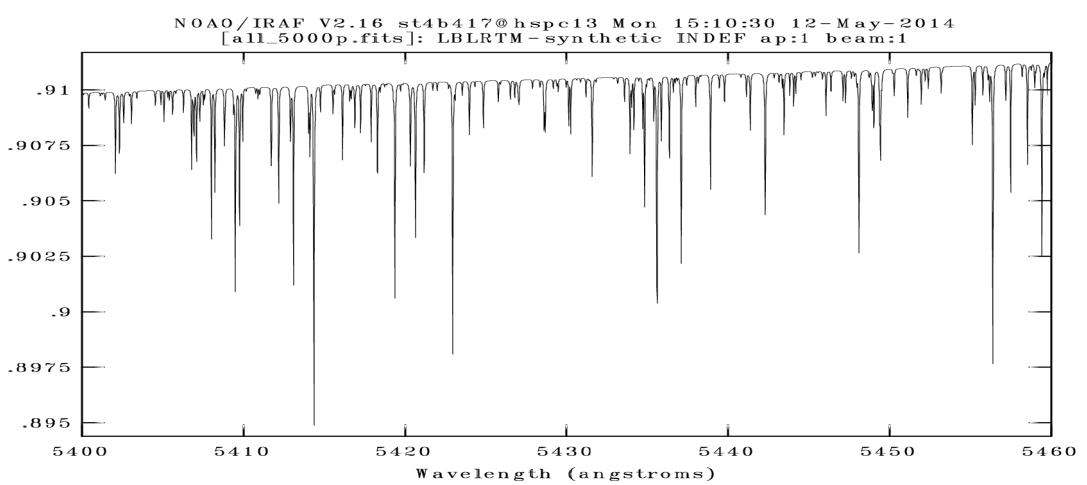
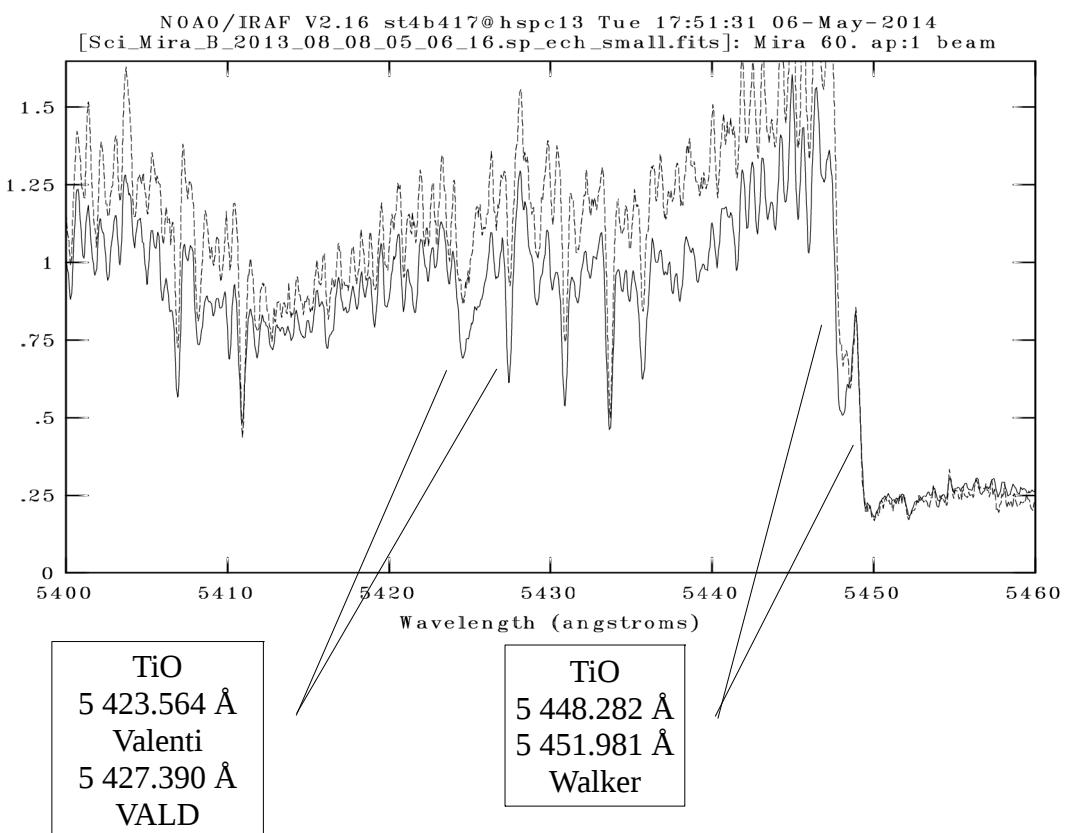


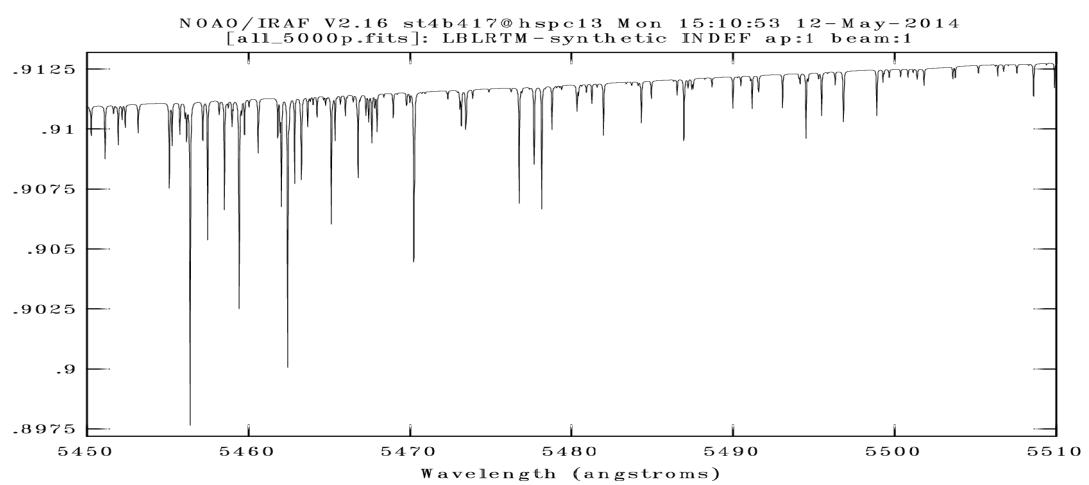
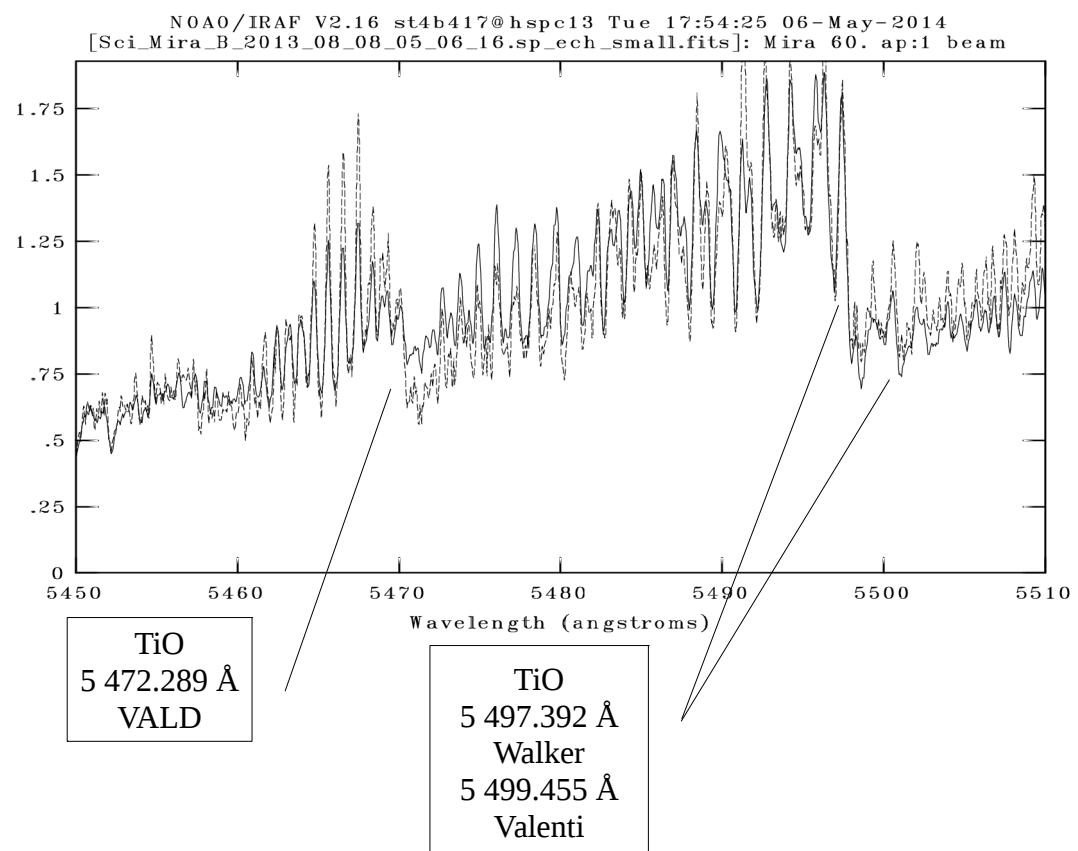


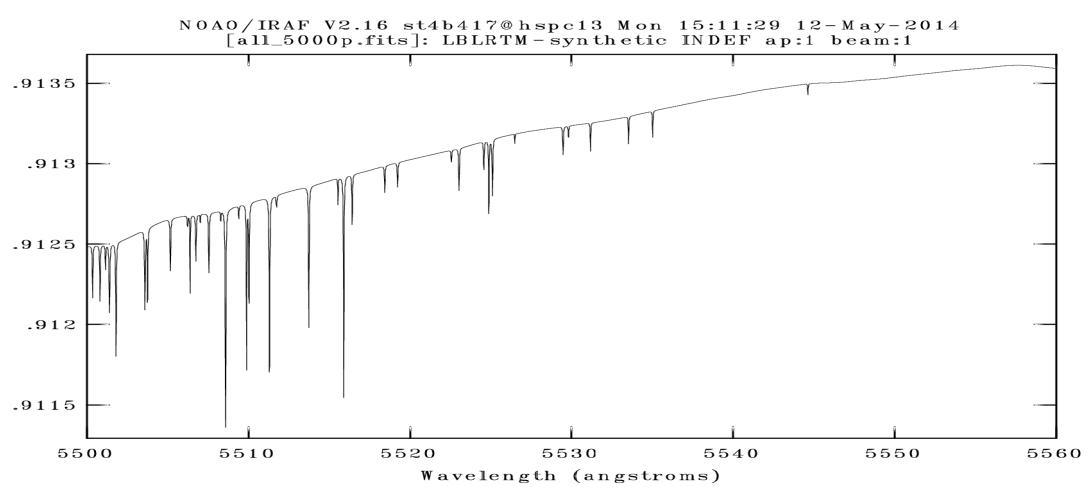
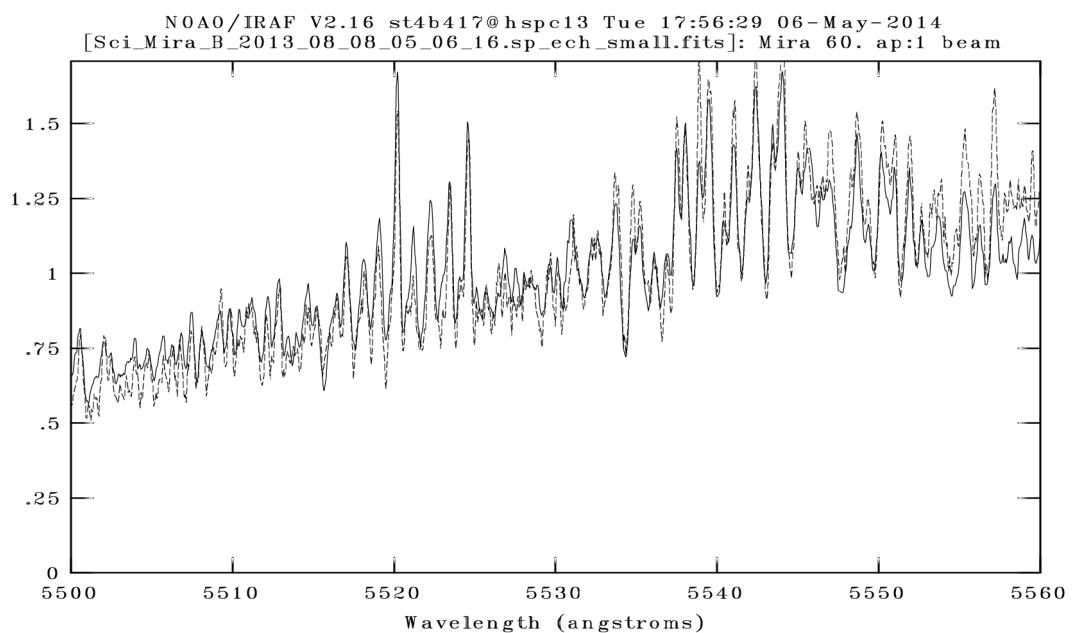


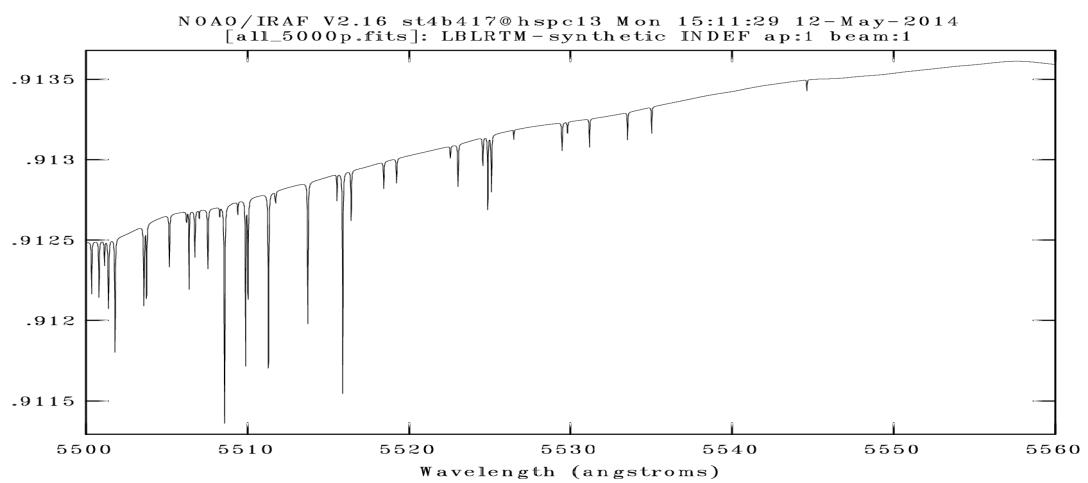
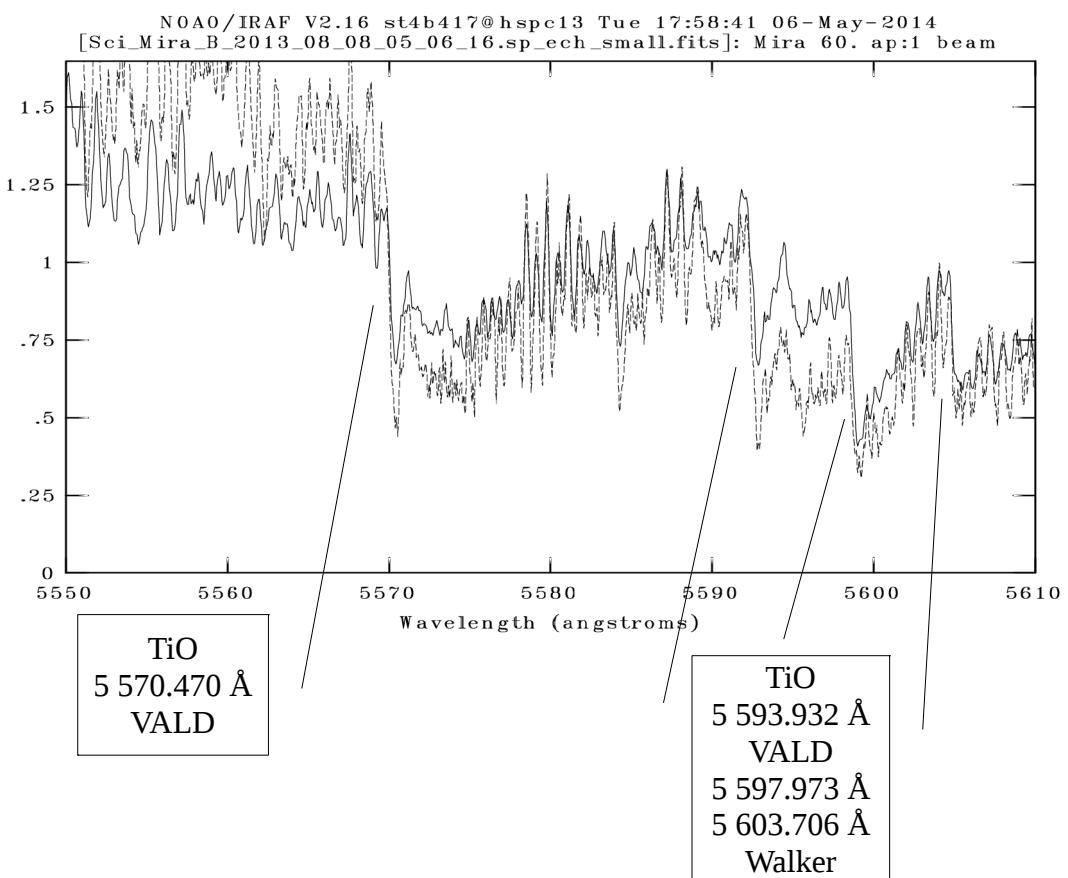


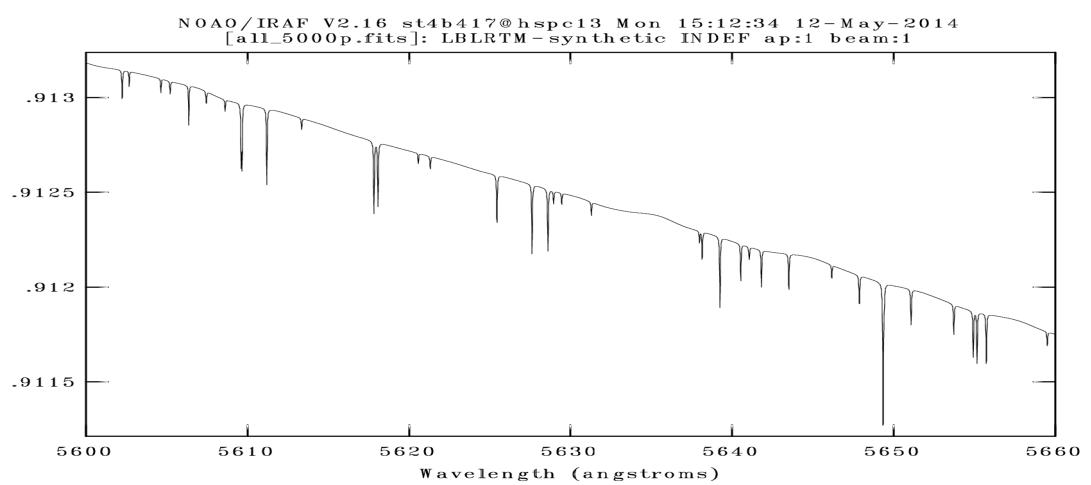
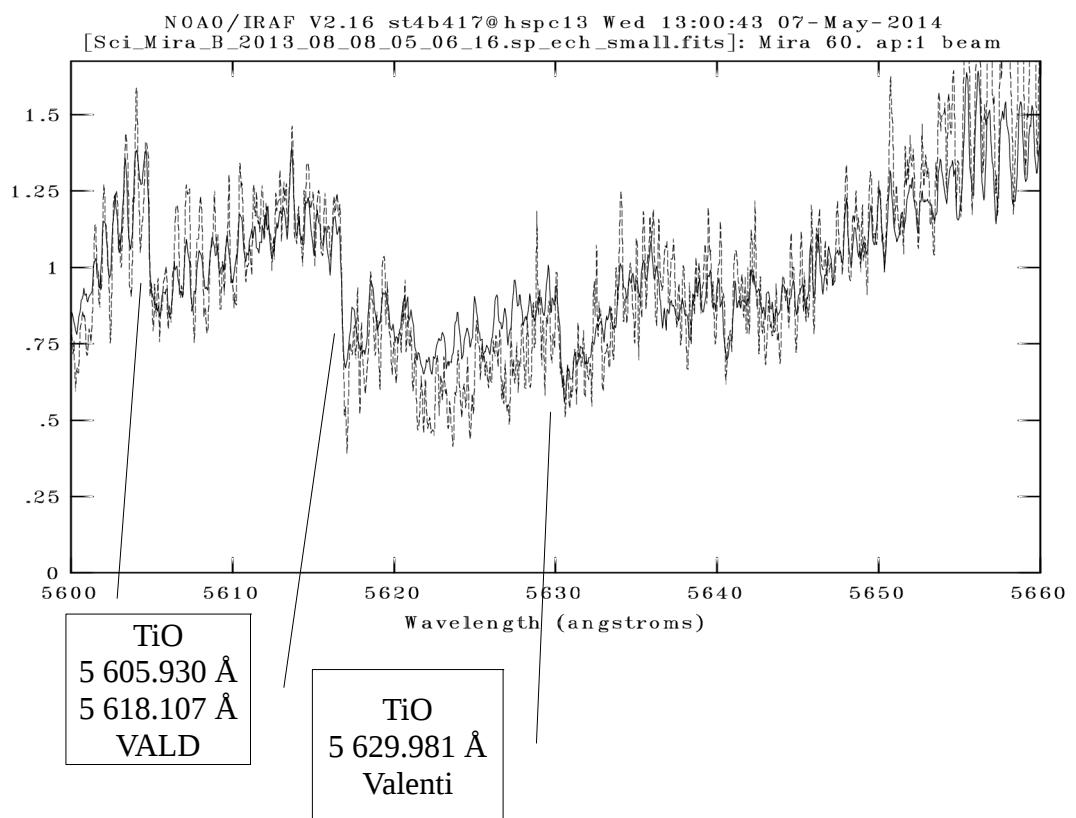


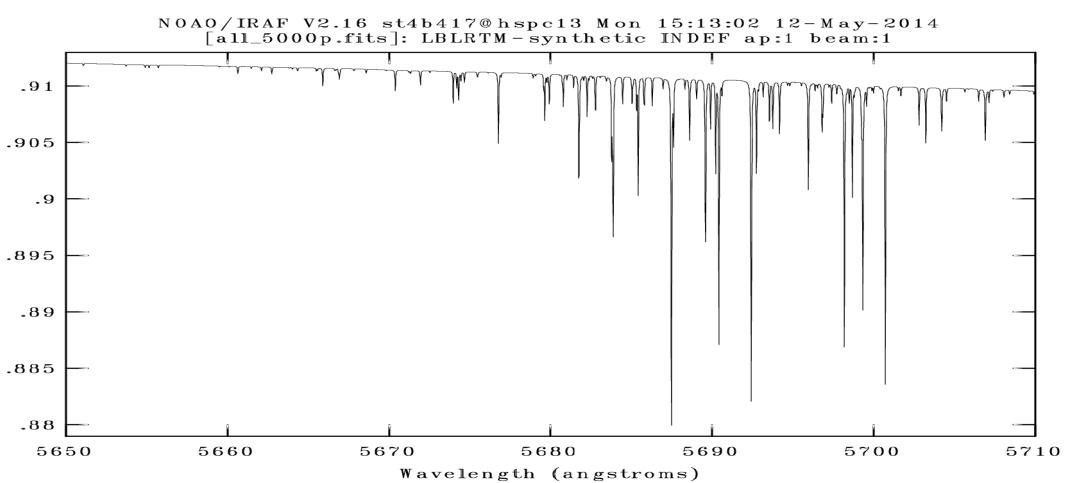
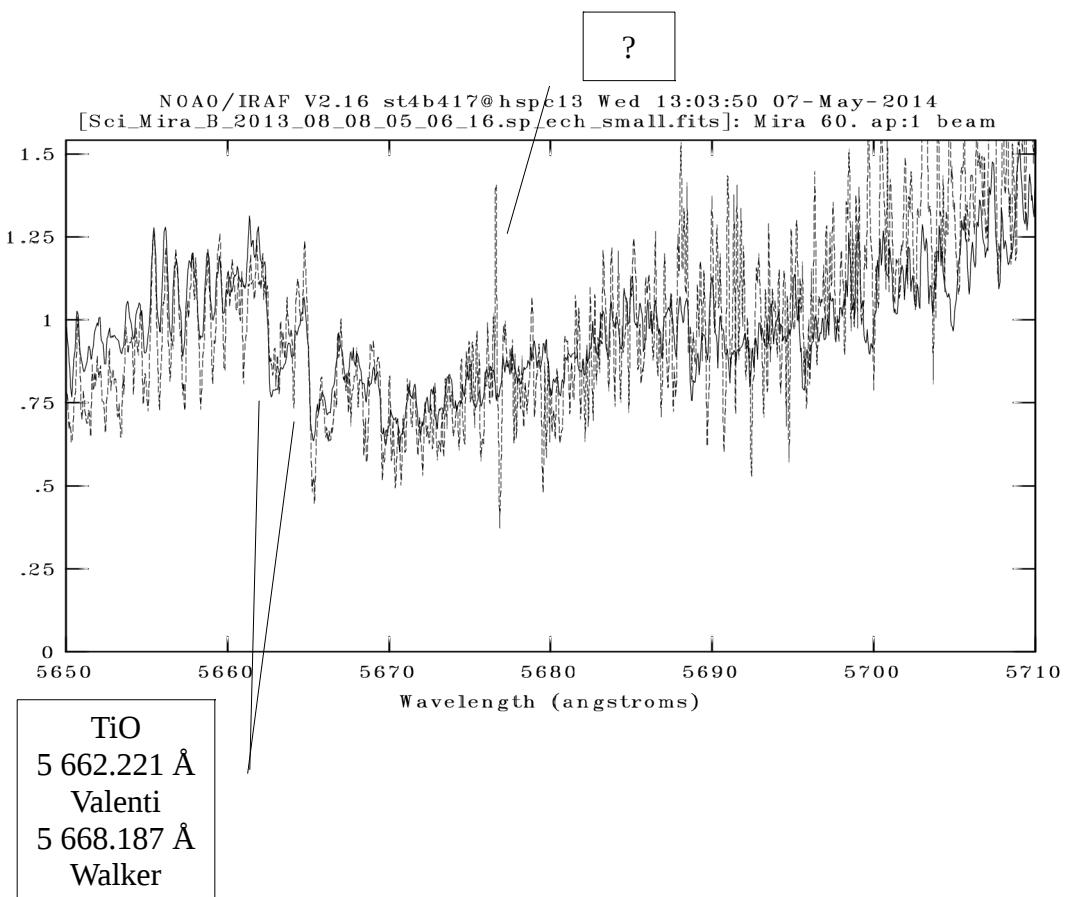




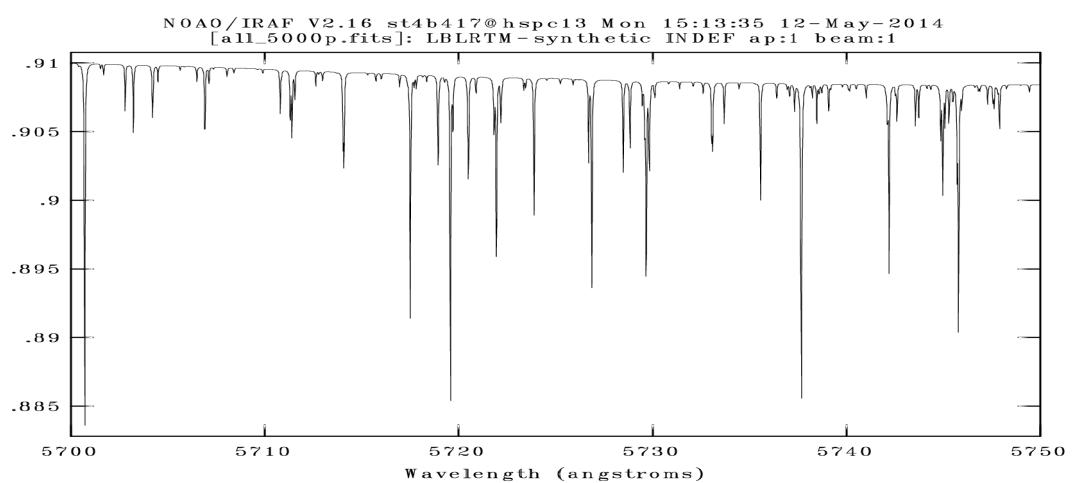
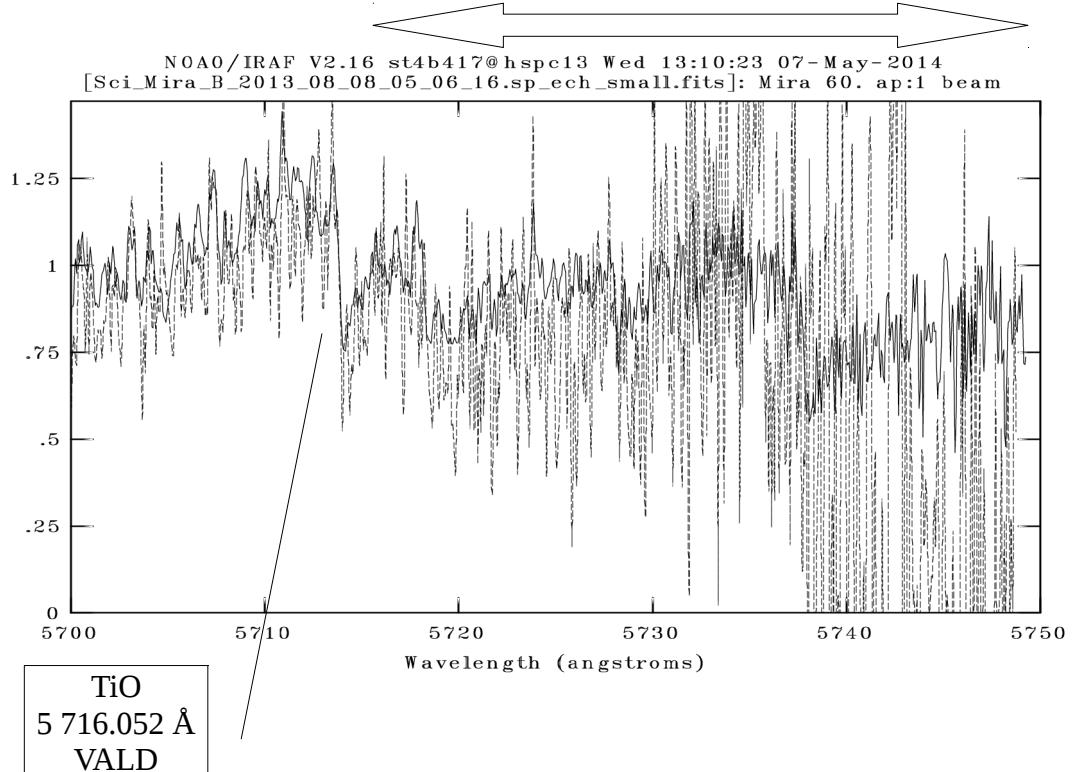




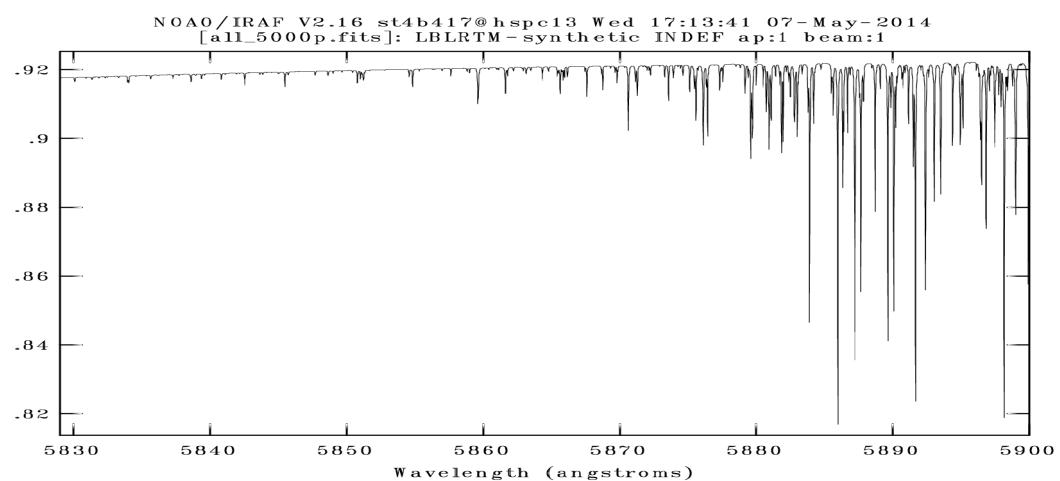
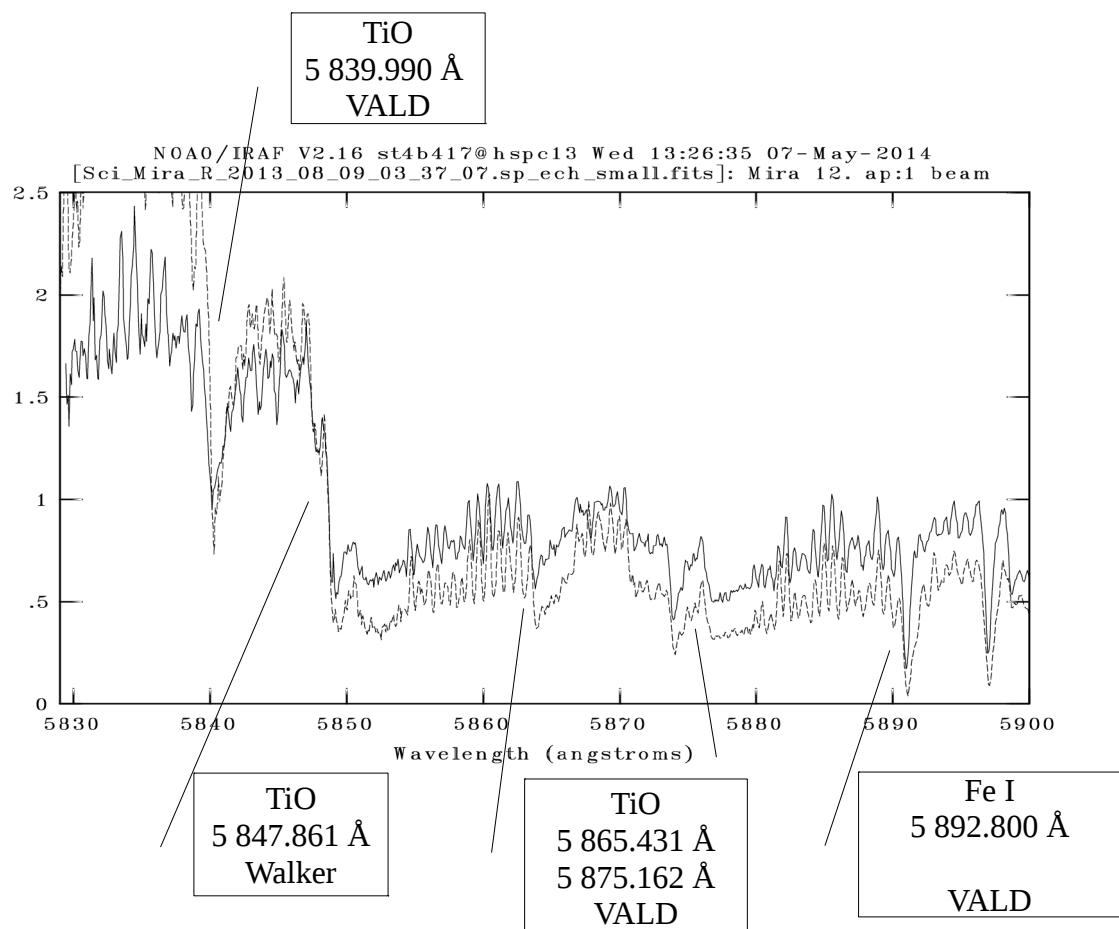


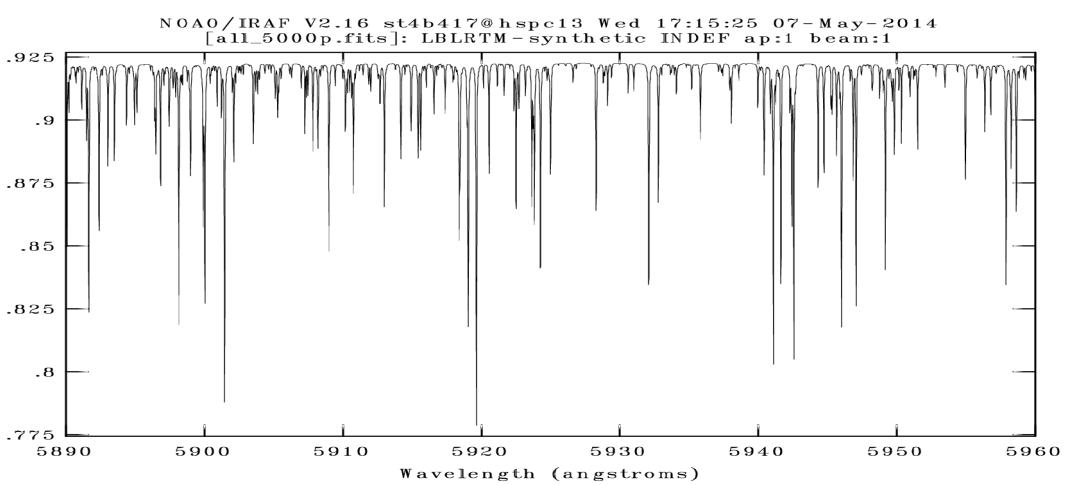
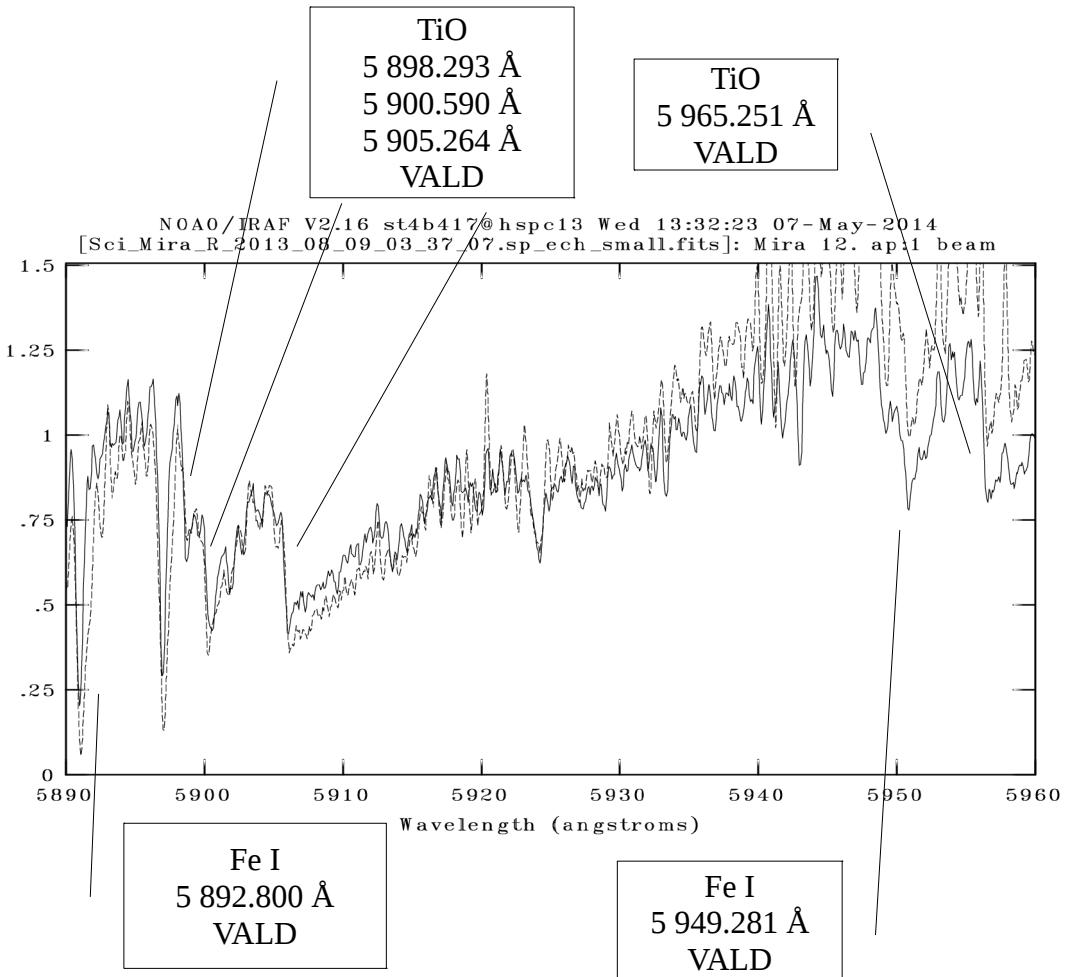


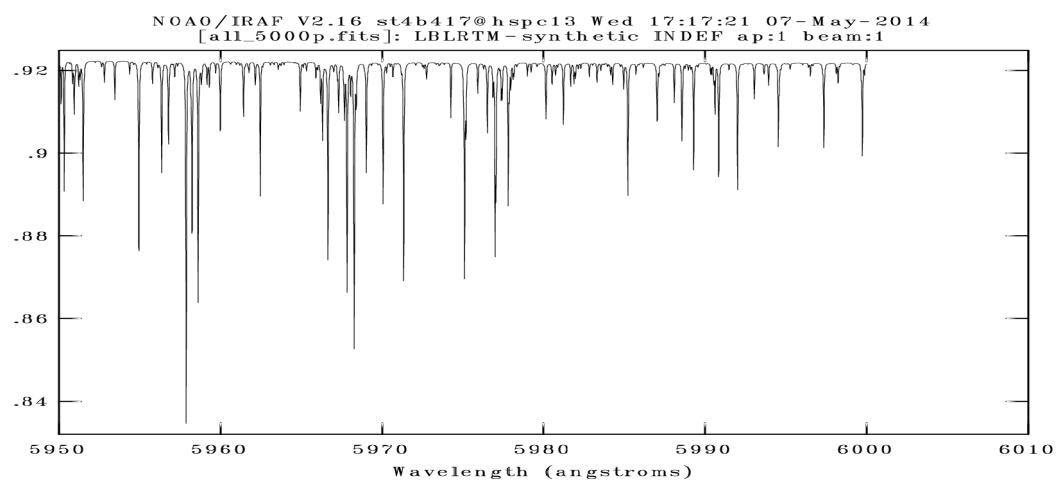
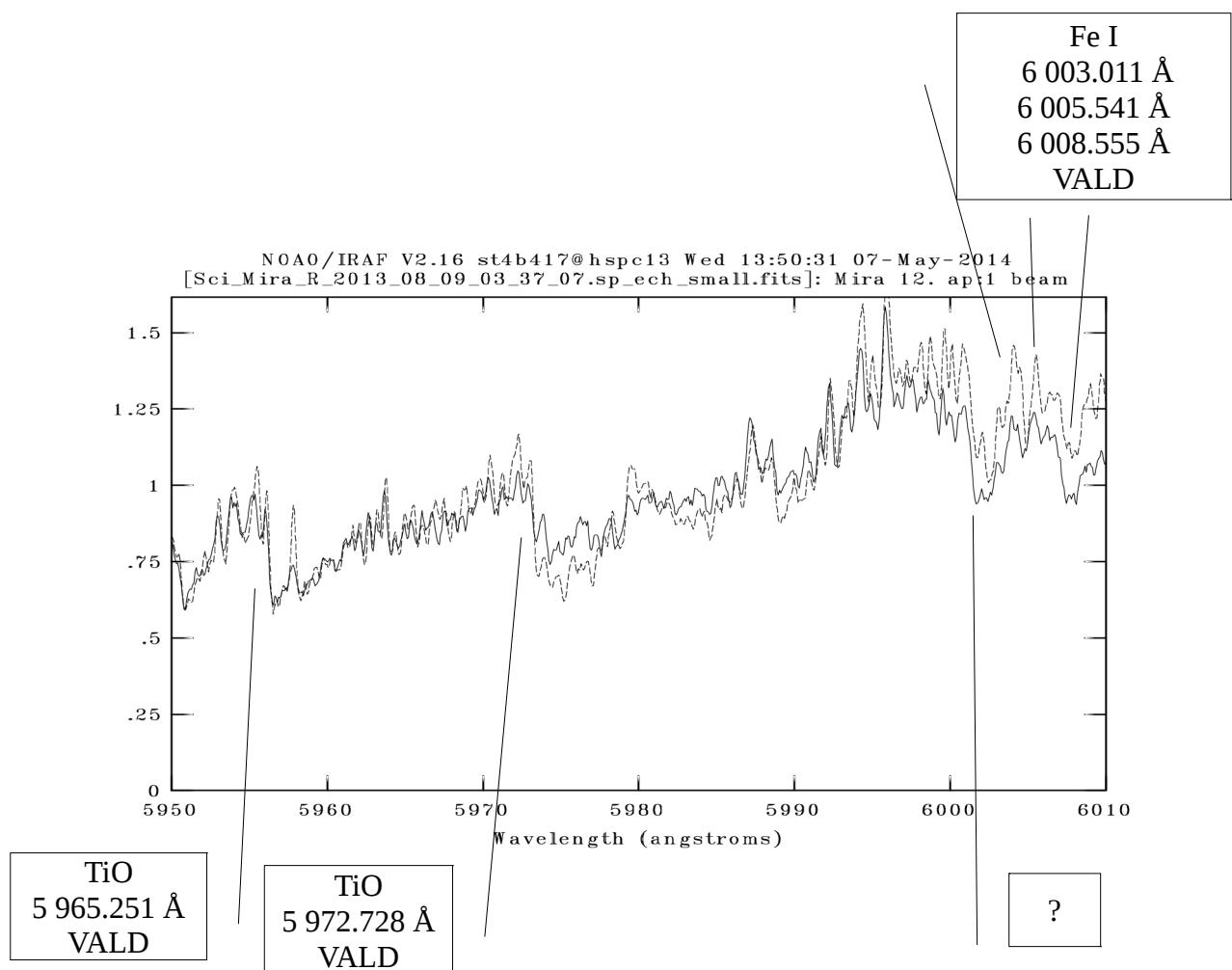
noise due to the transition from blue to red spectral range



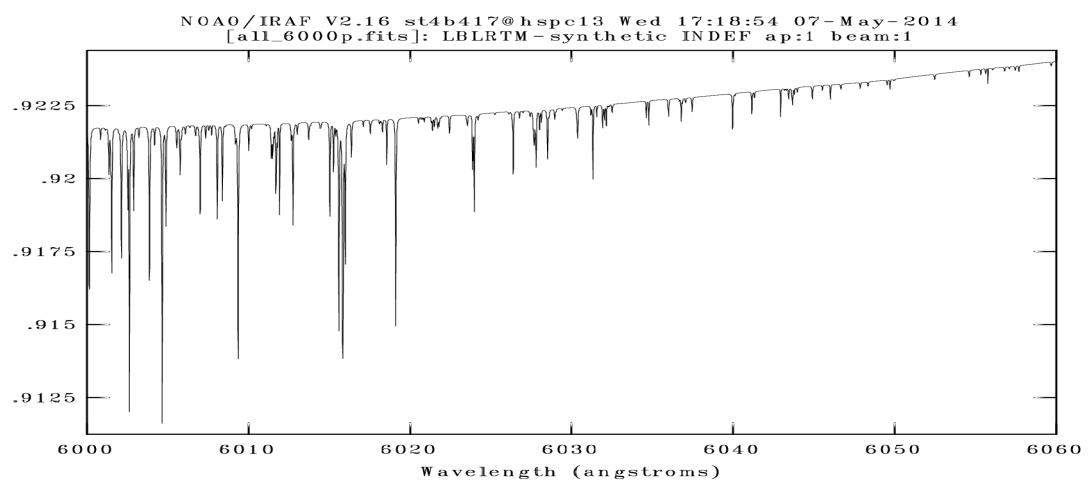
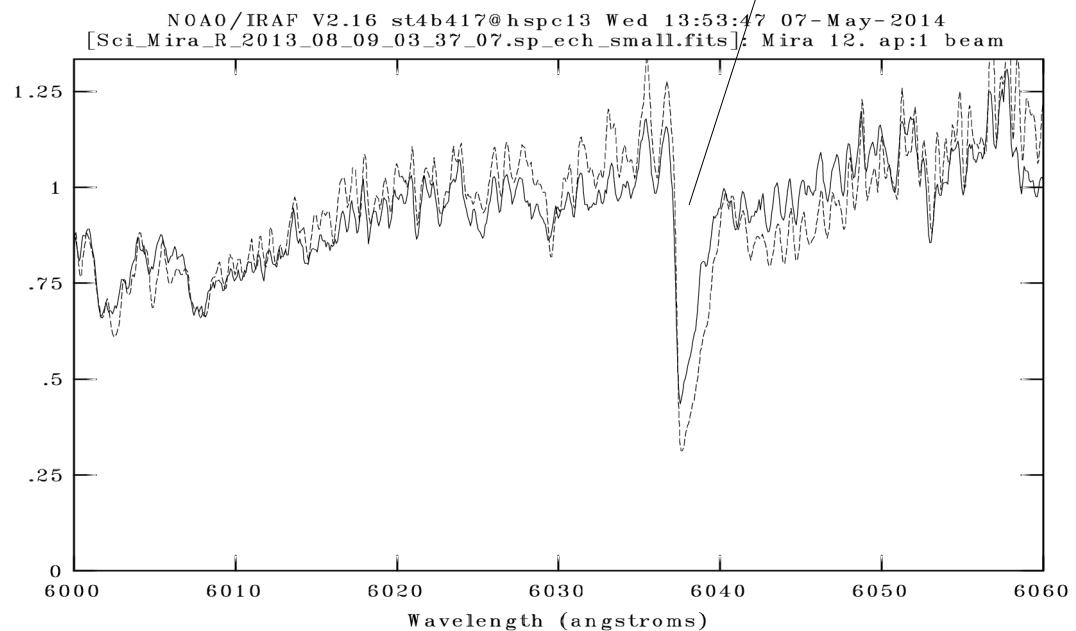
Appendix
B: Red spectrum (5829Å-8817Å)

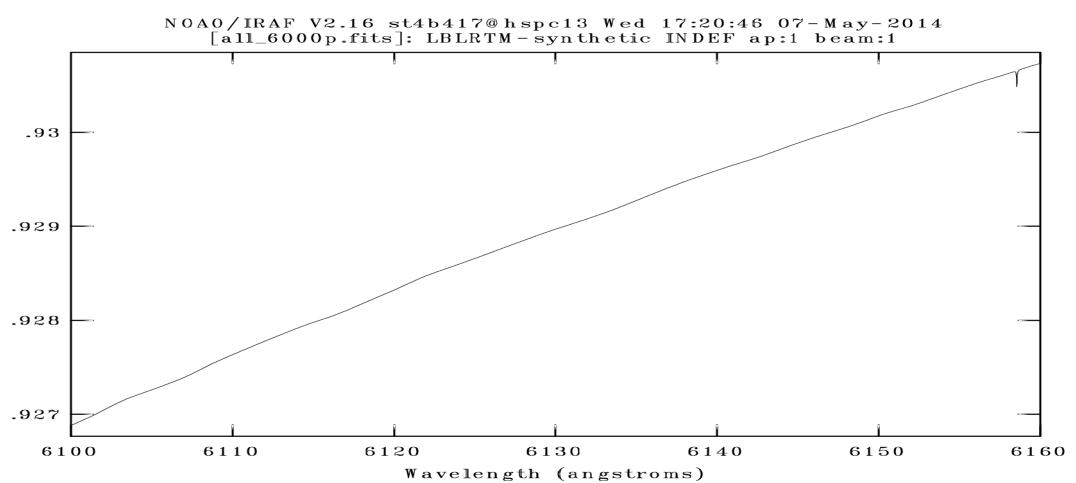
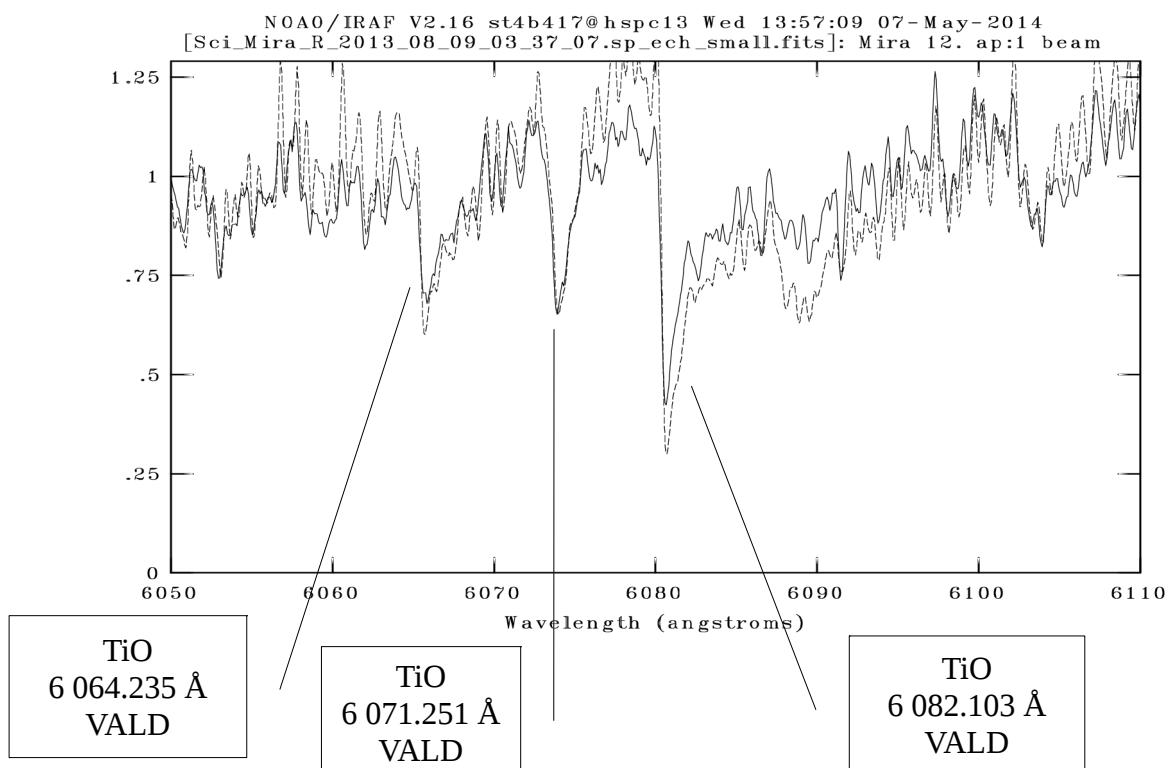


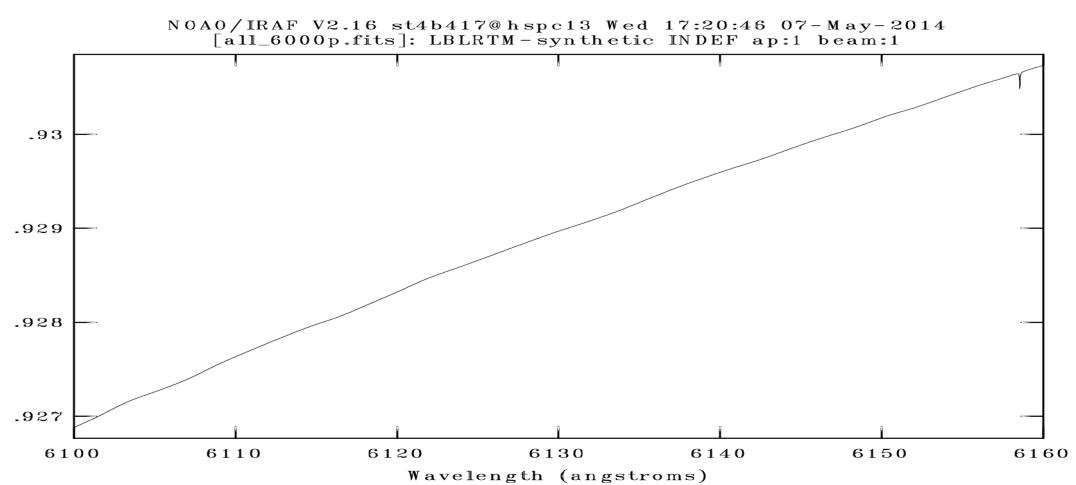
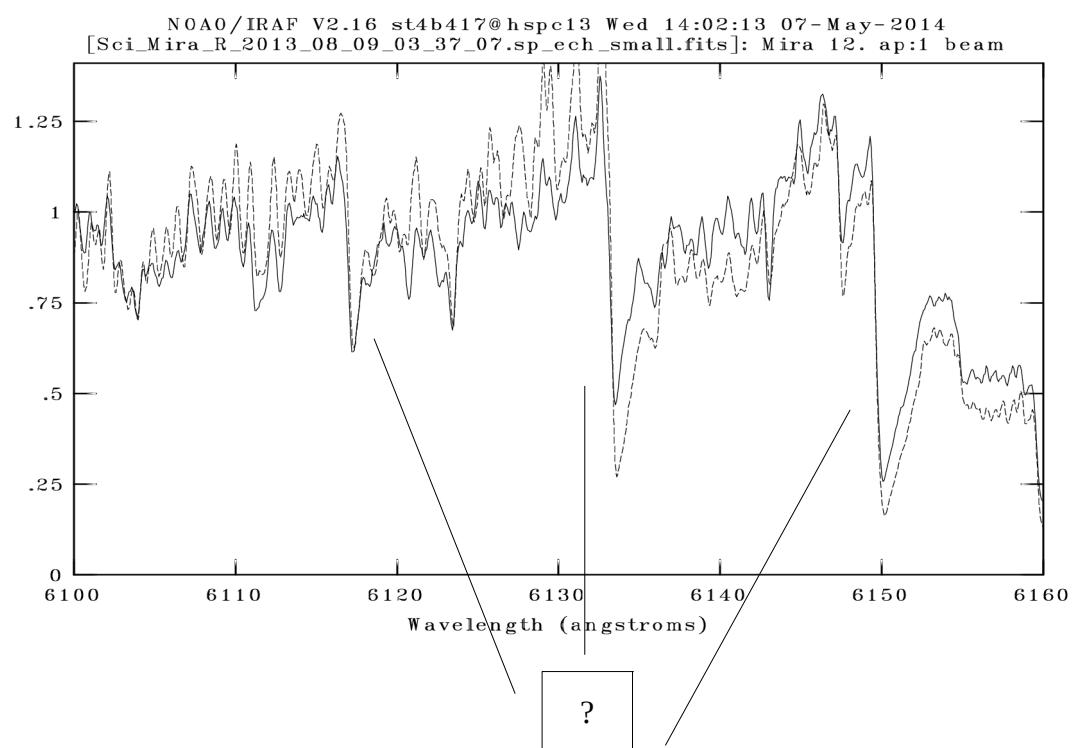




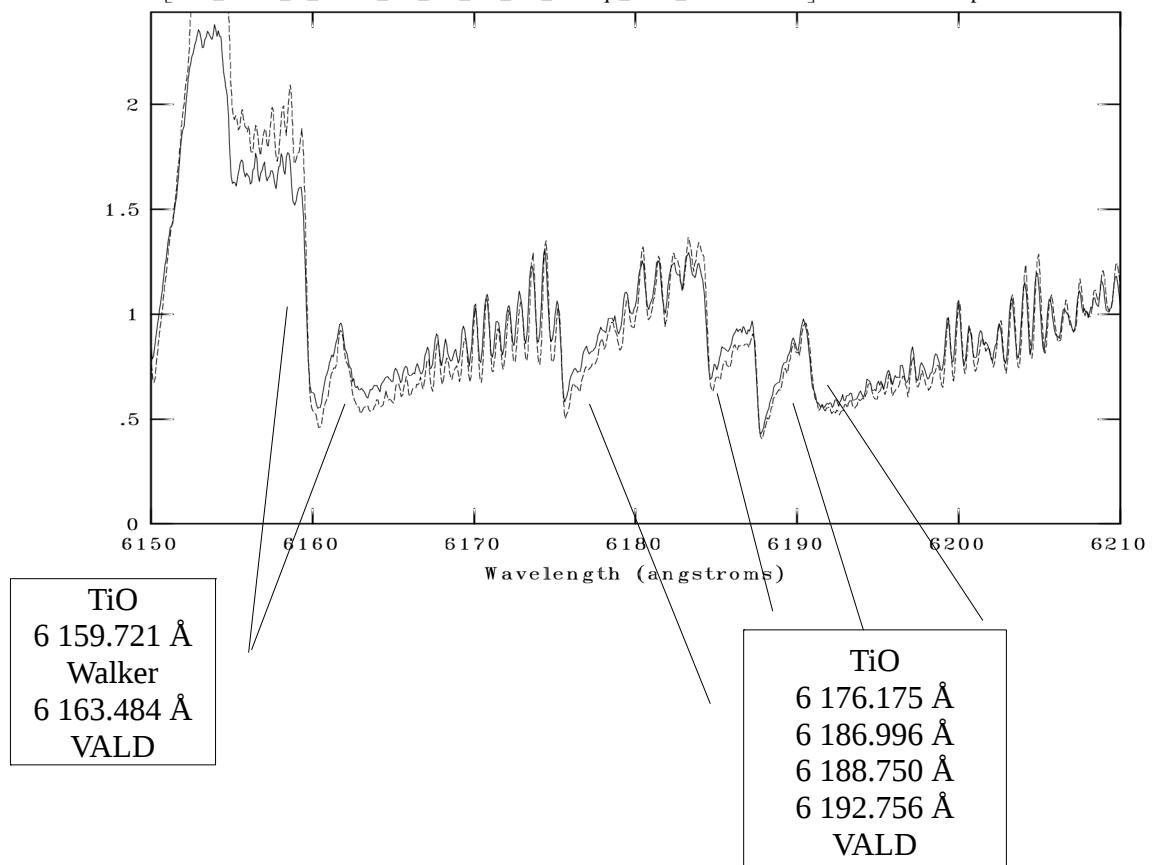
TiO
6 037.453 Å
VALD



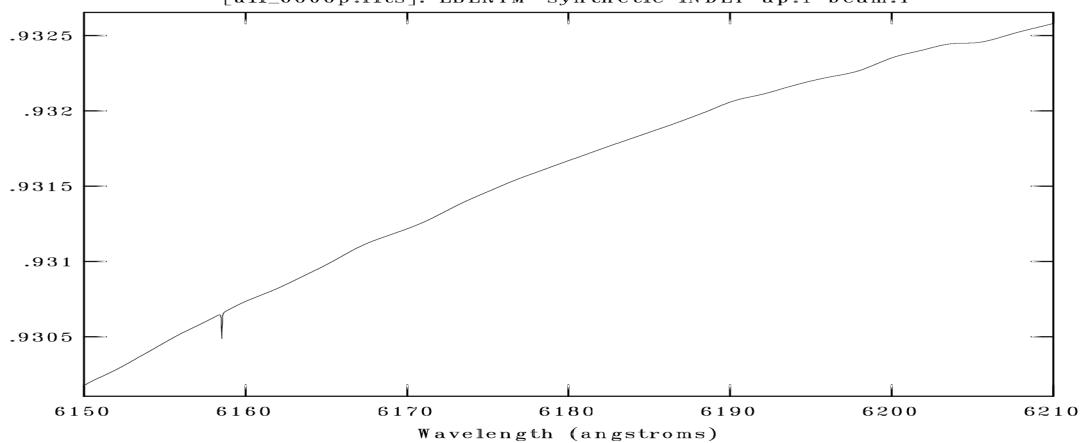


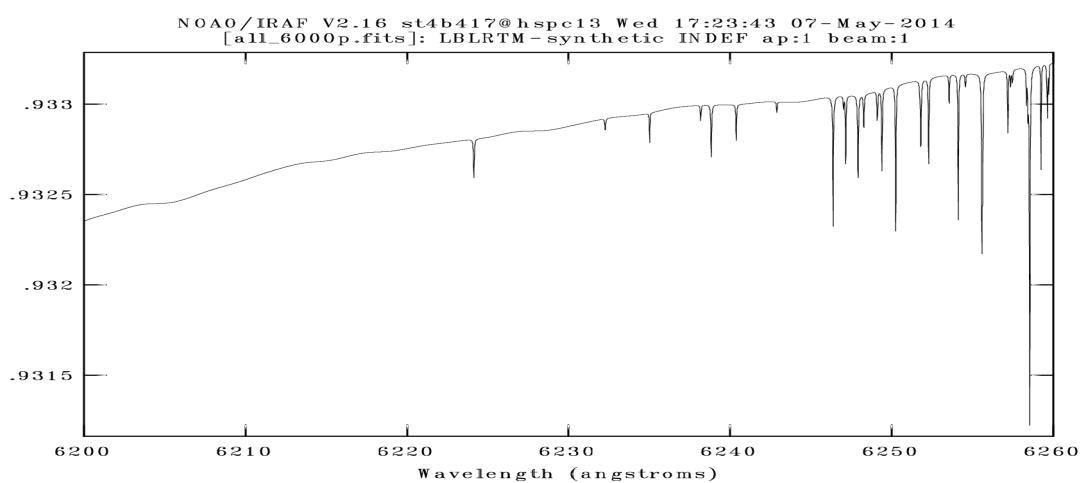
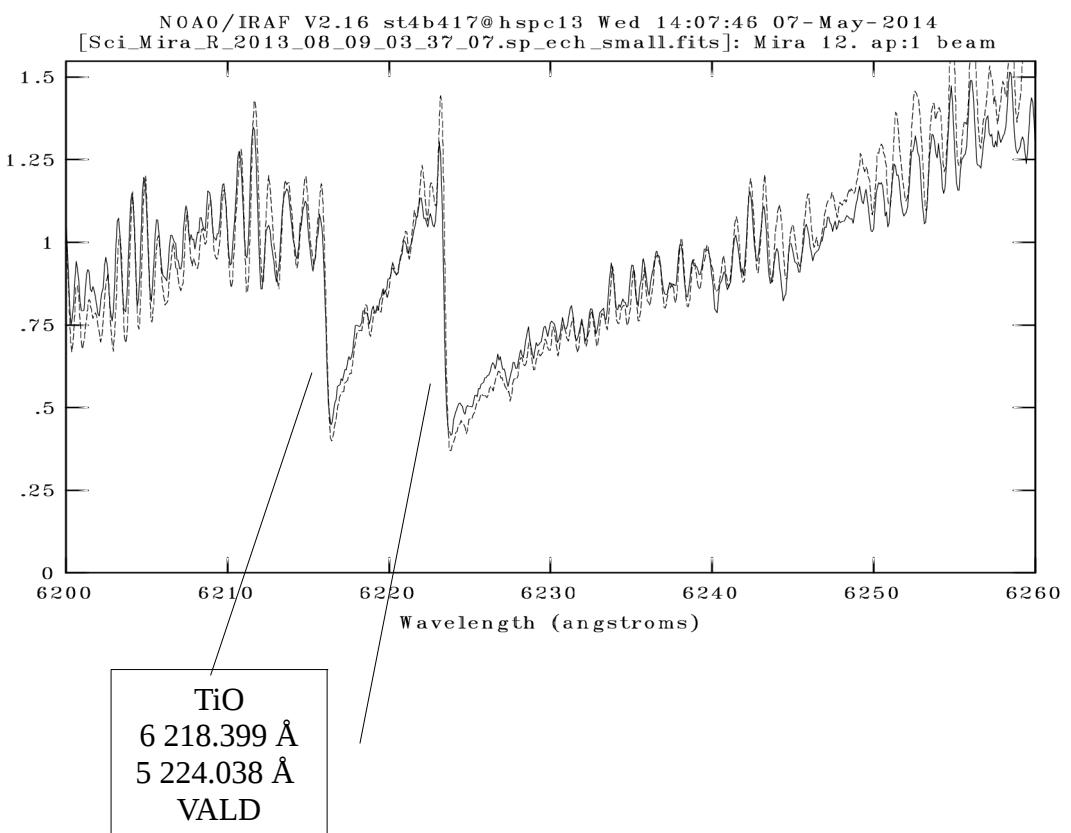


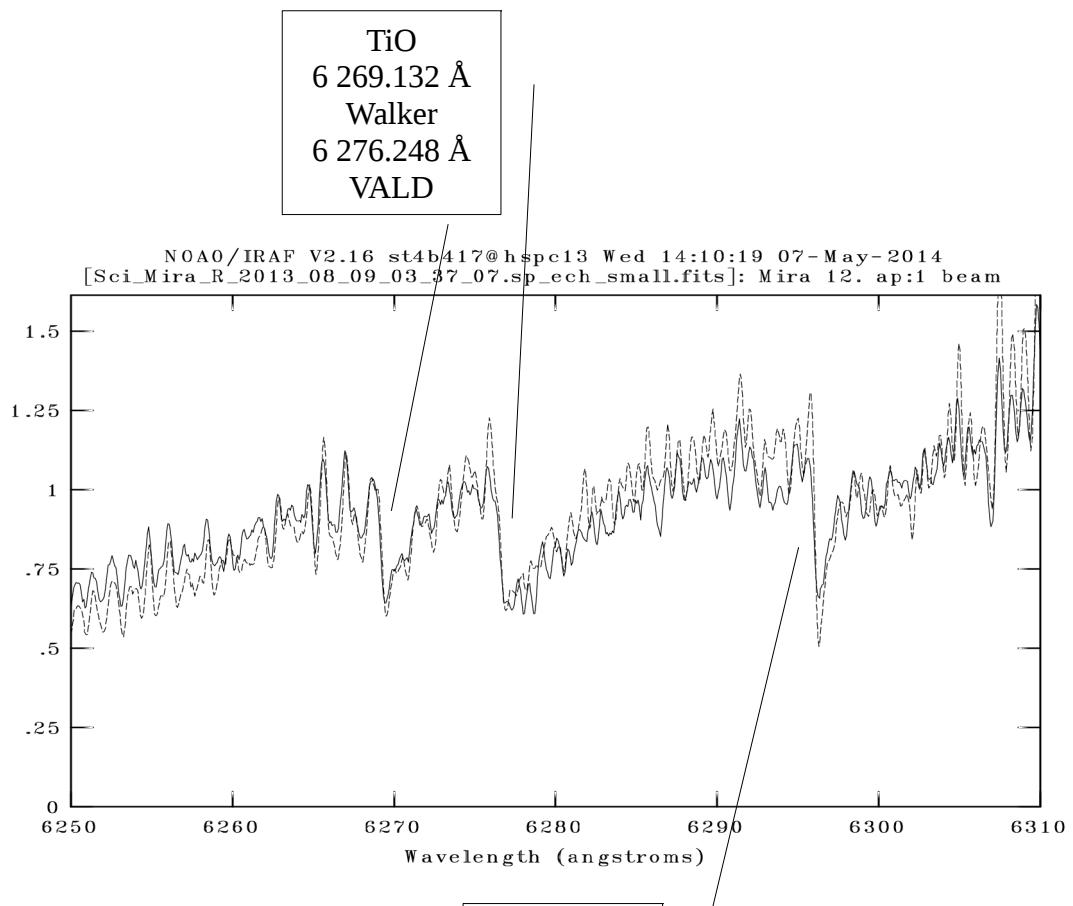
NOAO/IRAF V2.16 st4b417@hspc13 Wed 14:04:42 07-May-2014
[Sci_Mira_R_2013_08_09_03_37_07.sp_ech_small.fits]: Mira 12. ap:1 beam



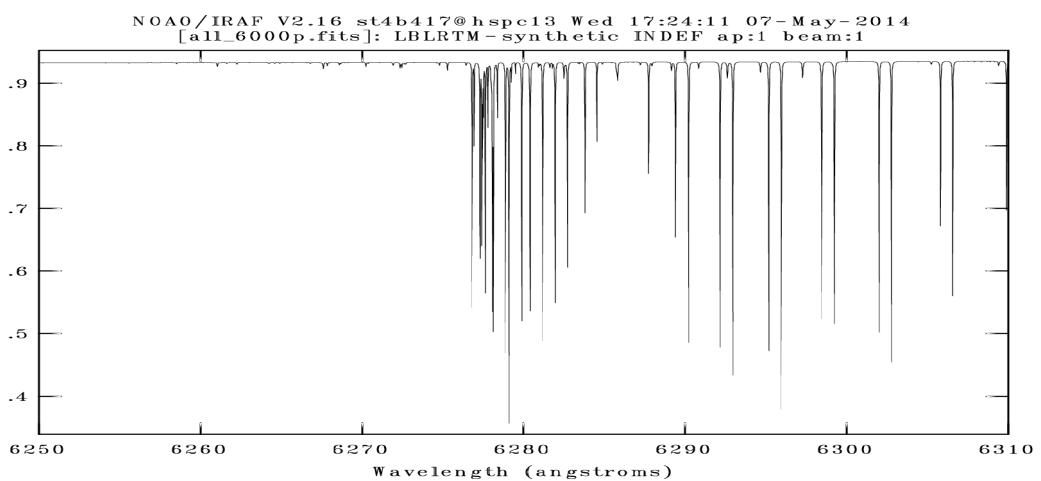
NOAO/IRAF V2.16 st4b417@hspc13 Wed 17:23:12 07-May-2014
[all_6000p.fits]: LBLRTM - synthetic INDEF ap:1 beam:1

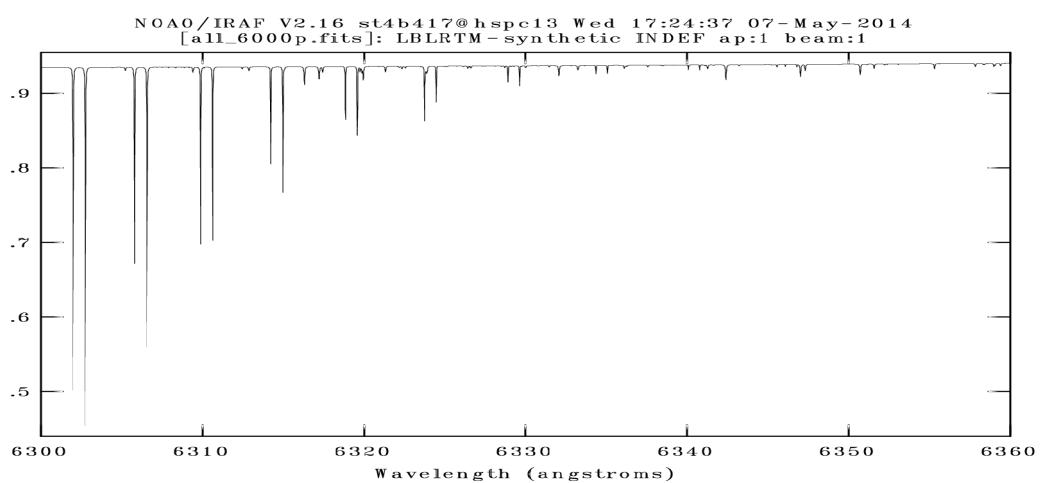
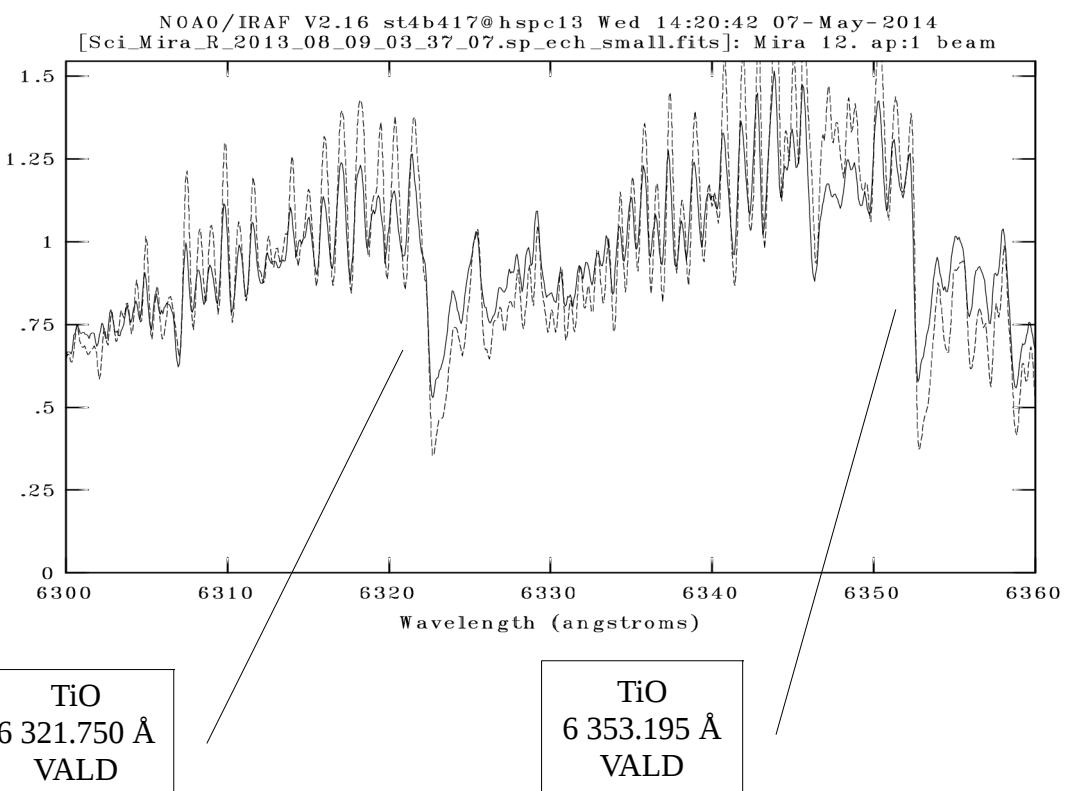


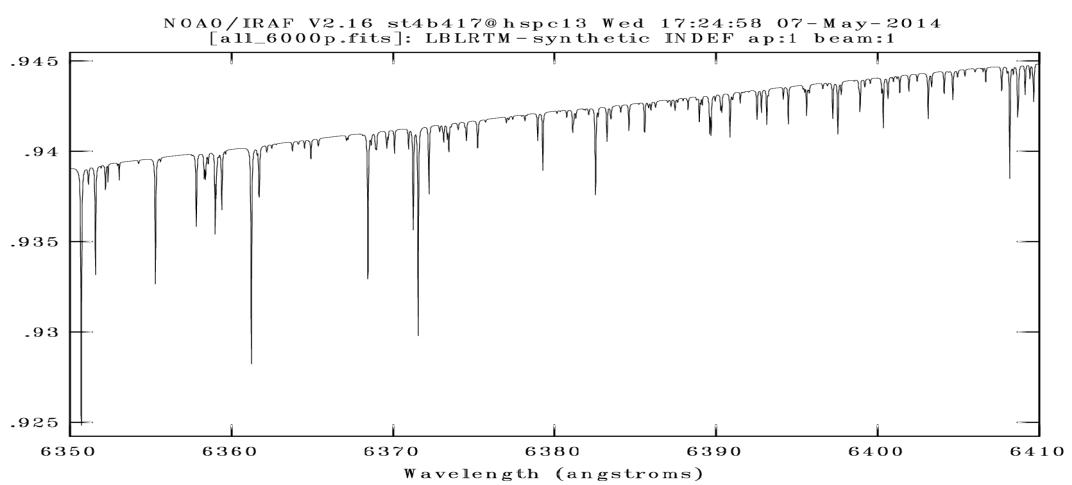
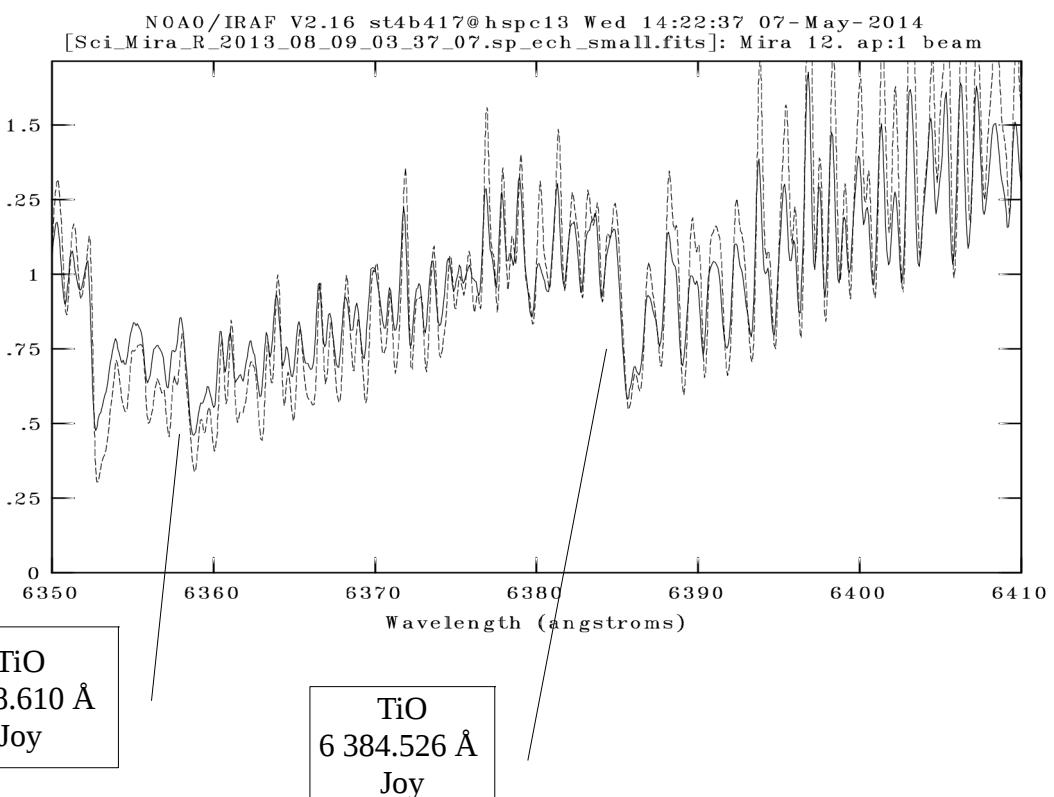


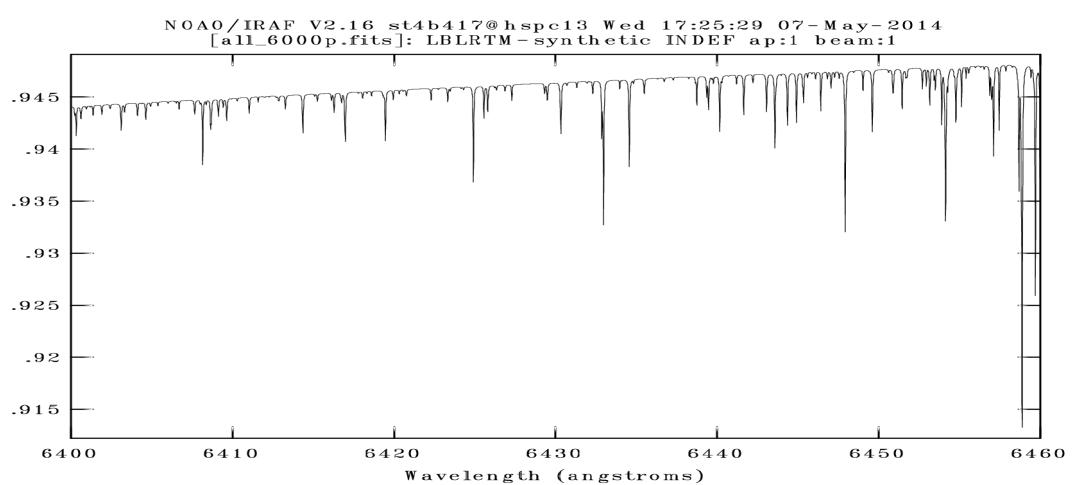
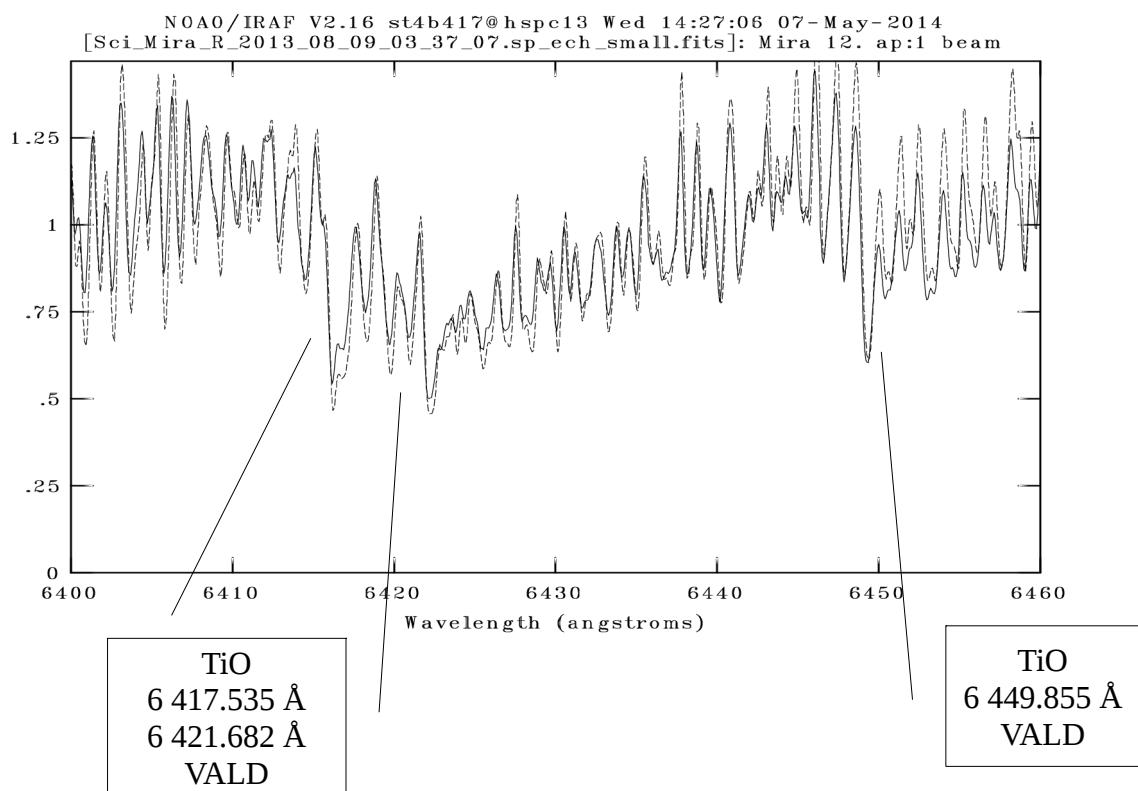


TiO
6 297.720 Å
VALD

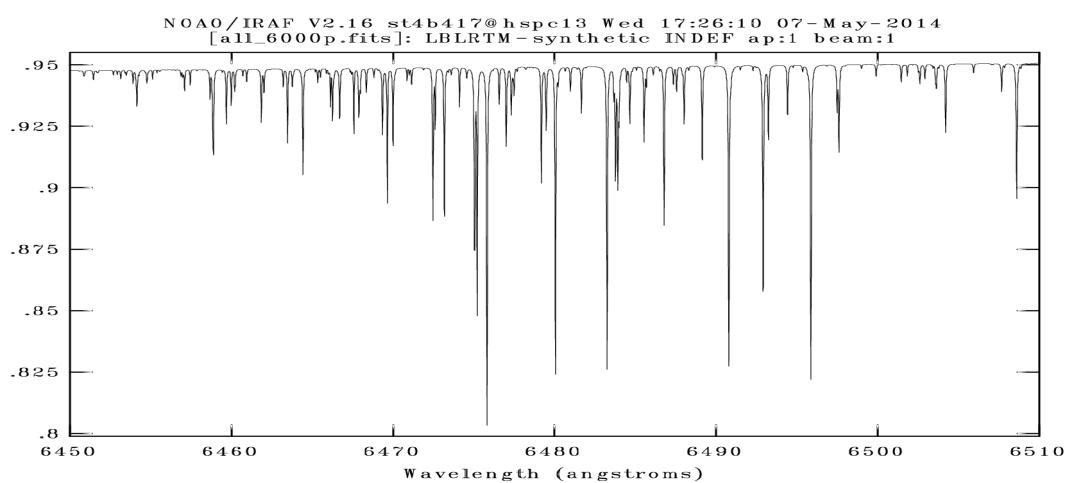
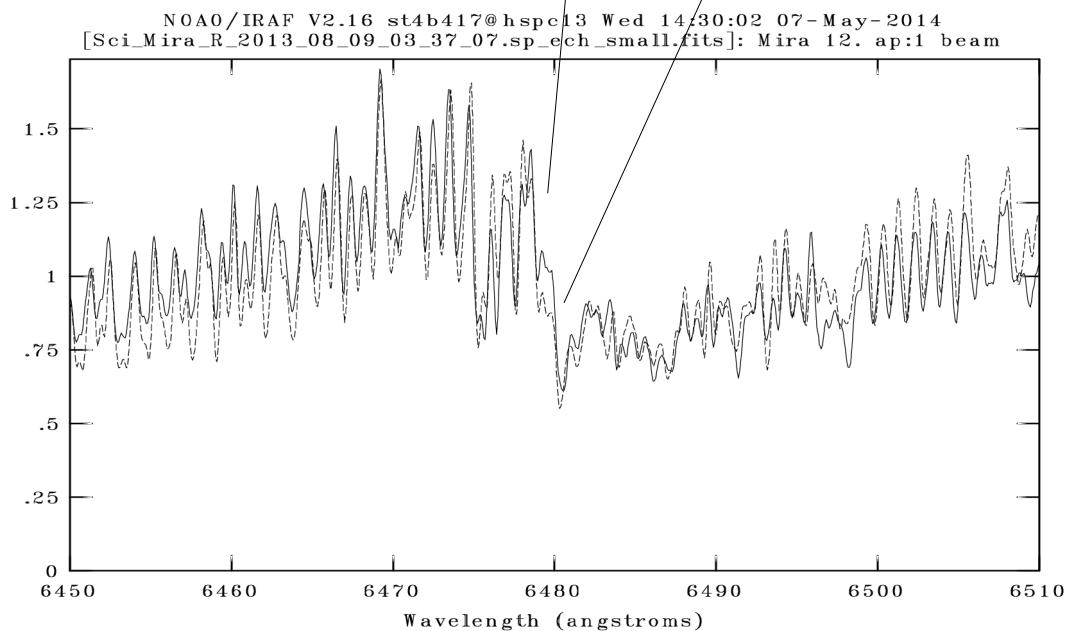


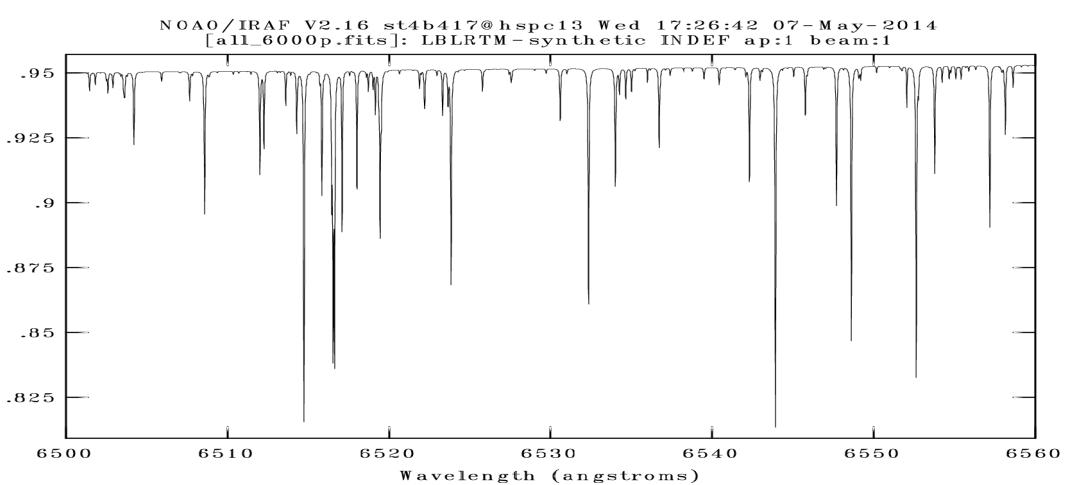
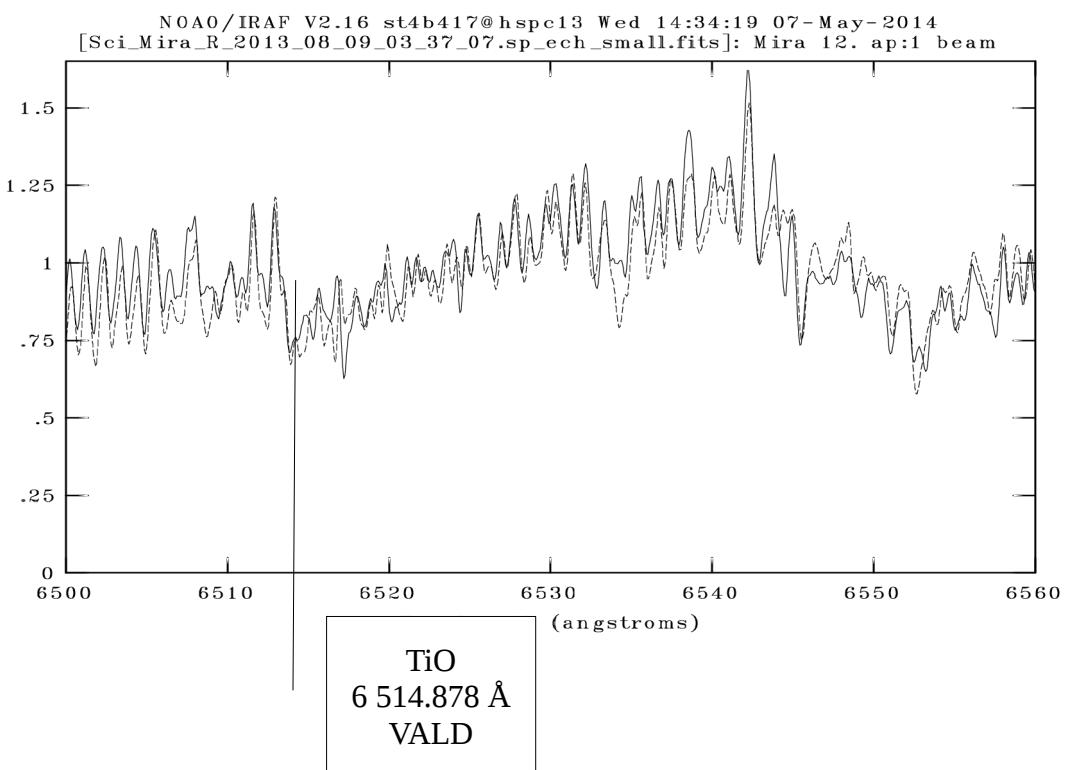


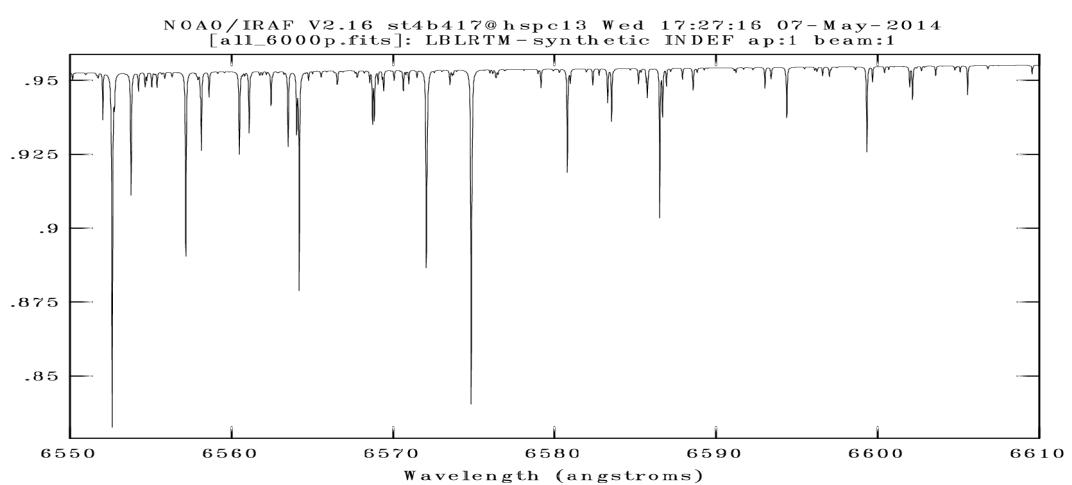
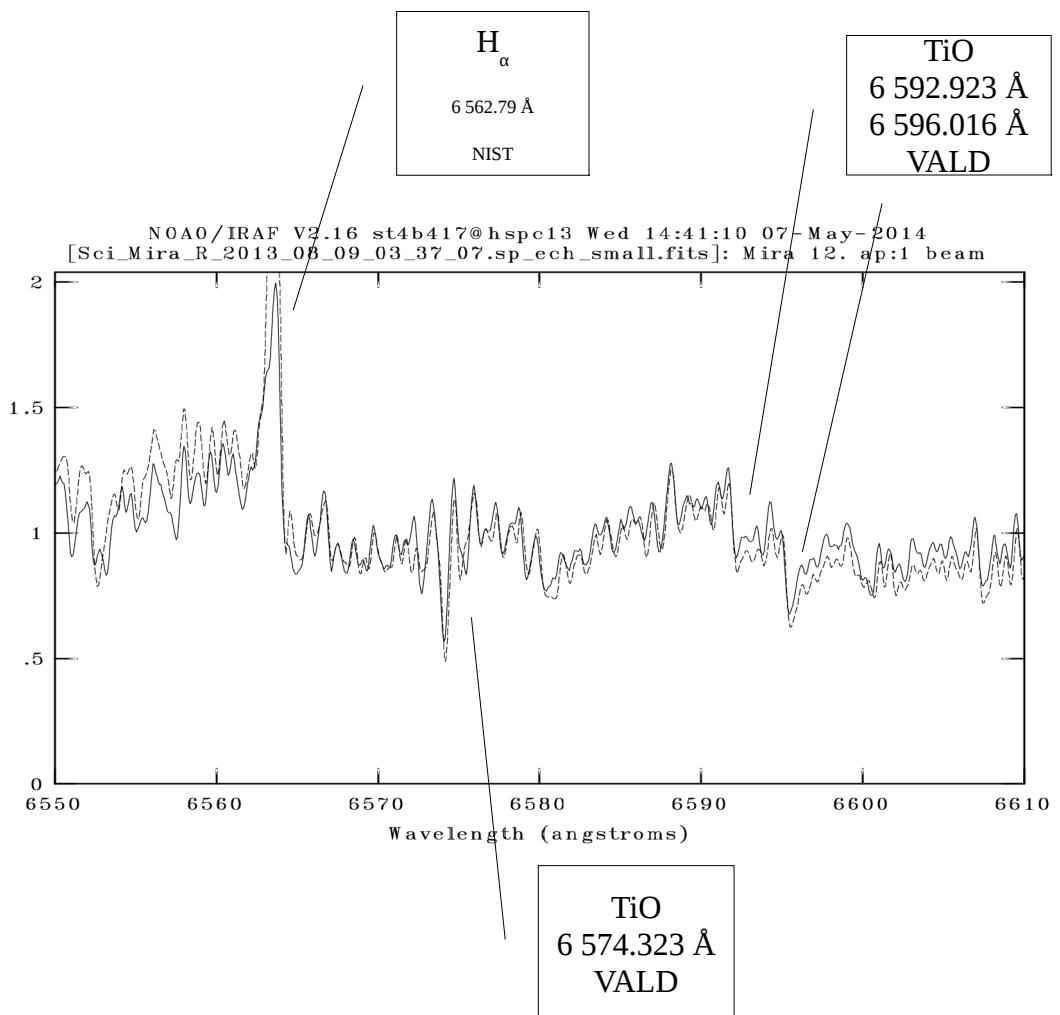


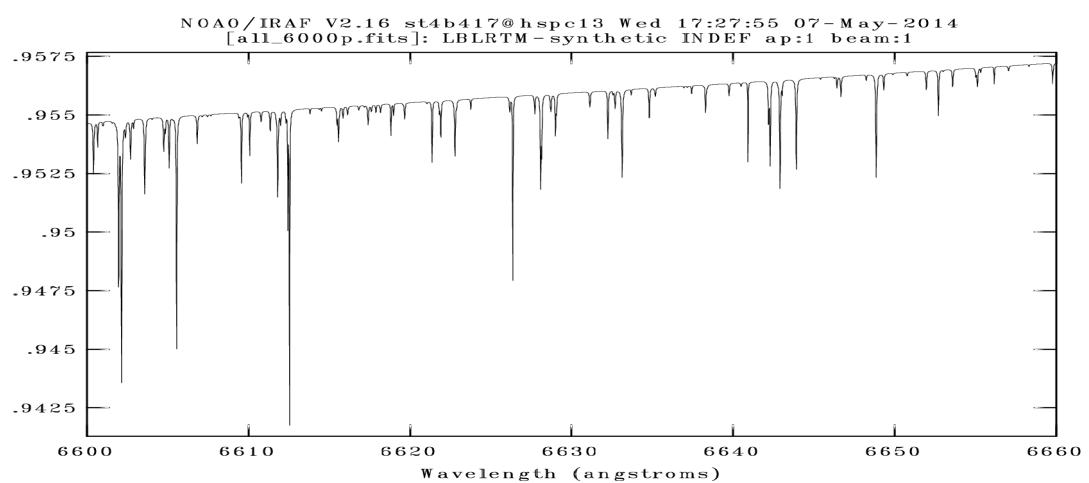
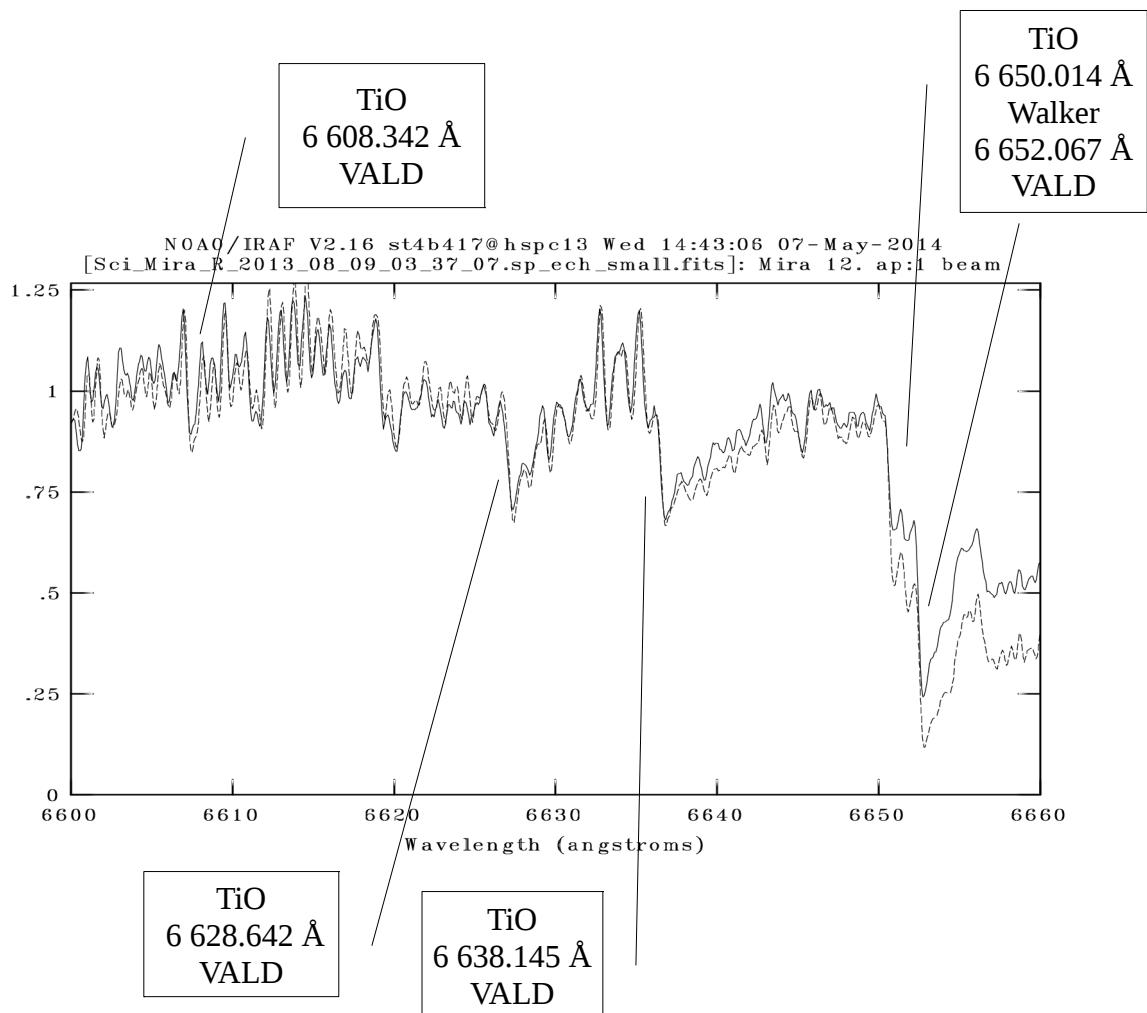


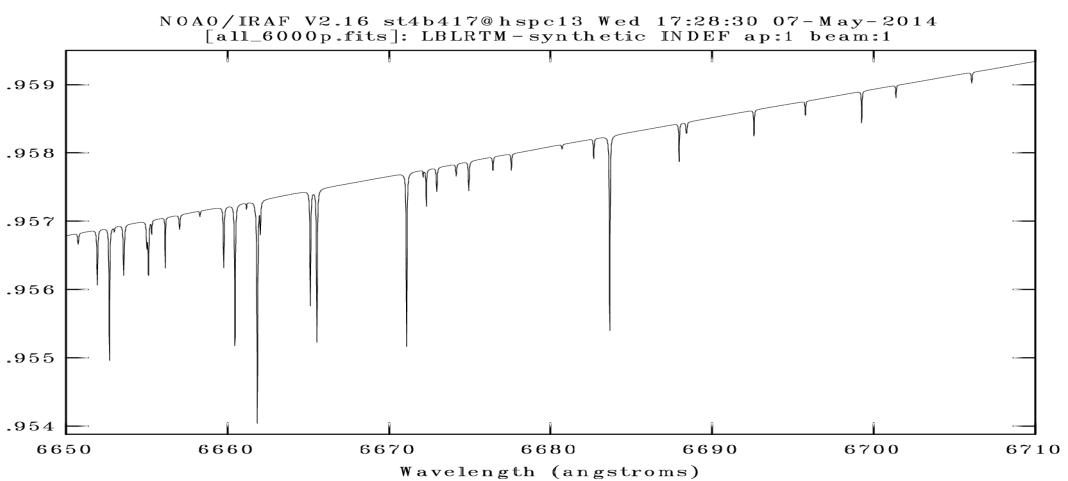
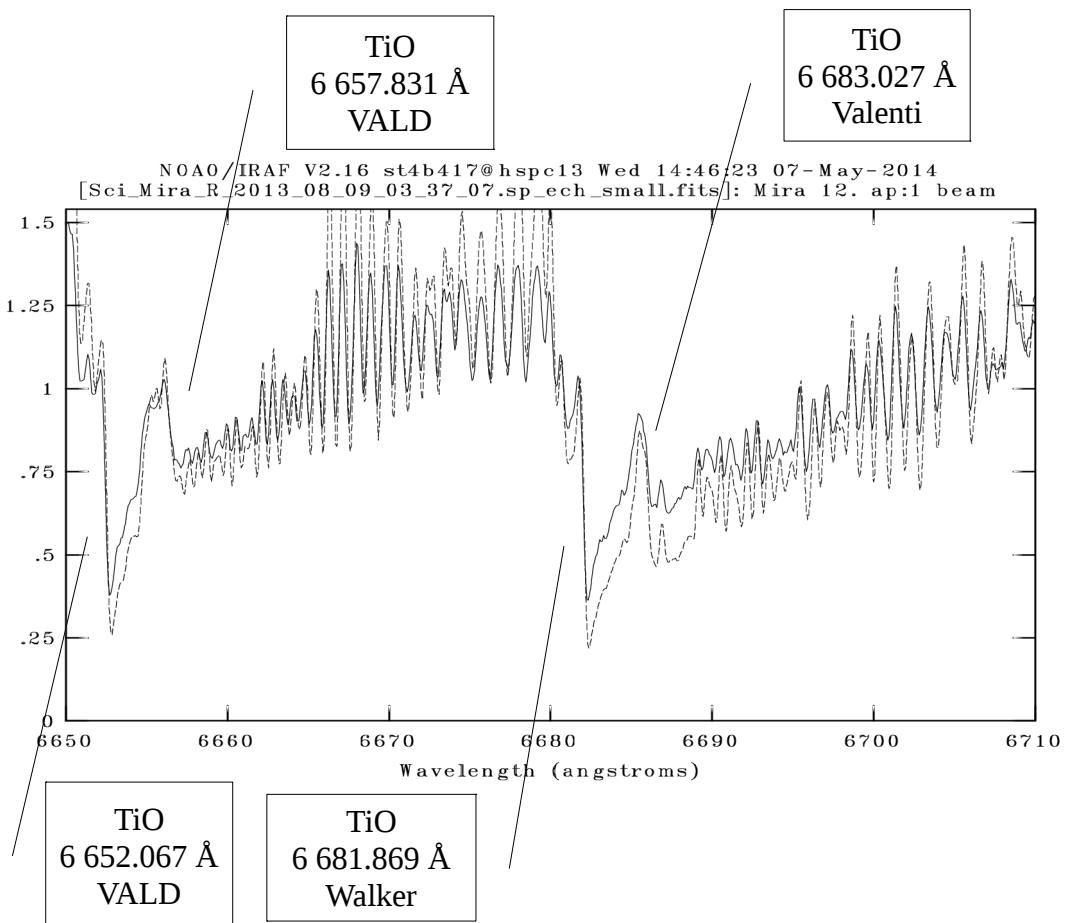
TiO
 6 478.108 Å
 Walker
 6 479.972 Å
 VALD

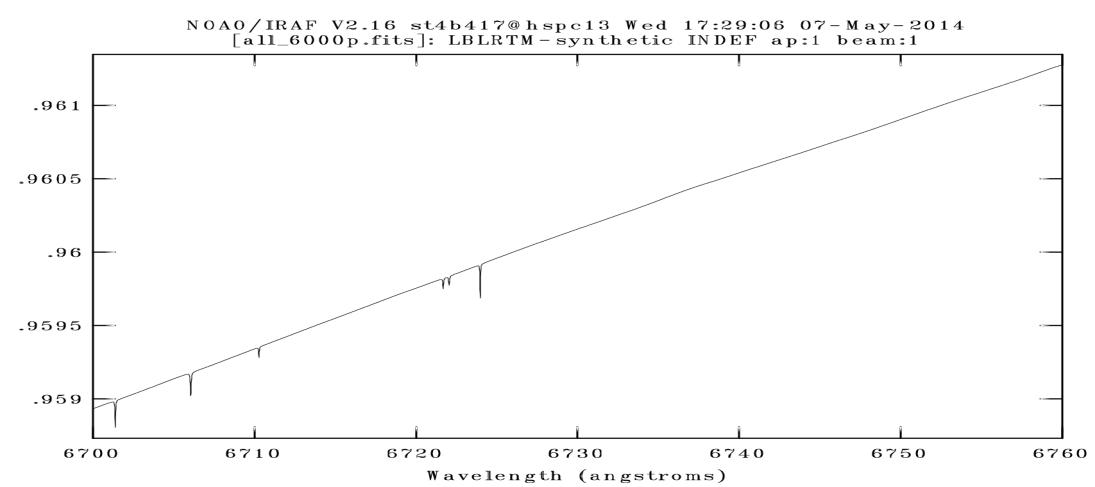
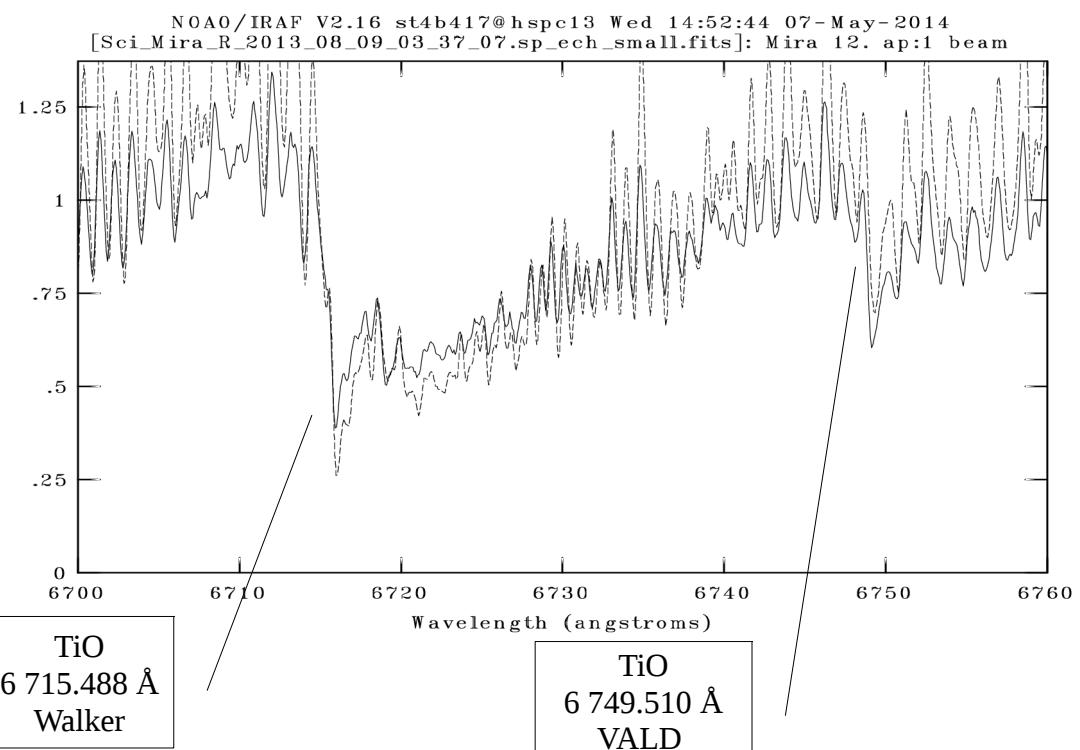


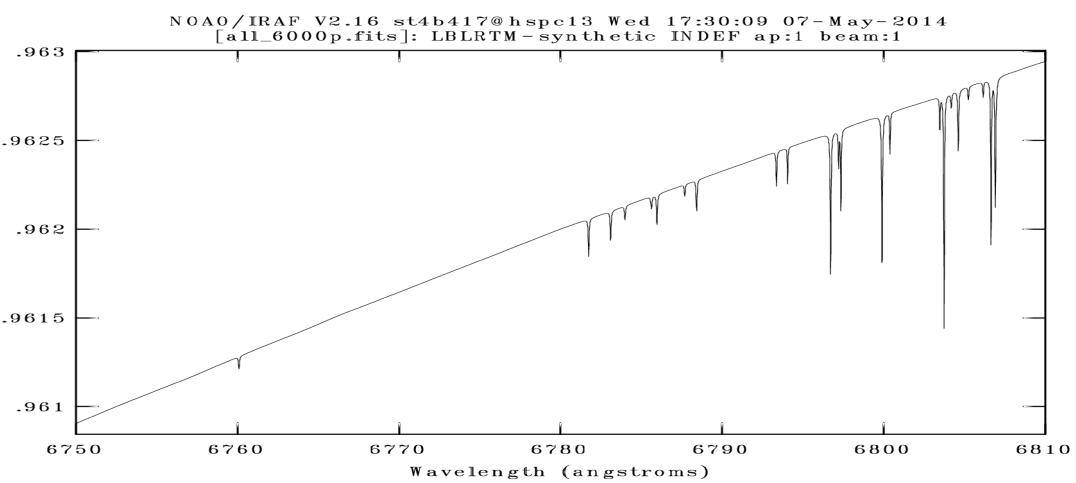
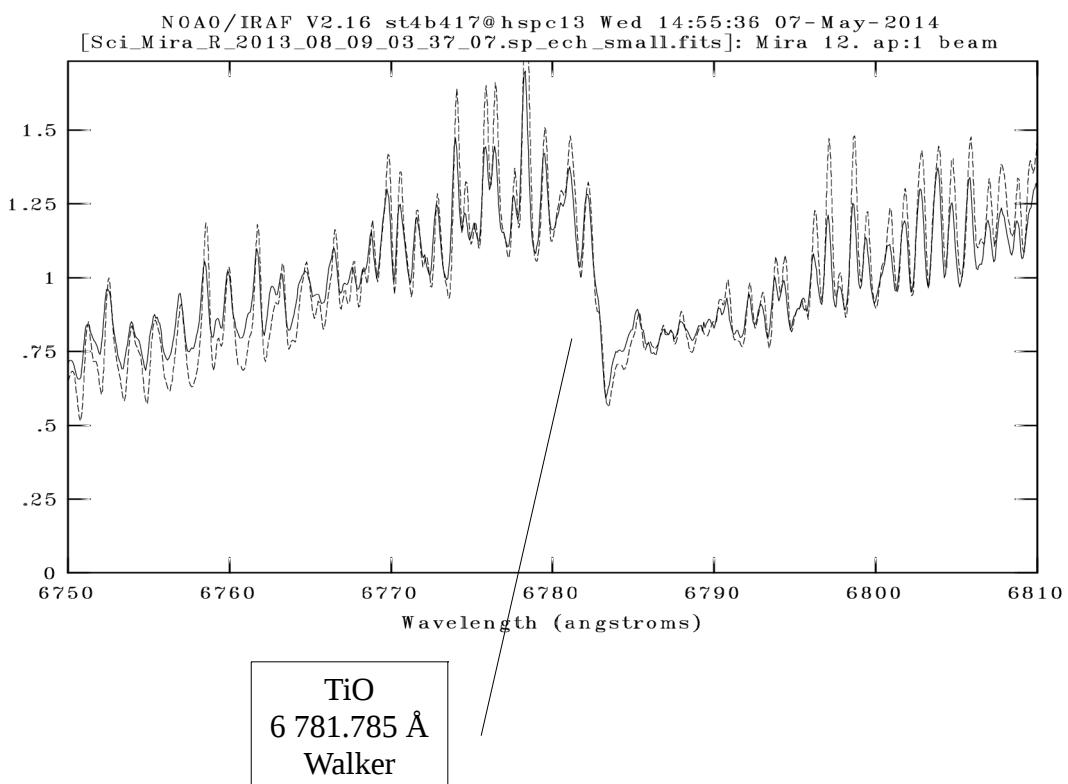


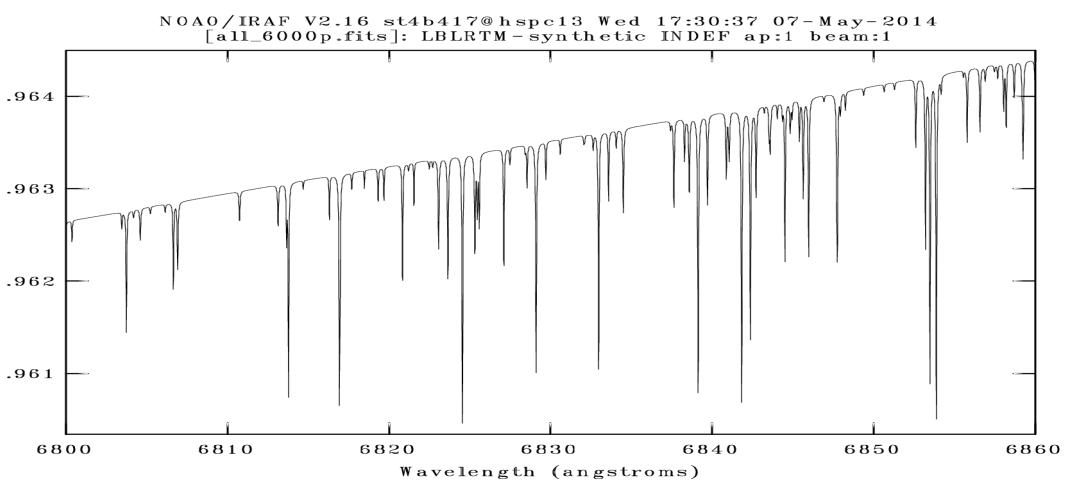
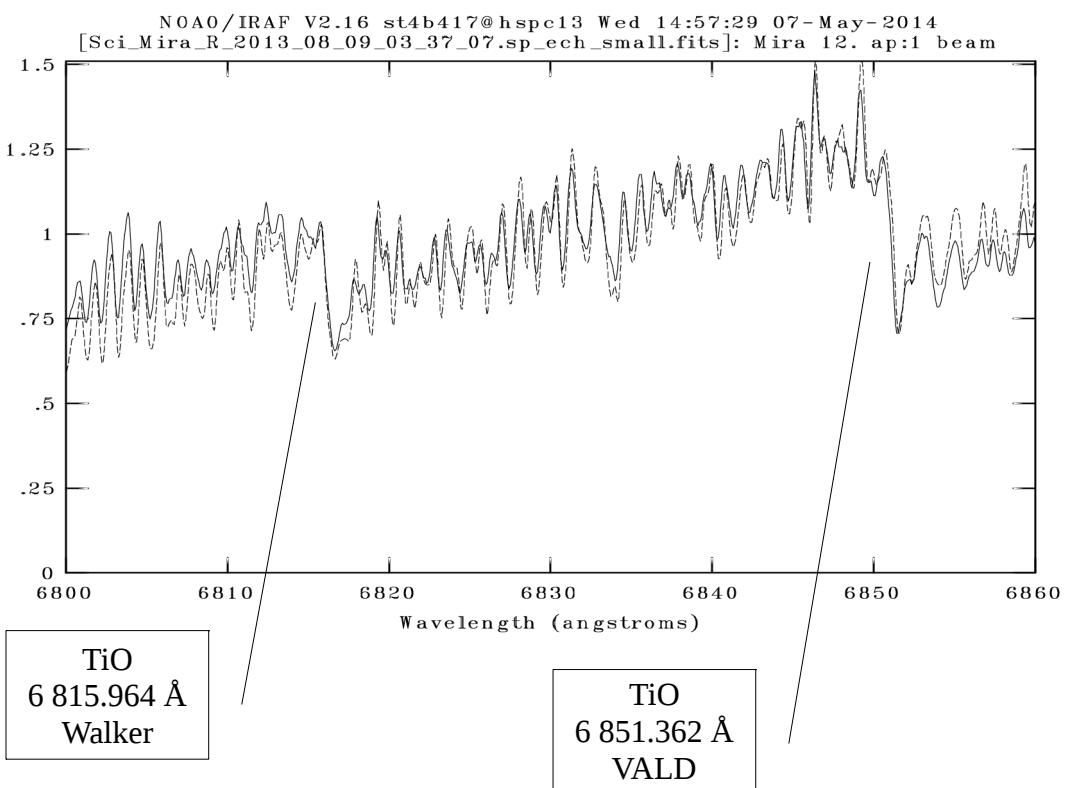


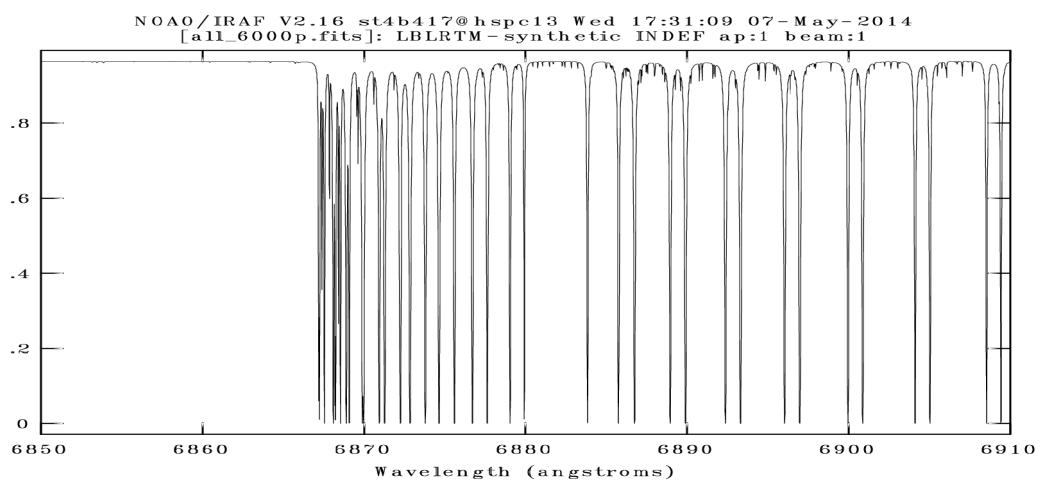
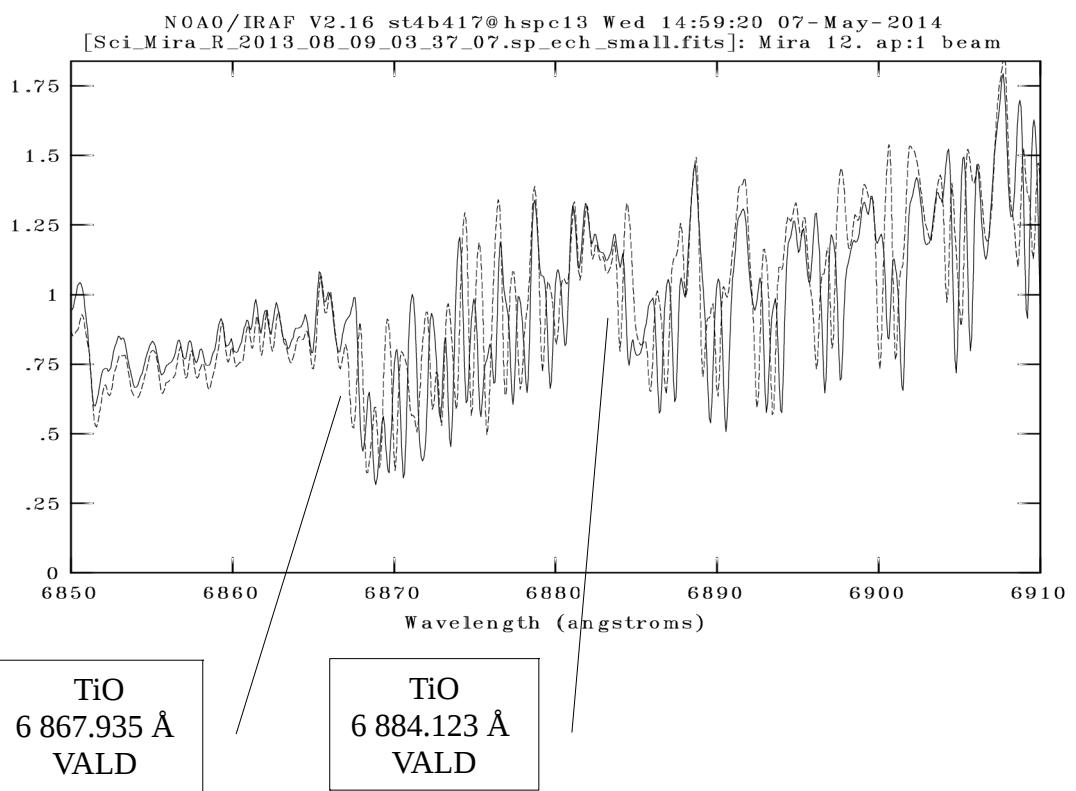




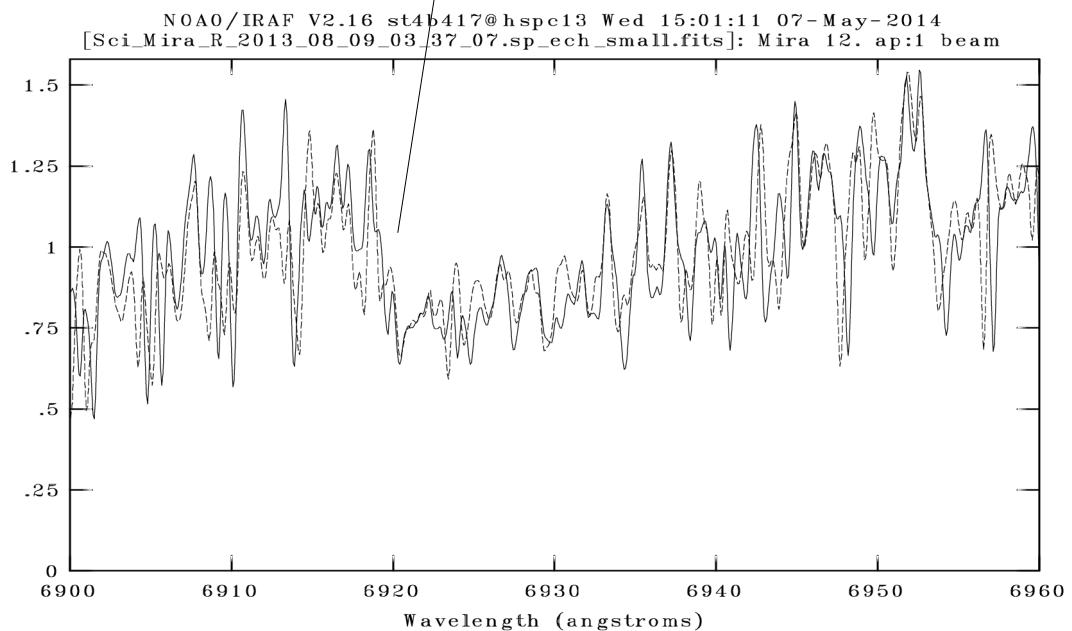




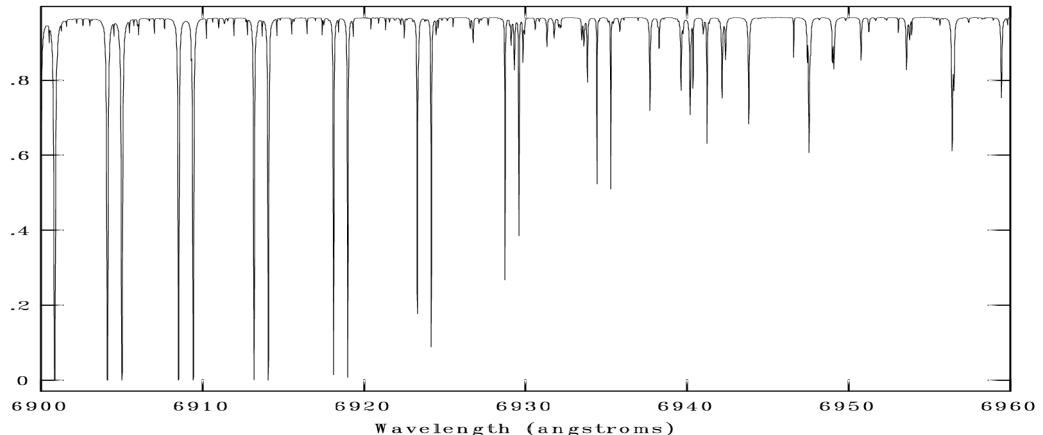




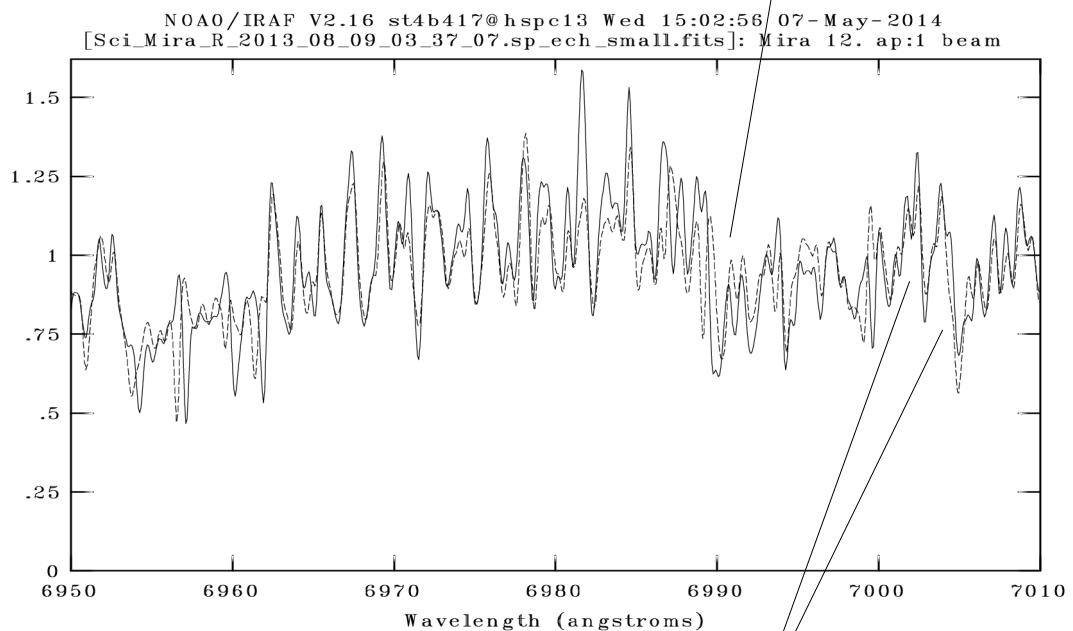
TiO
6919.344 Å
VALD



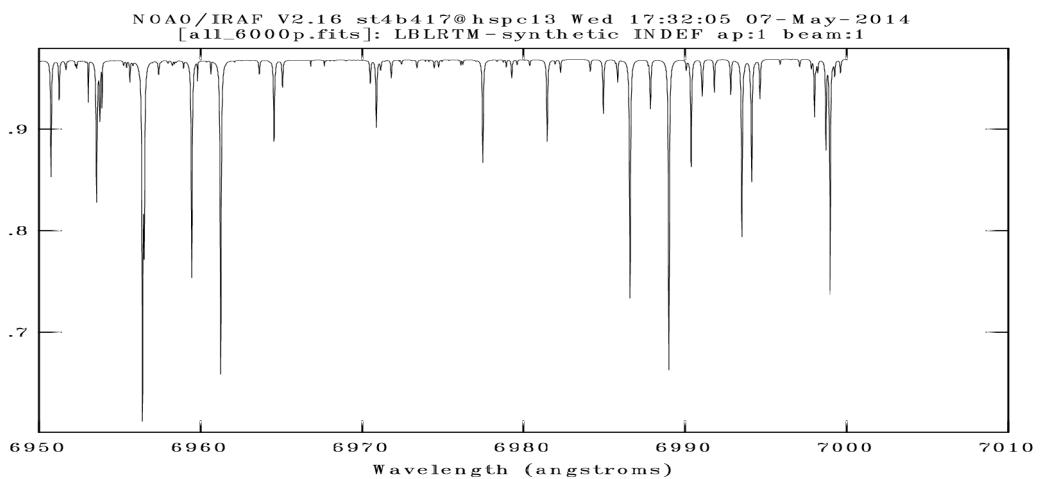
NOAO/IRAF V2.16 st4b417@hspc13 Wed 17:31:36 07-May-2014
[all_6000p.fits]: LBLRTM - synthetic INDEF ap:1 beam:1

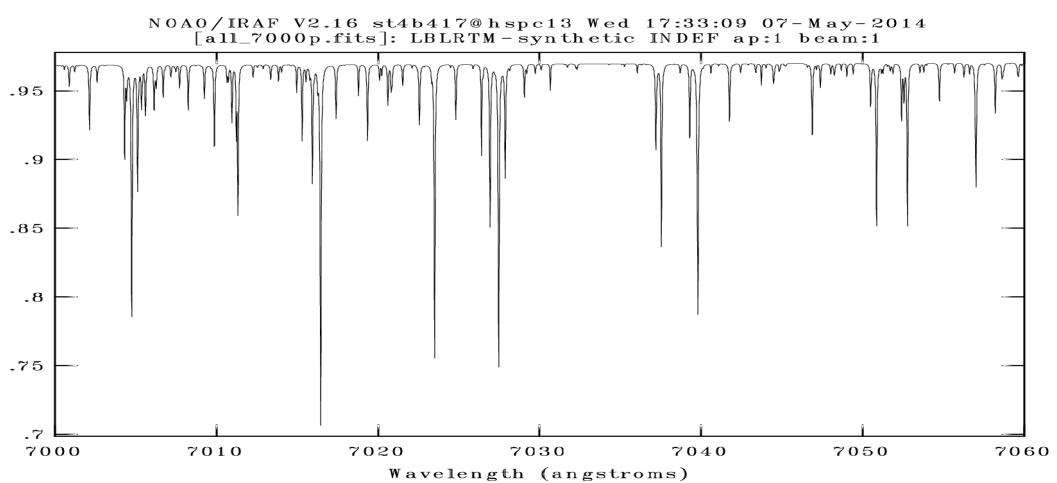
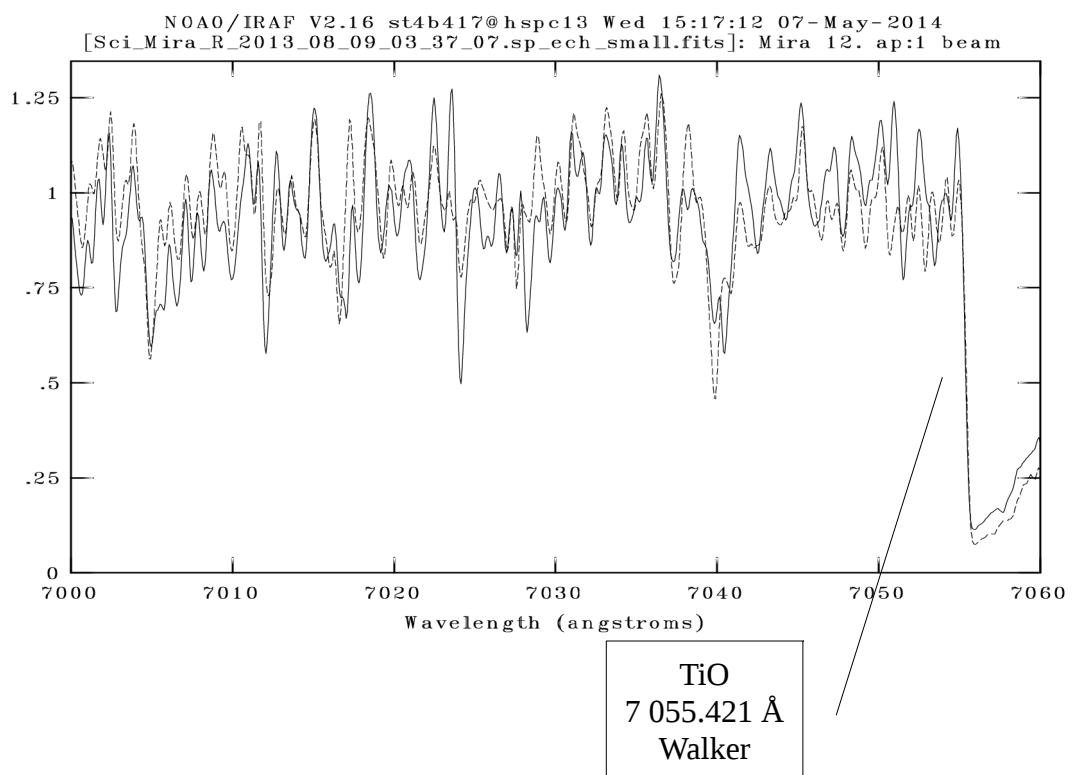


TiO
6989.646 Å
VALD

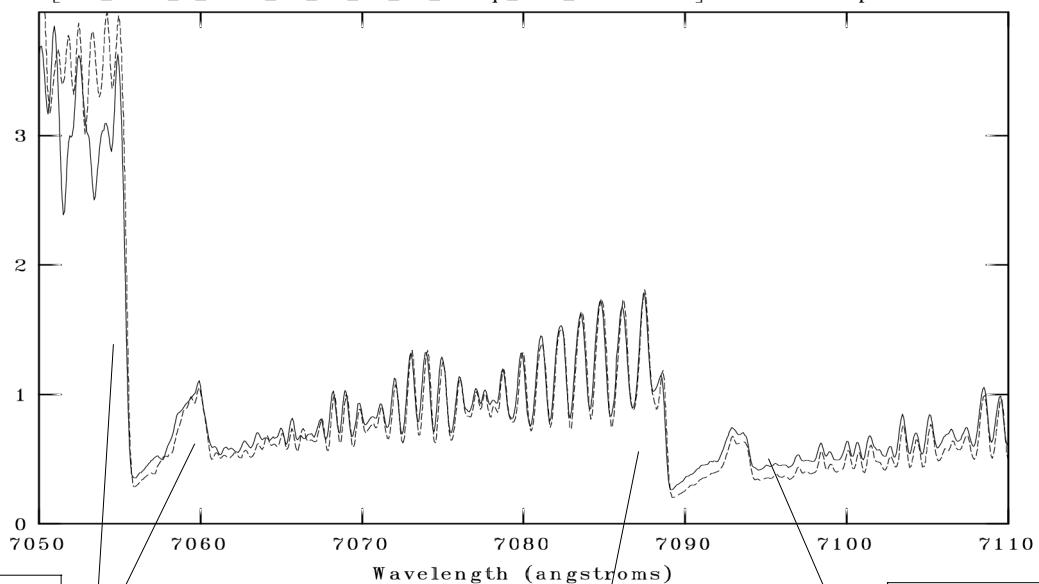


TiO
7002.101 Å
7005.324 Å
VALD

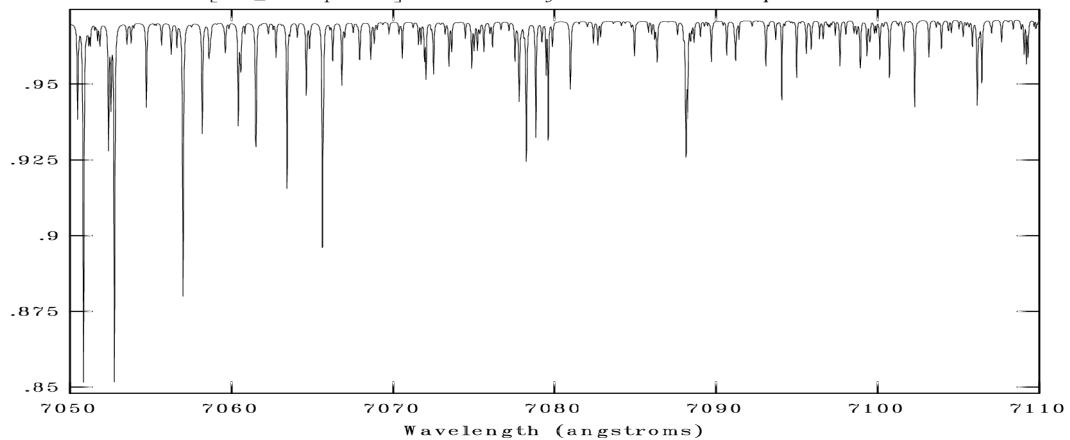


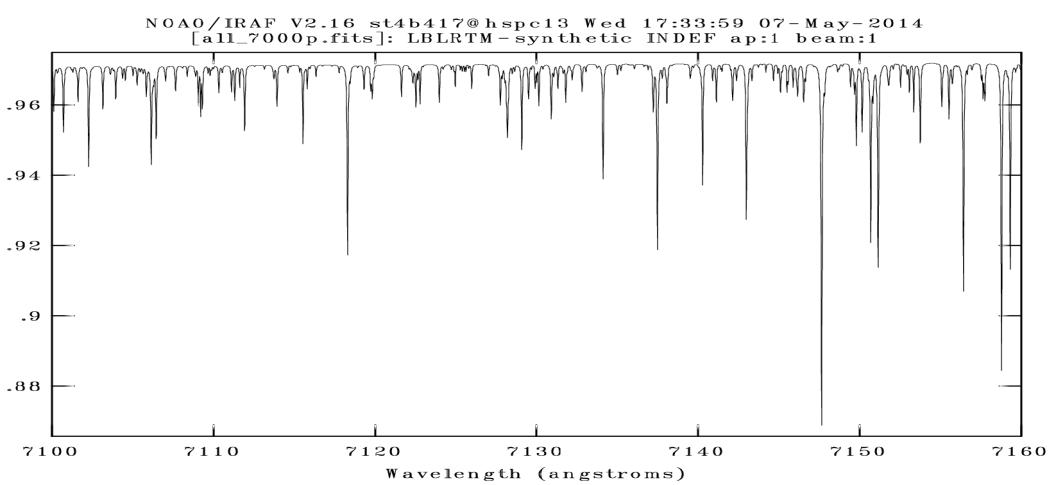
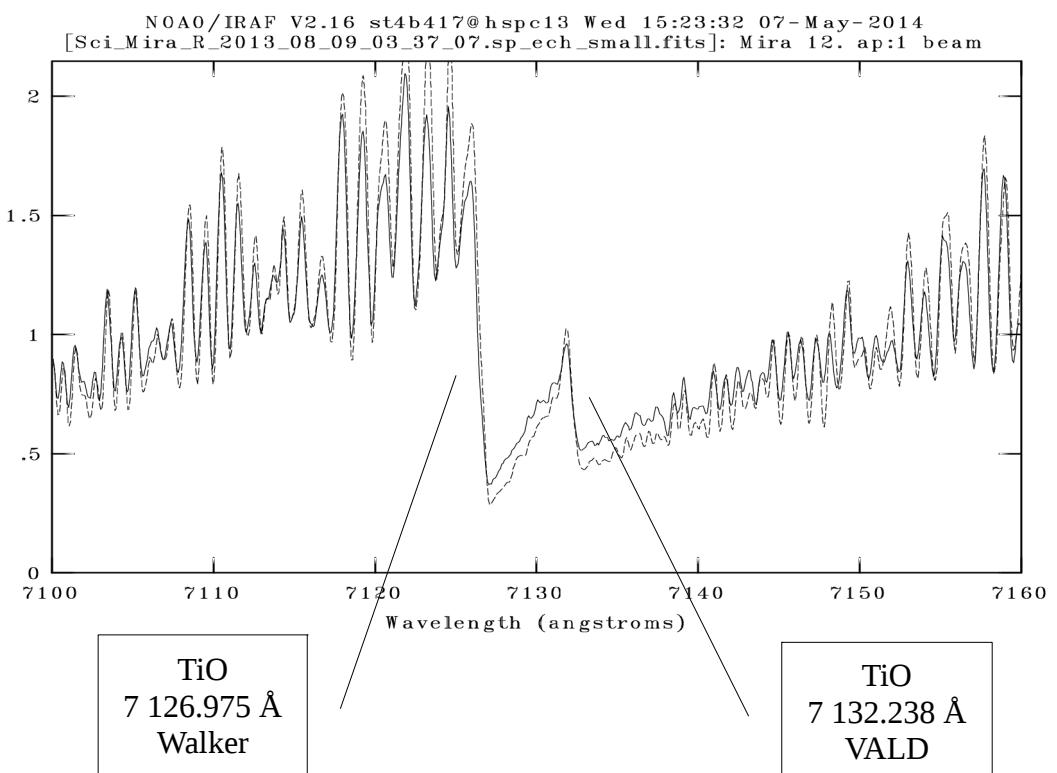


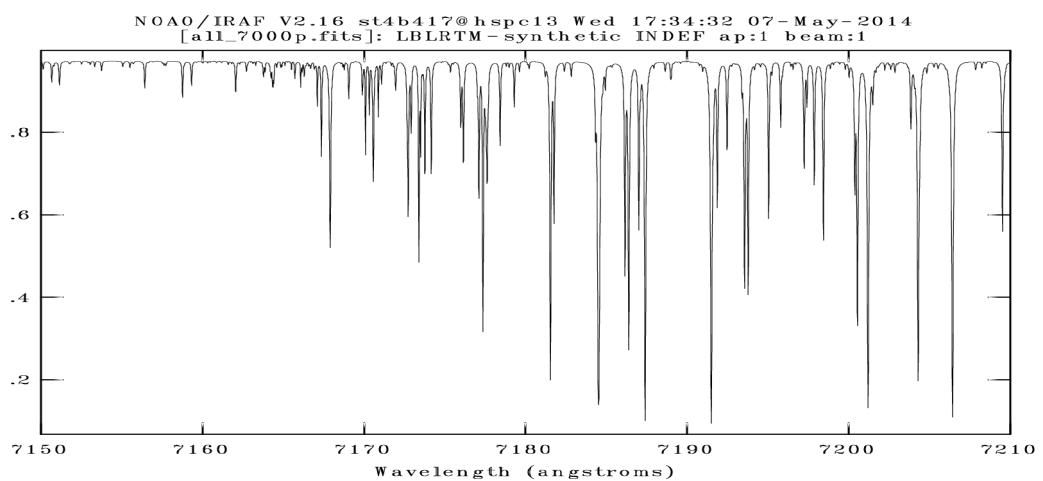
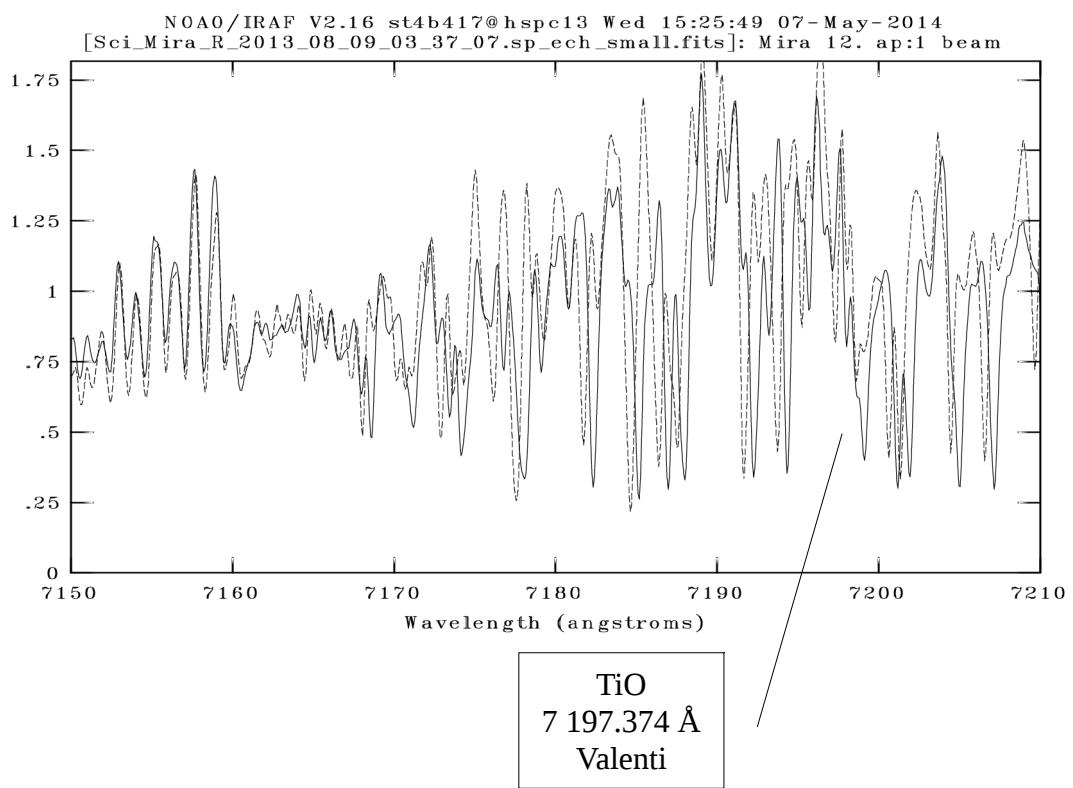
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[Sci_Mira_R_2013_08_09_03_37_07.sp_ech_small.fits]: Mira 12. ap:1 beam

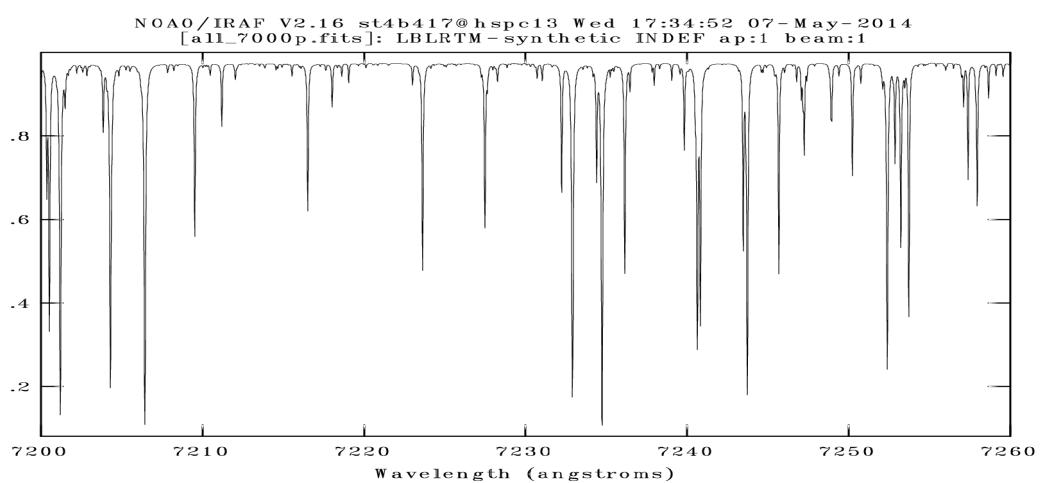
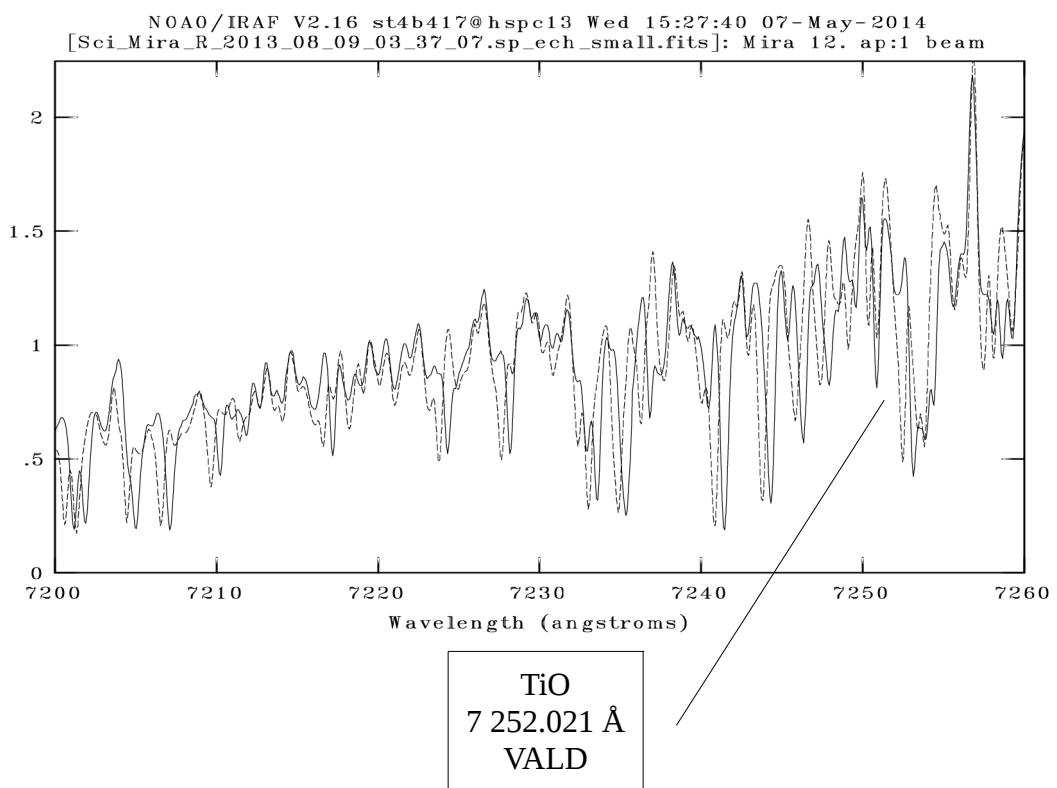


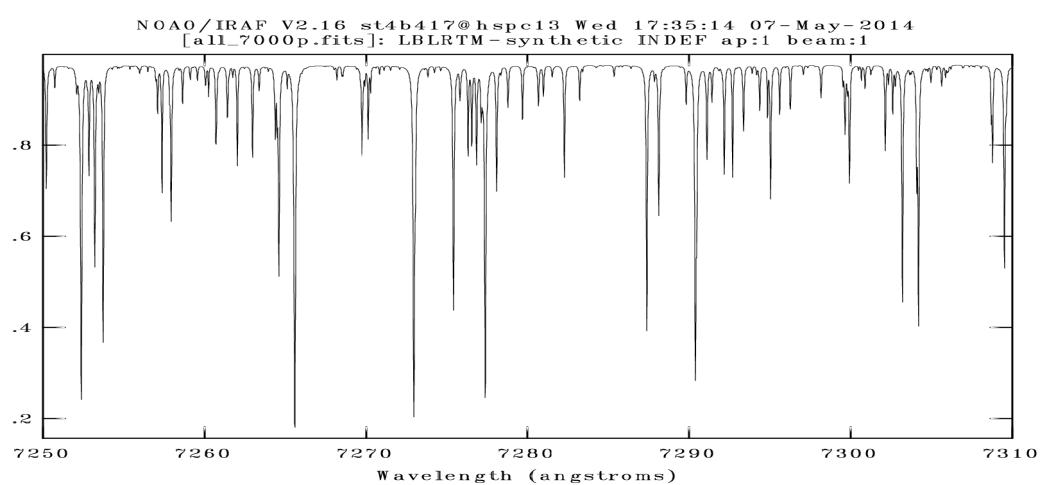
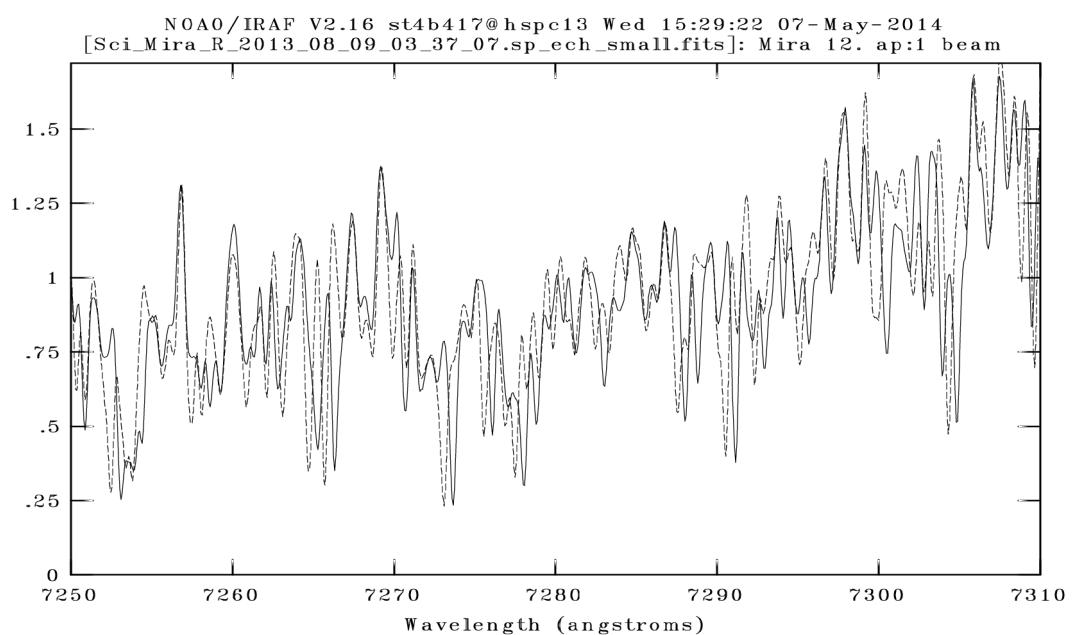
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[all_7000p.fits]: LBLRTM - synthetic INDEF ap:1 beam:1



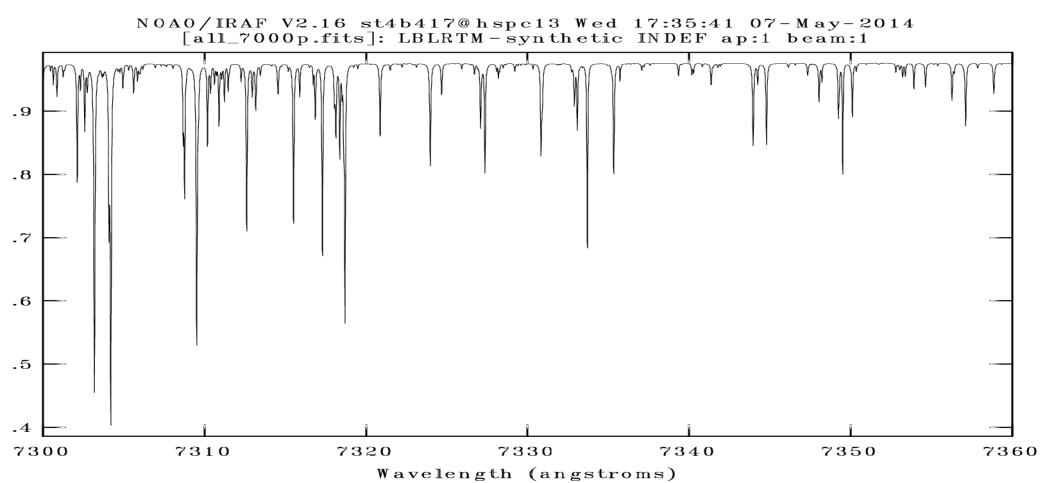
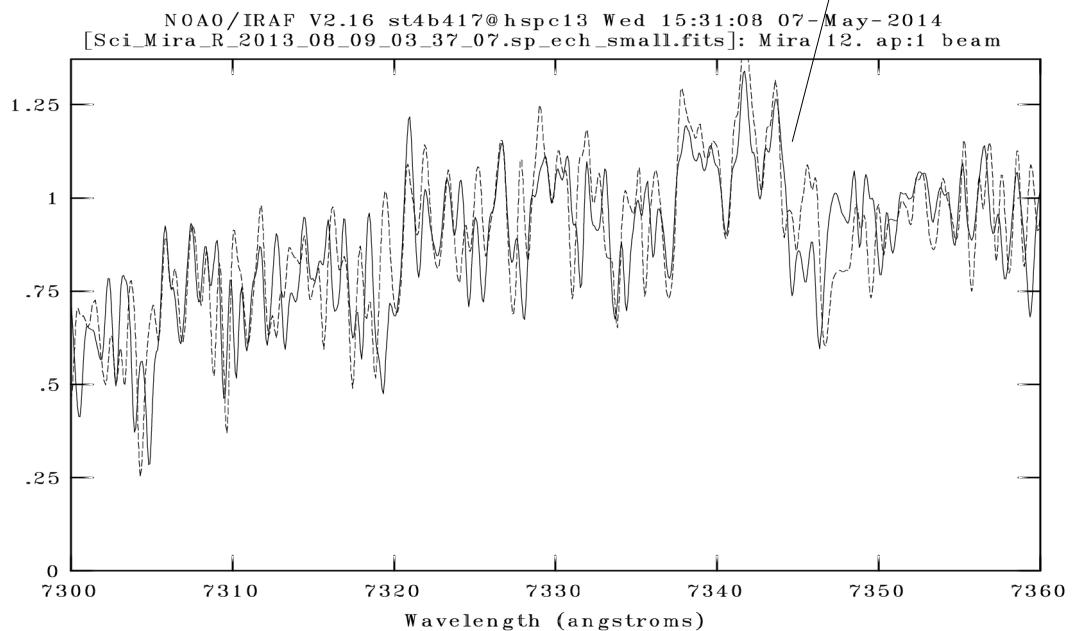


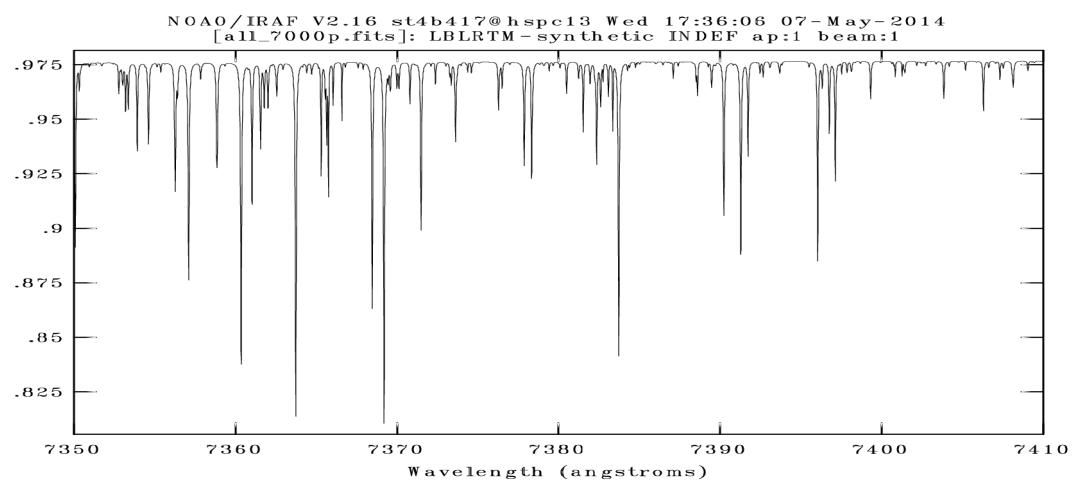
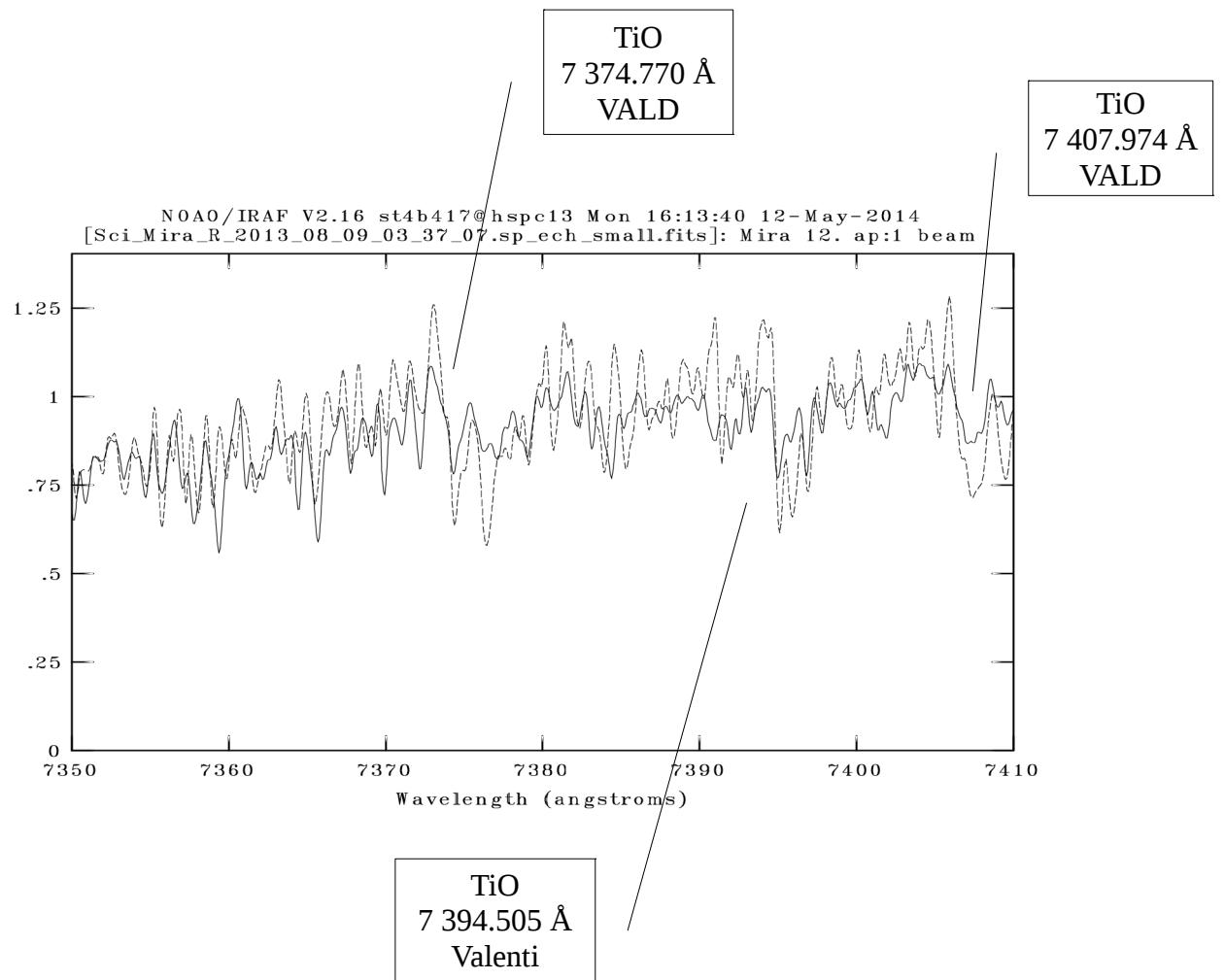


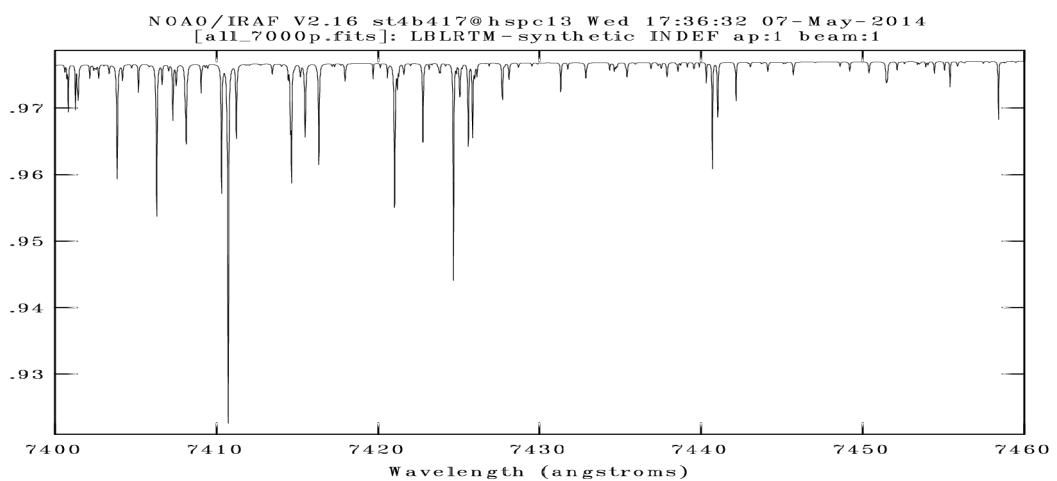
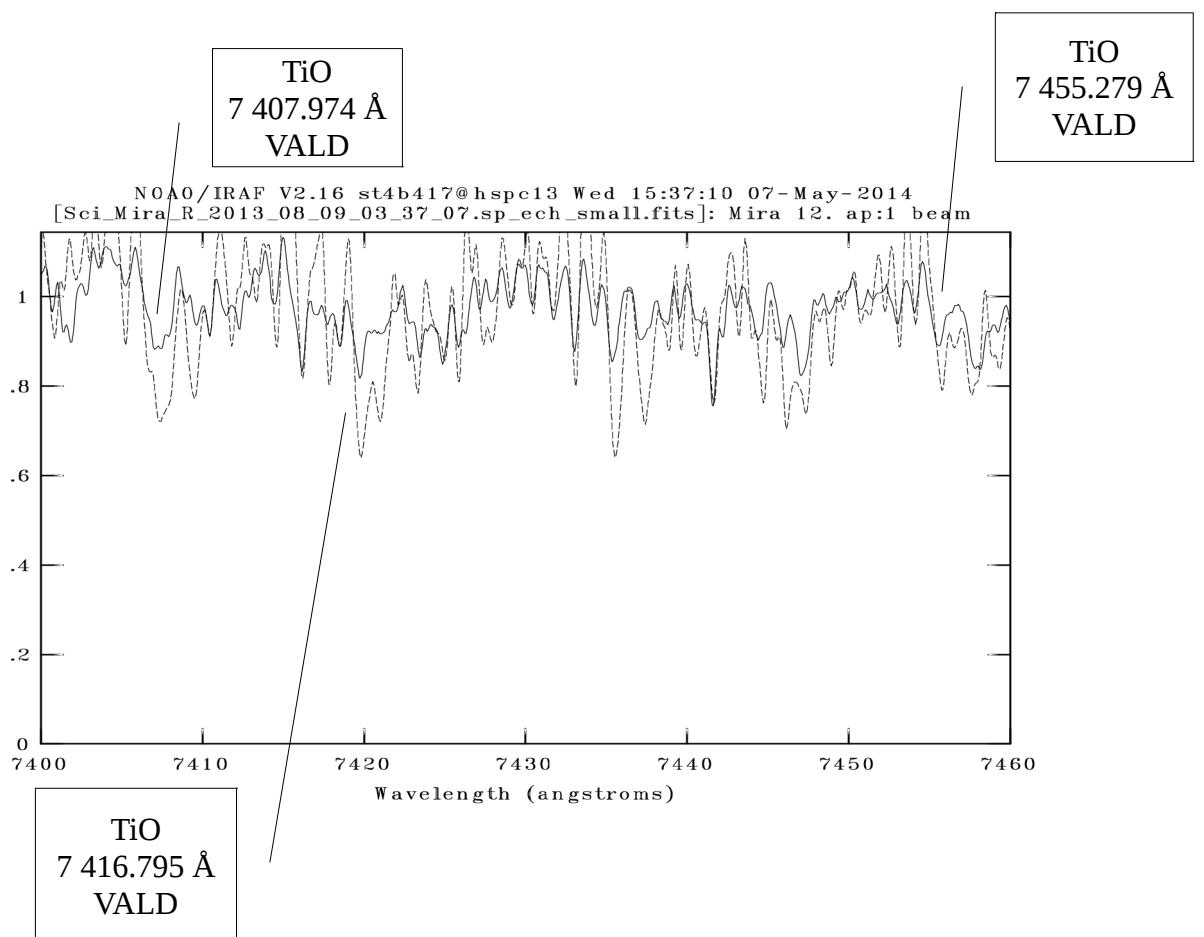


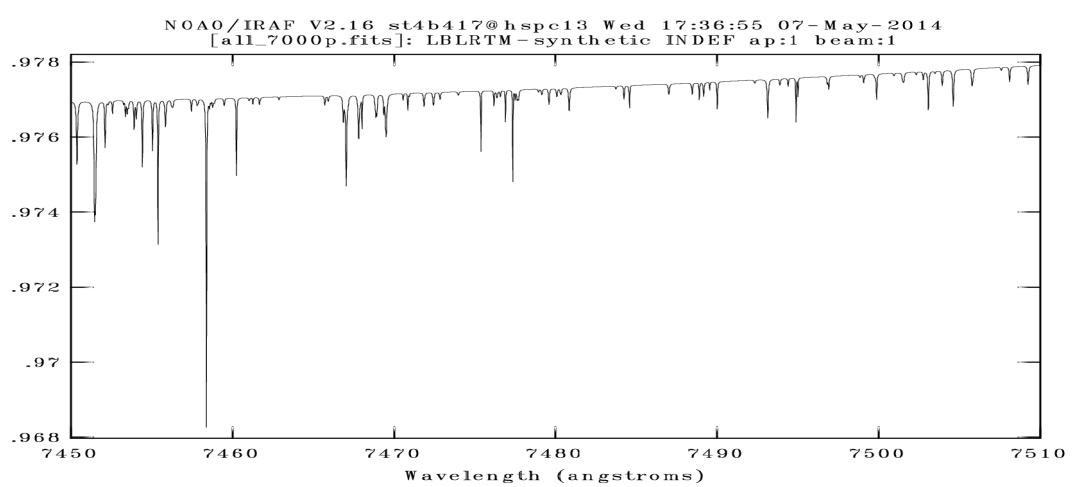
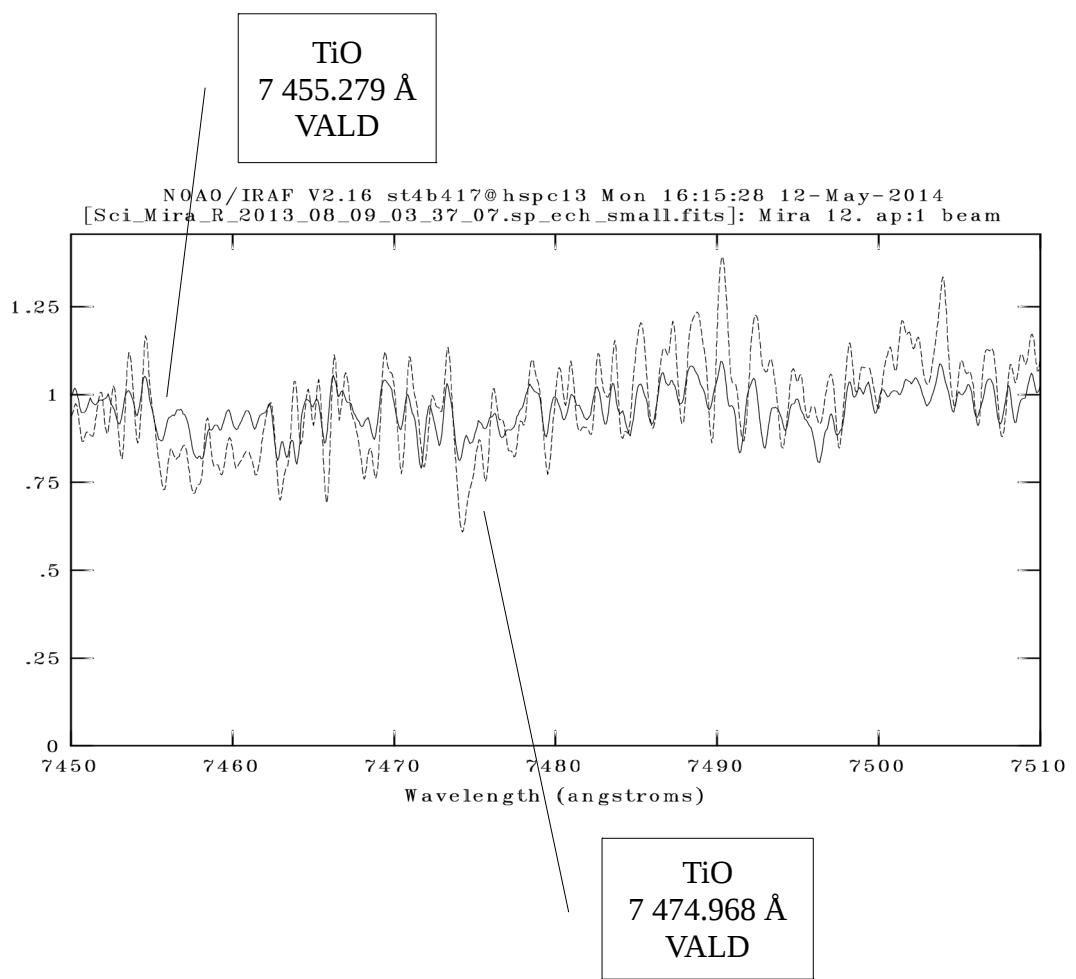


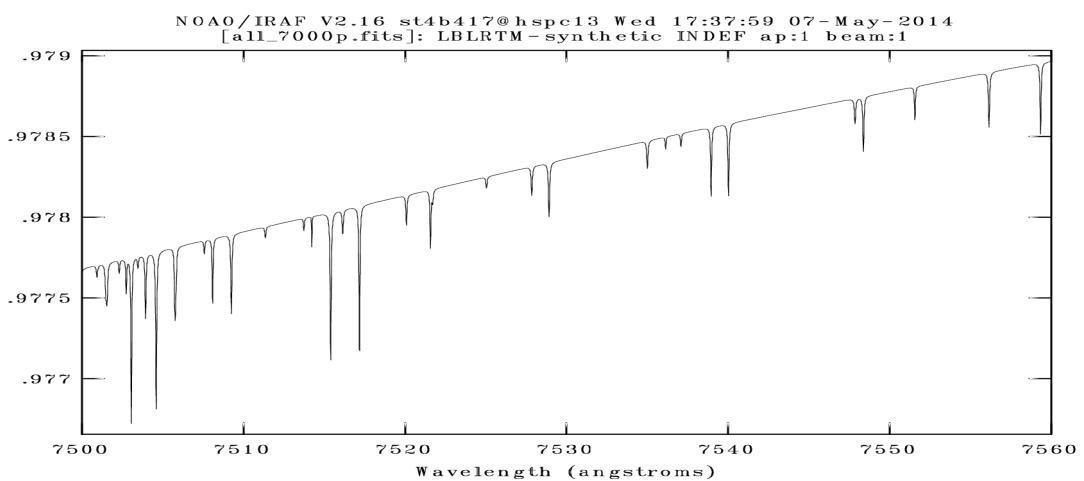
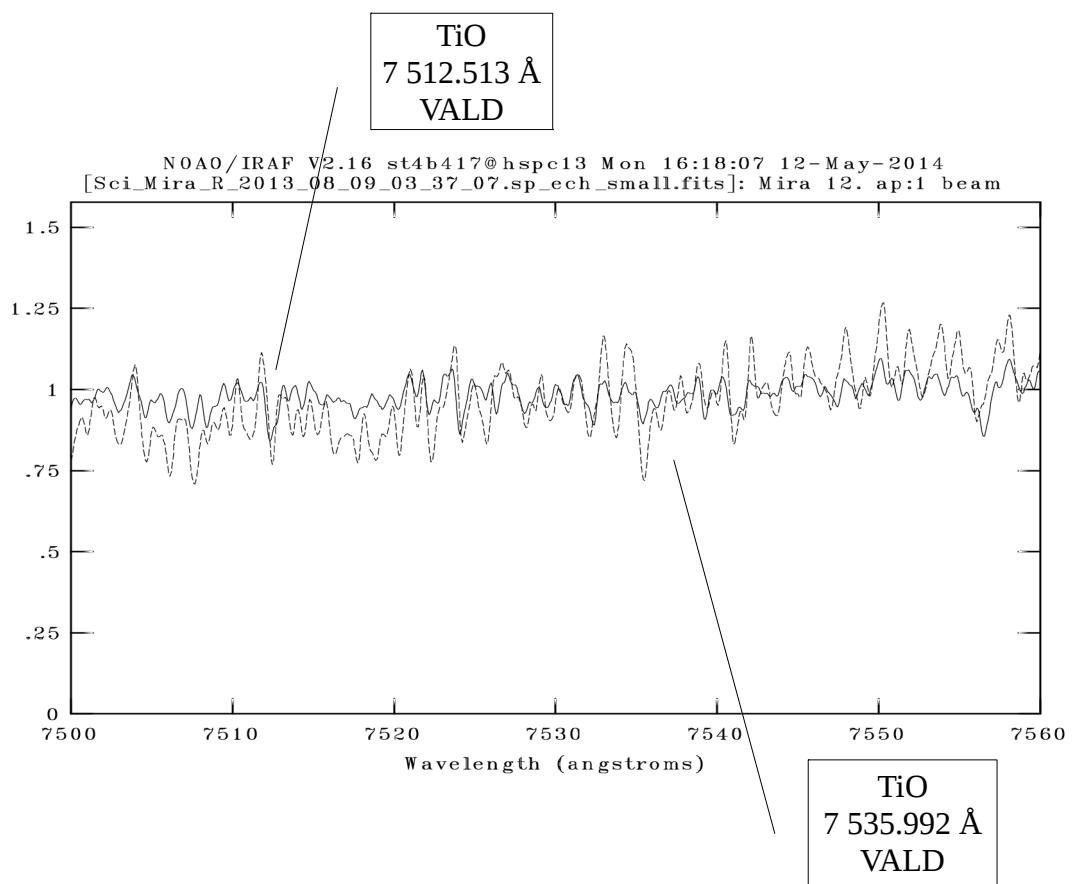
TiO
7344.710 Å
VALD



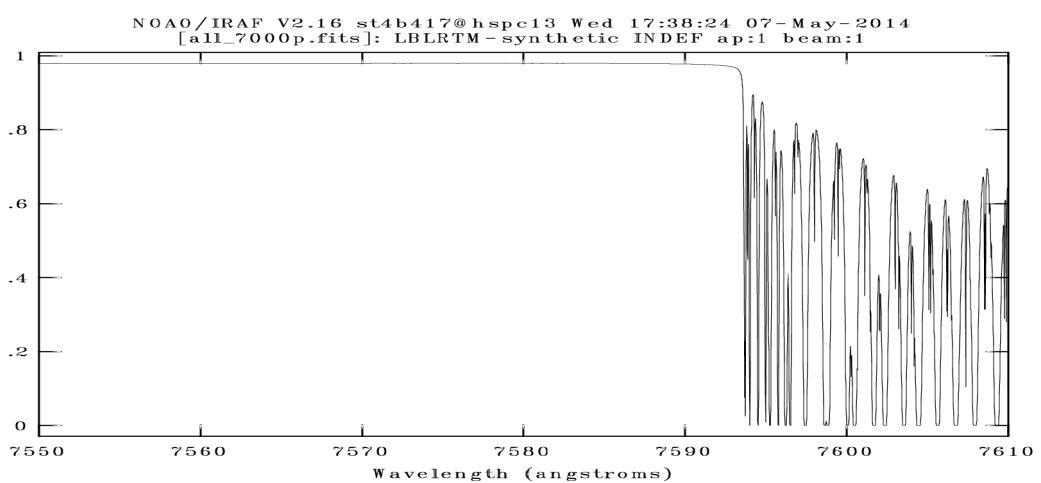
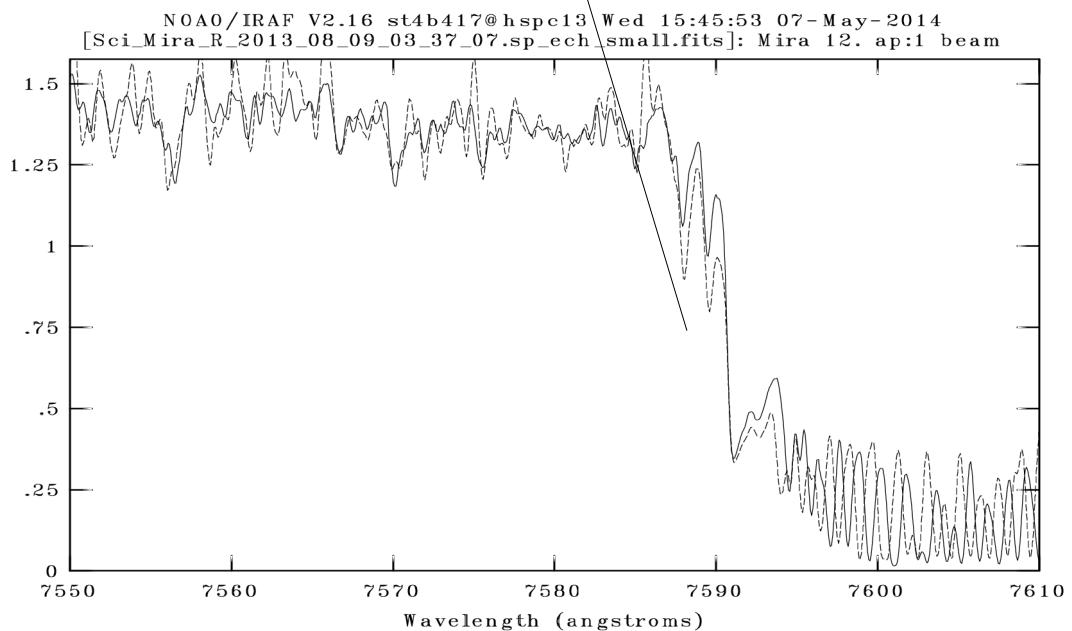


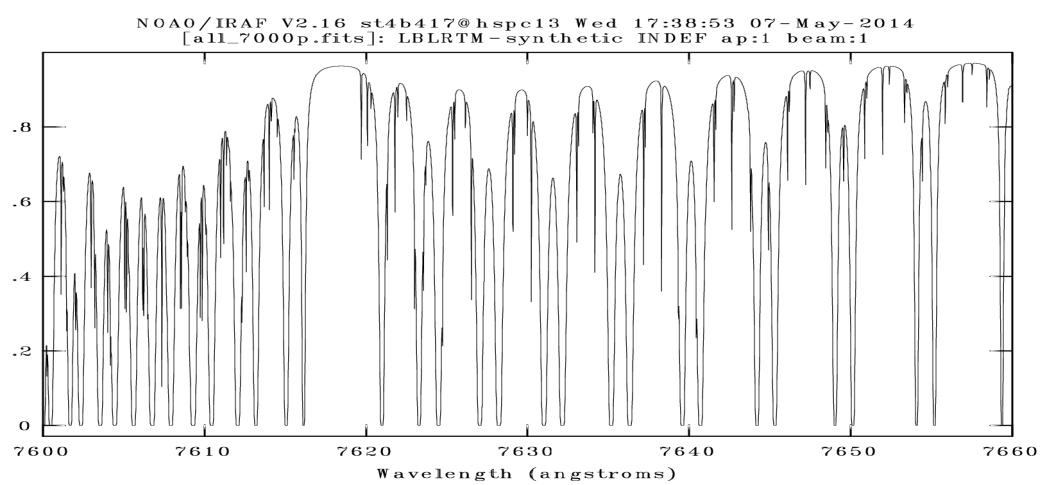
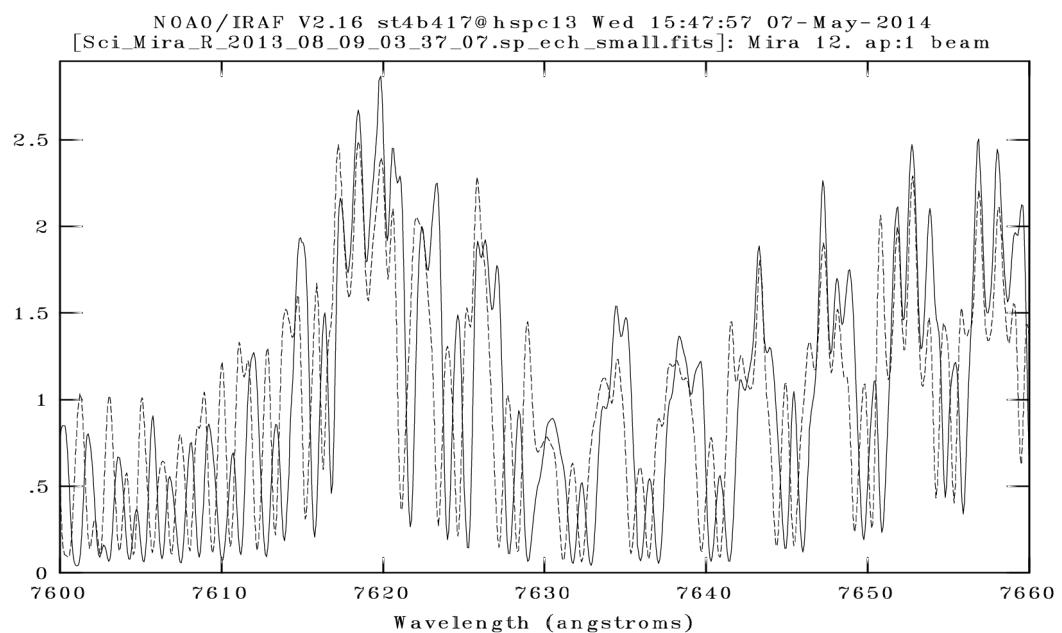


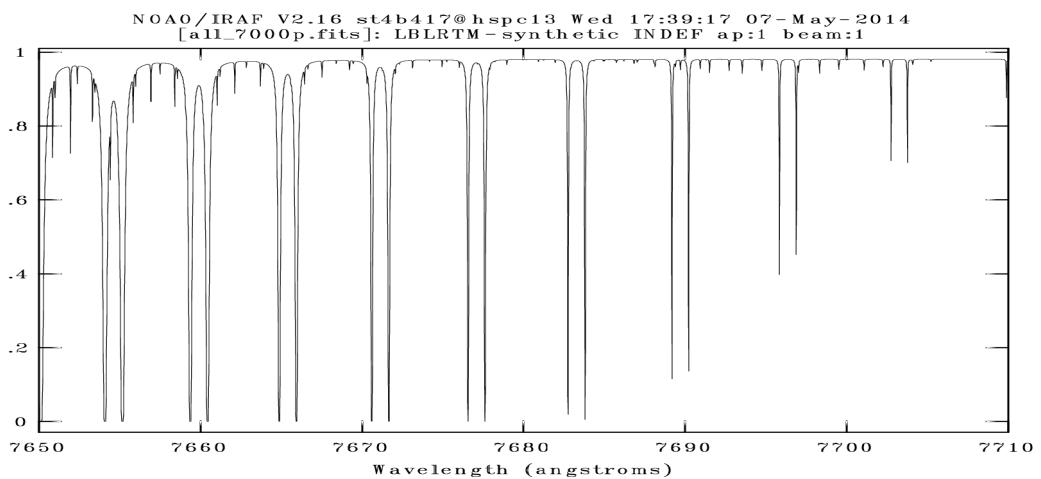
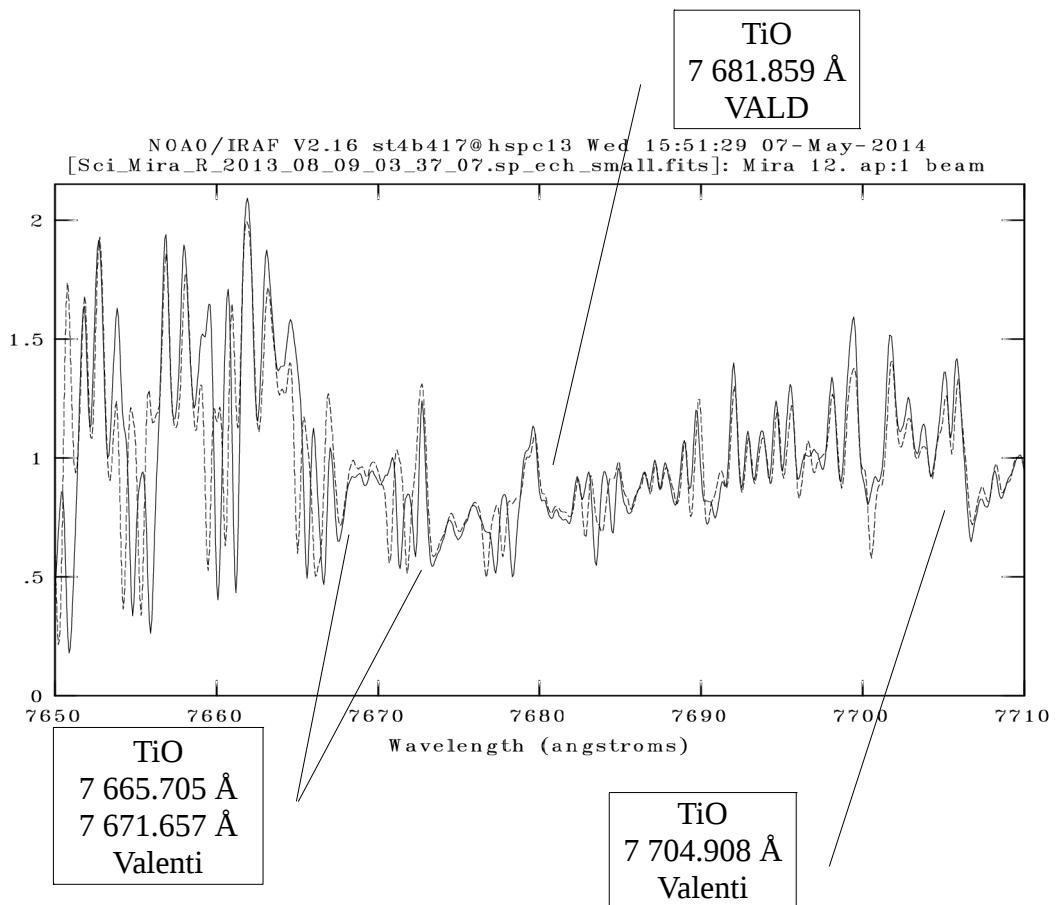


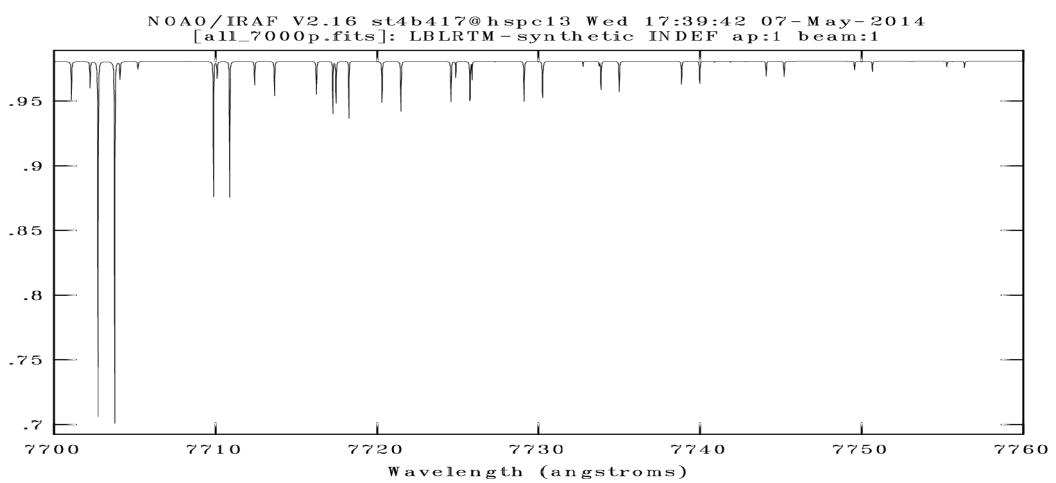
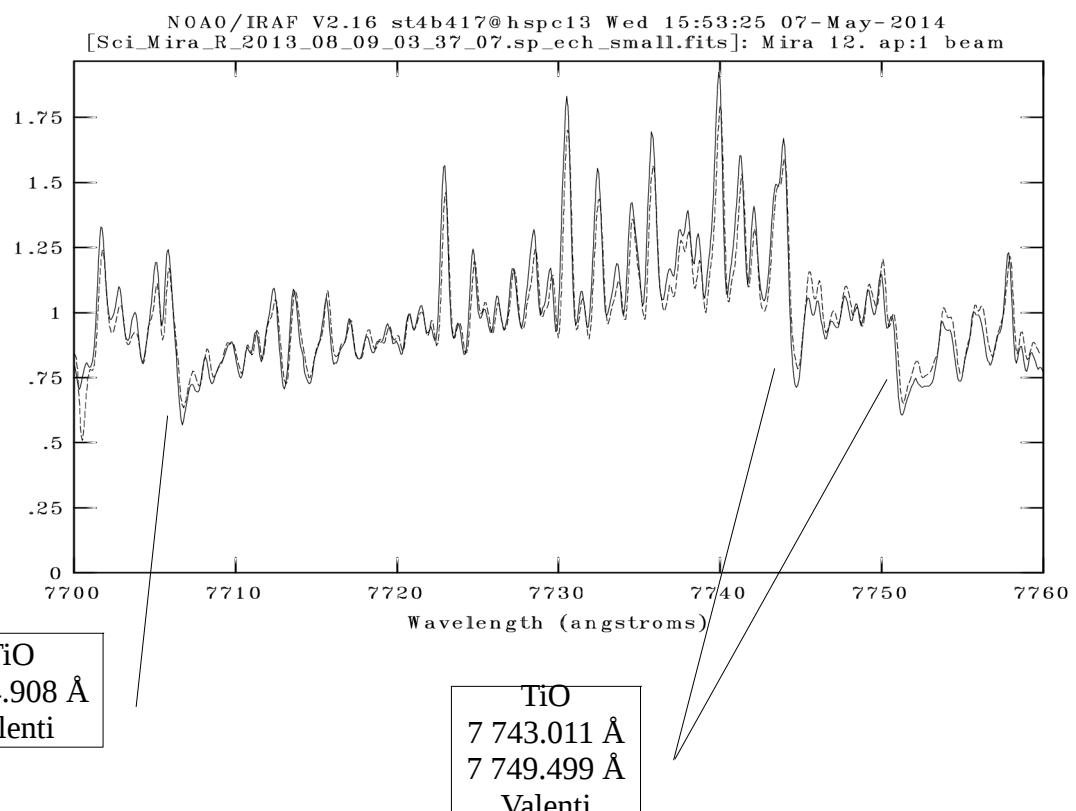


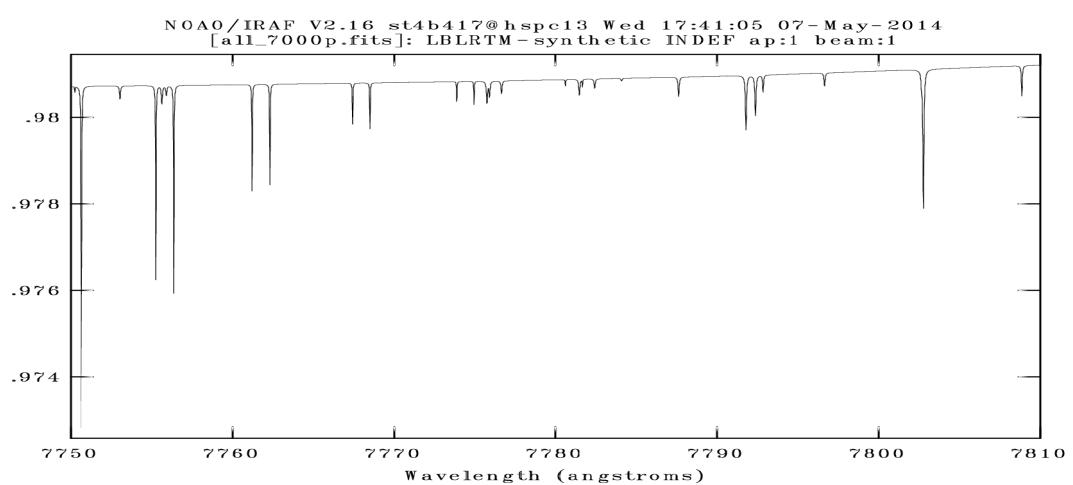
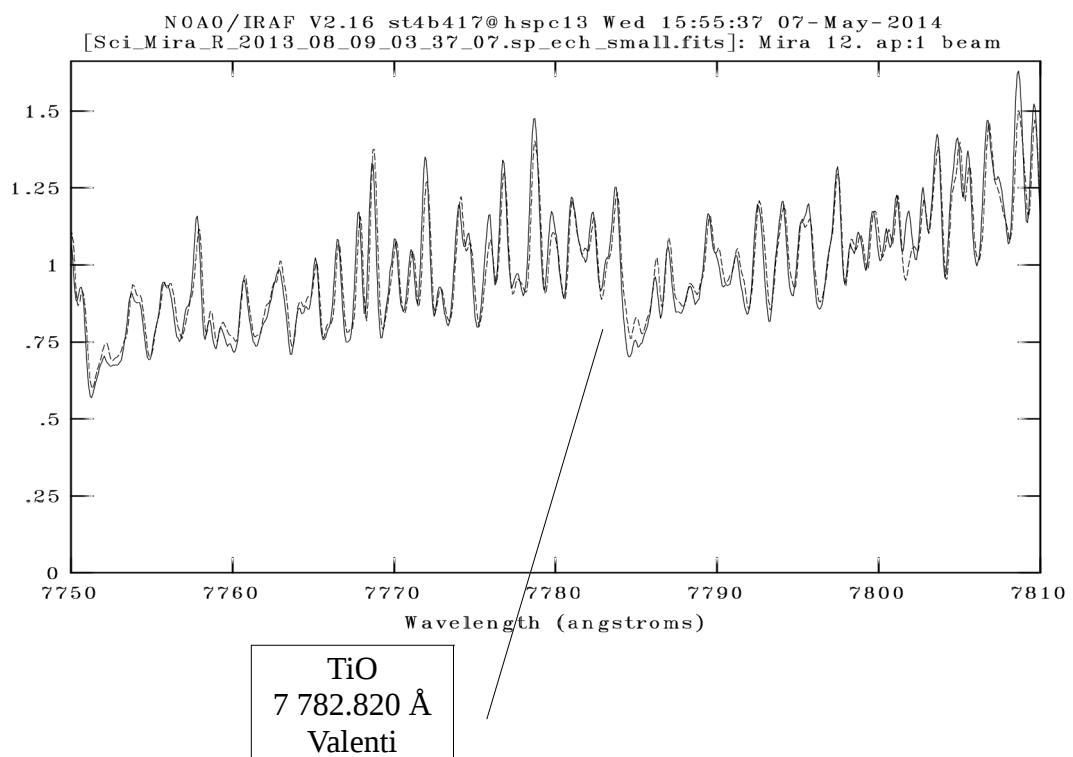
TiO
7 589.219 Å
Valenti



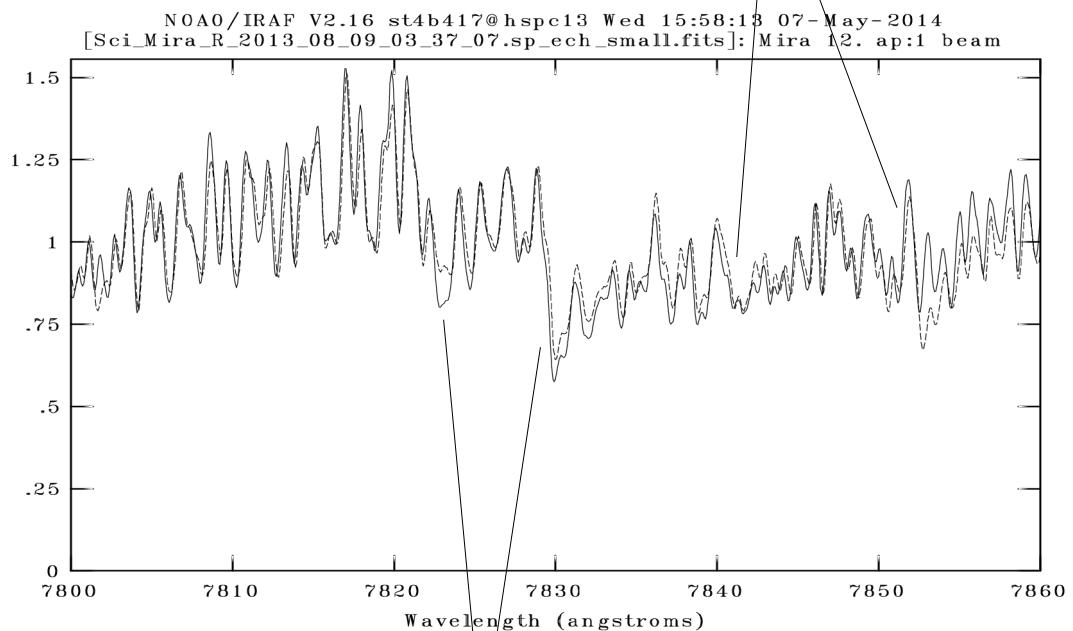




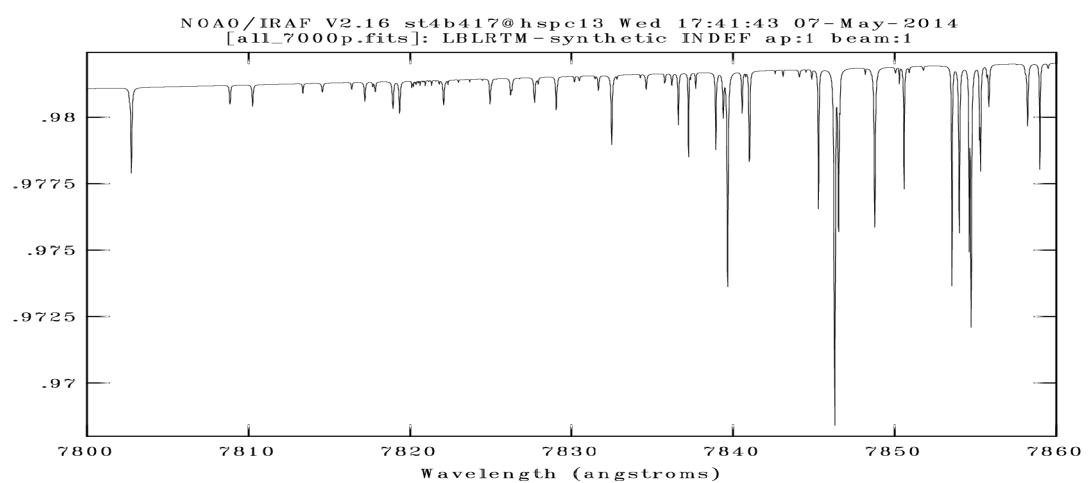


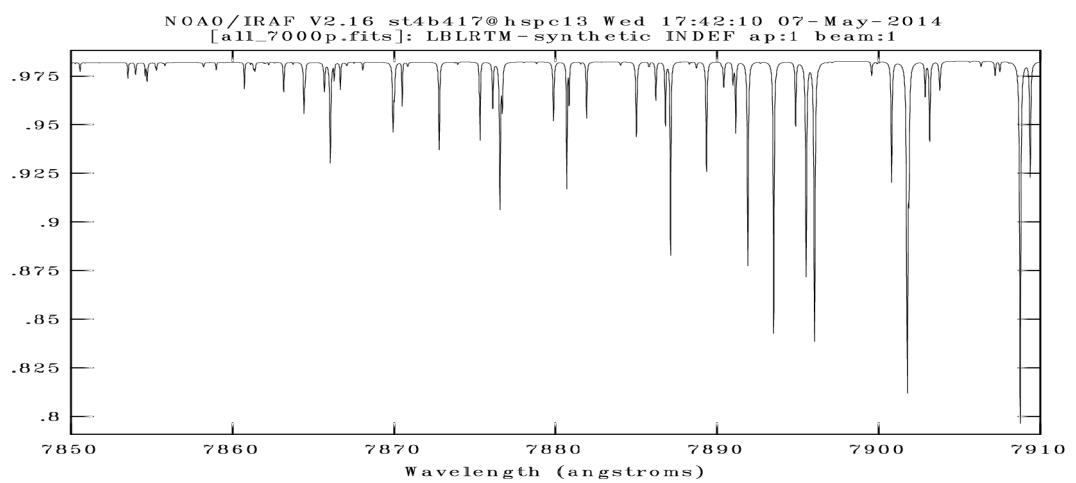
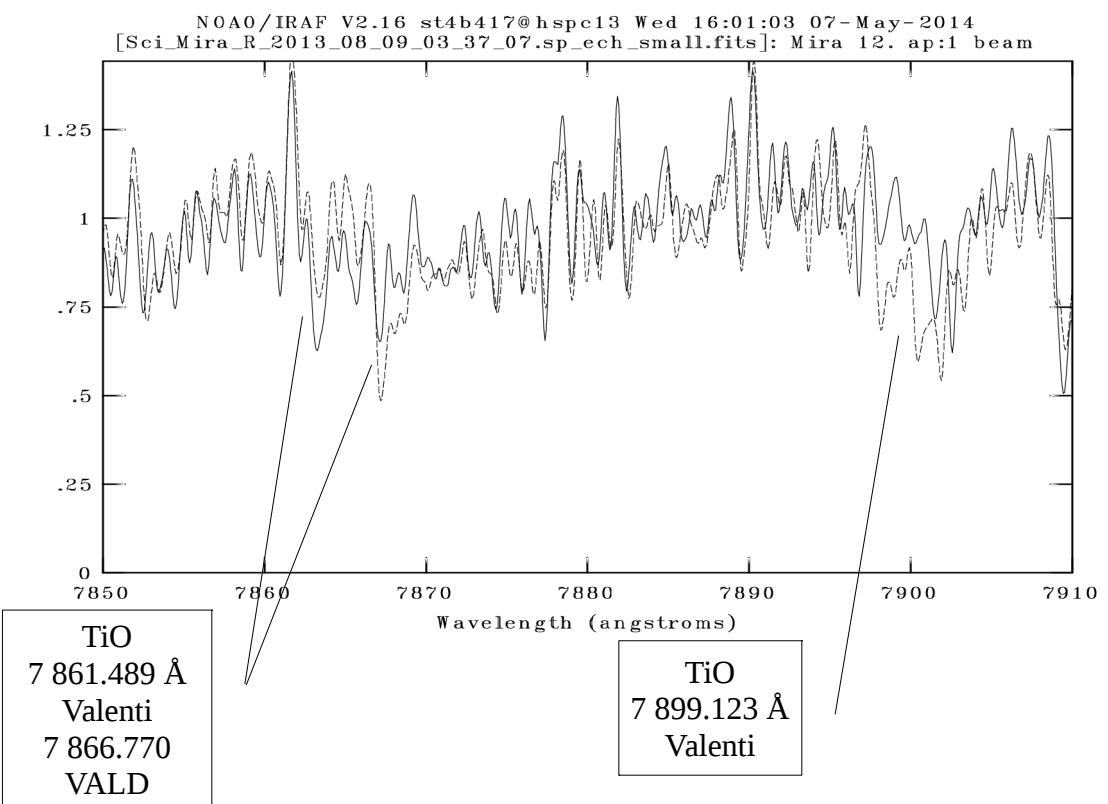


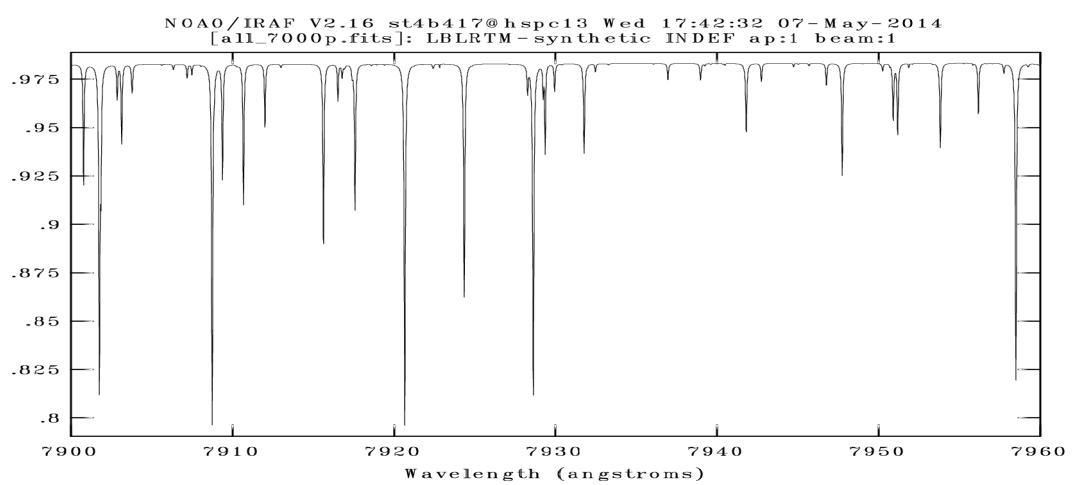
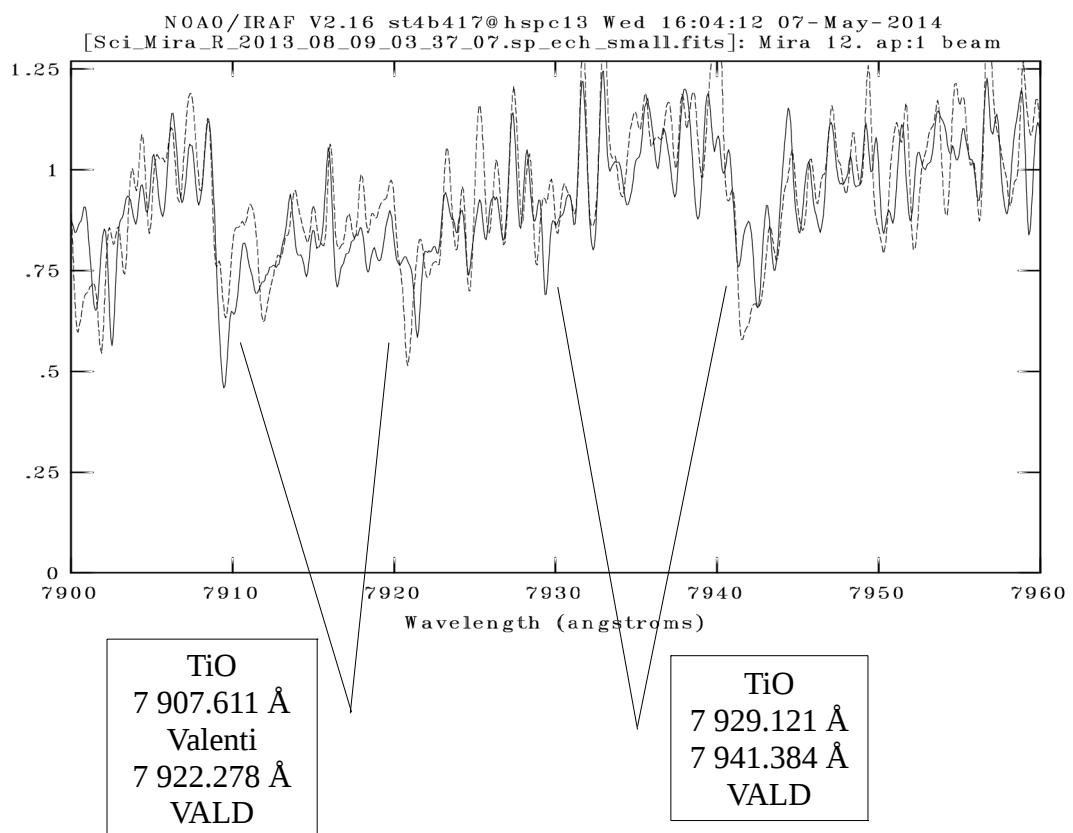
TiO
 7839.219 Å
 Valenti
 7853.335 Å
 VALD

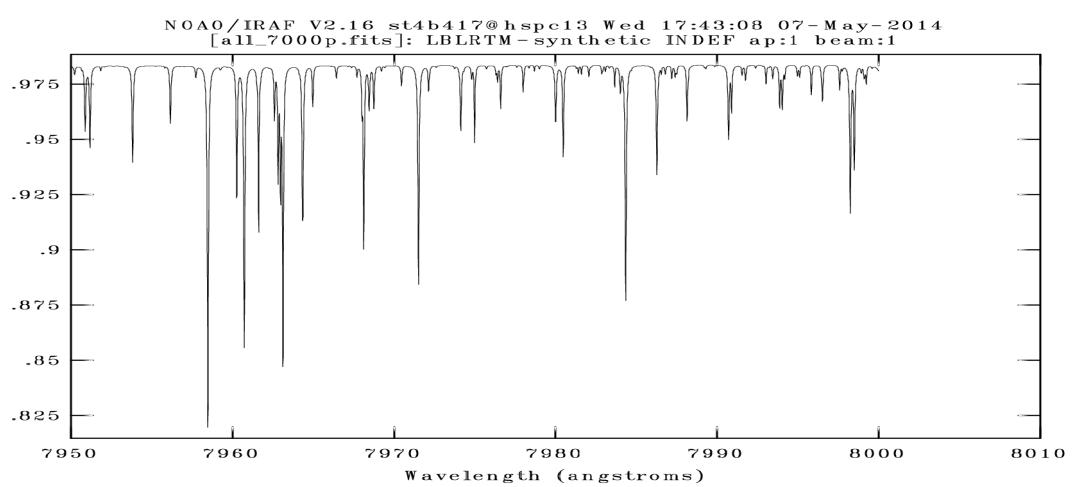
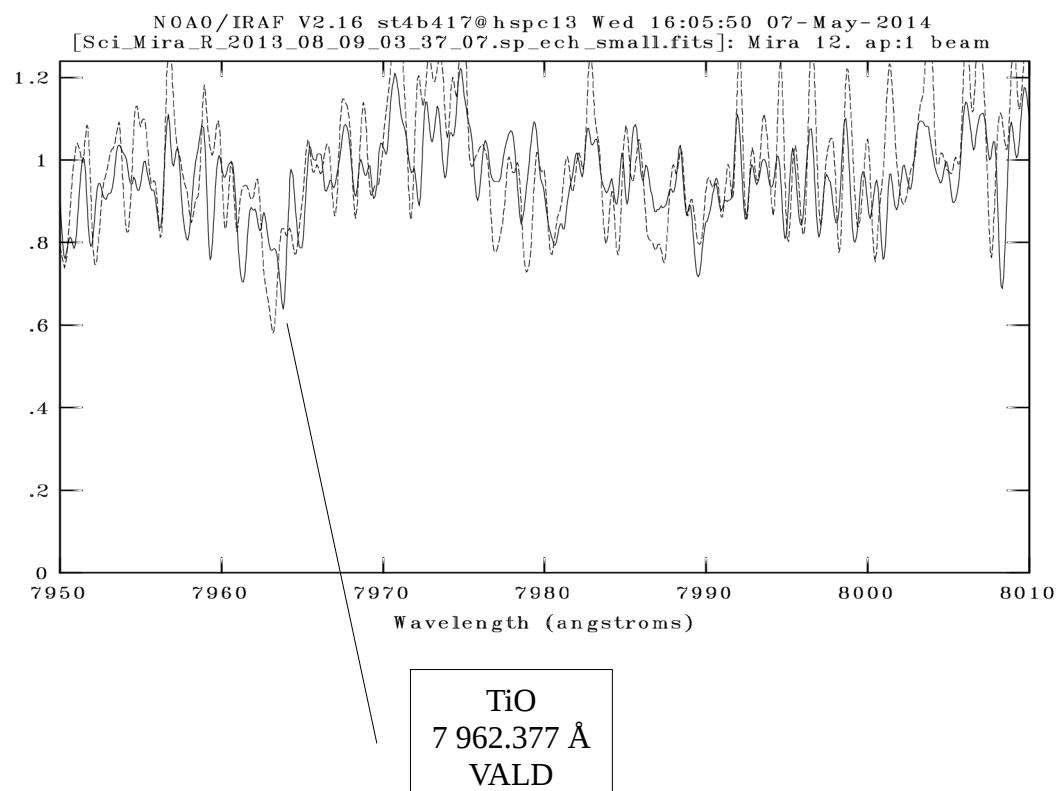


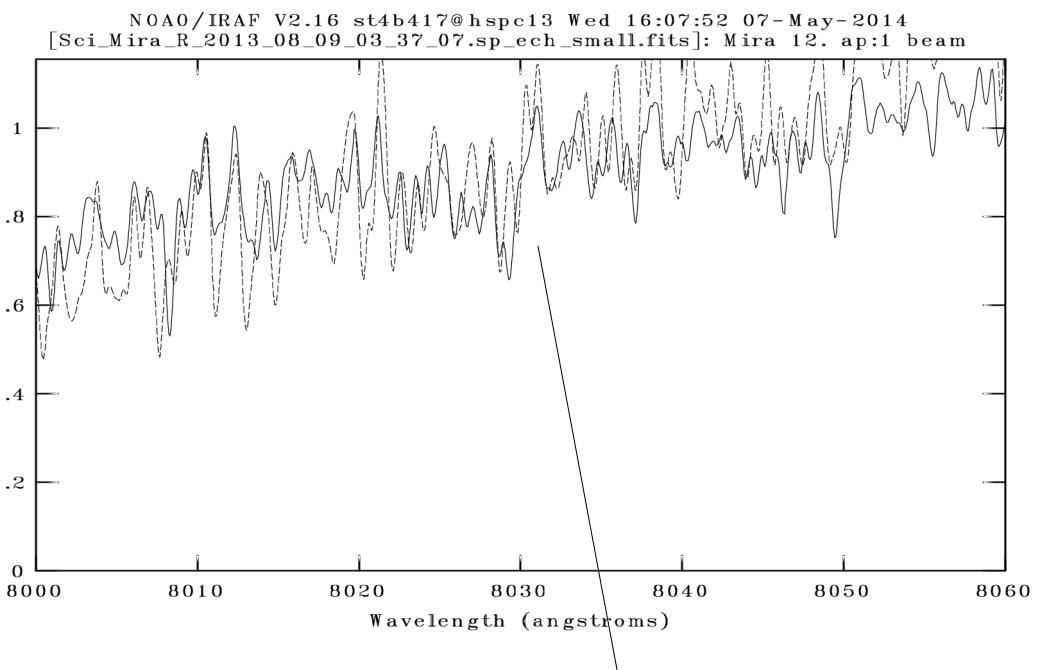
TiO
 7821.000 Å
 7828.142 Å
 Valenti



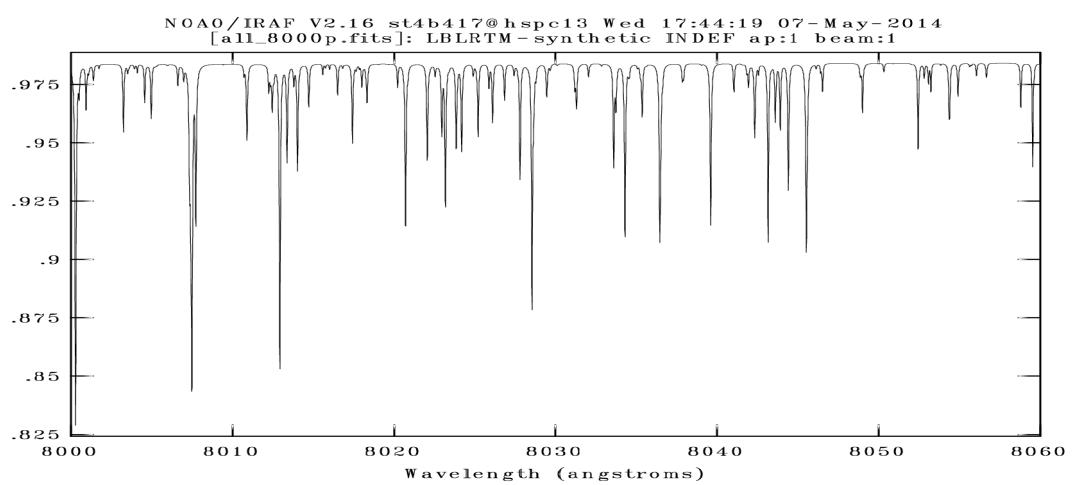


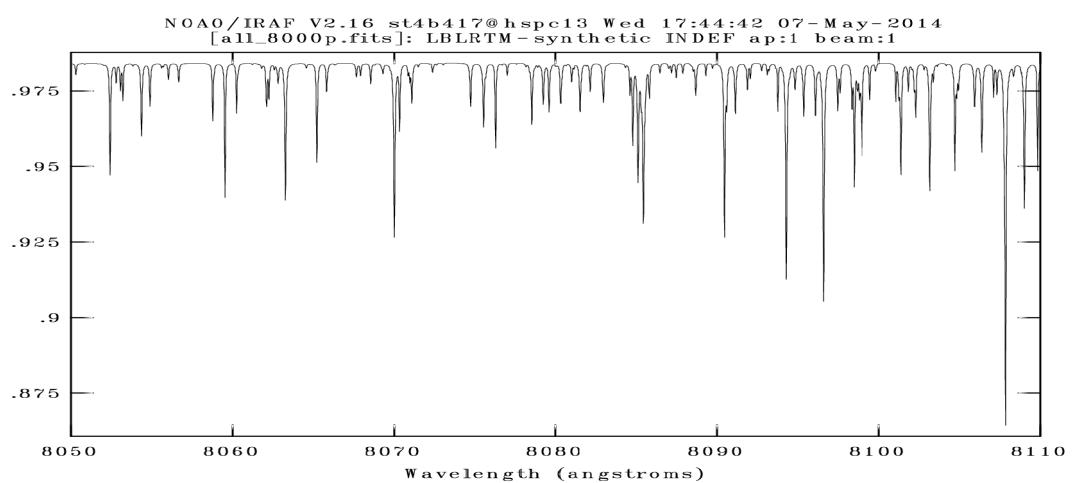
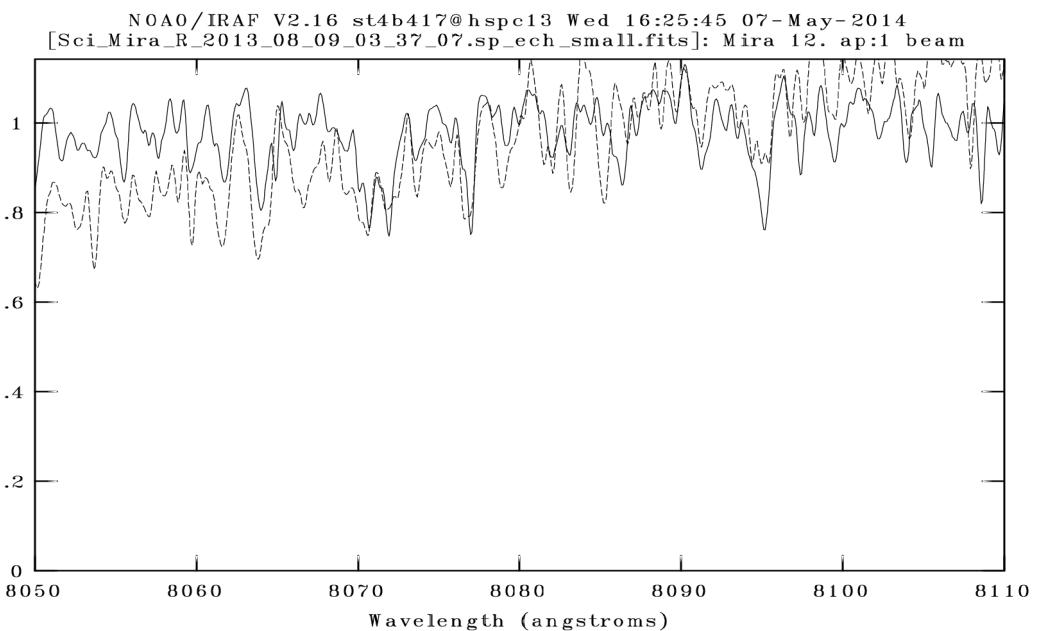




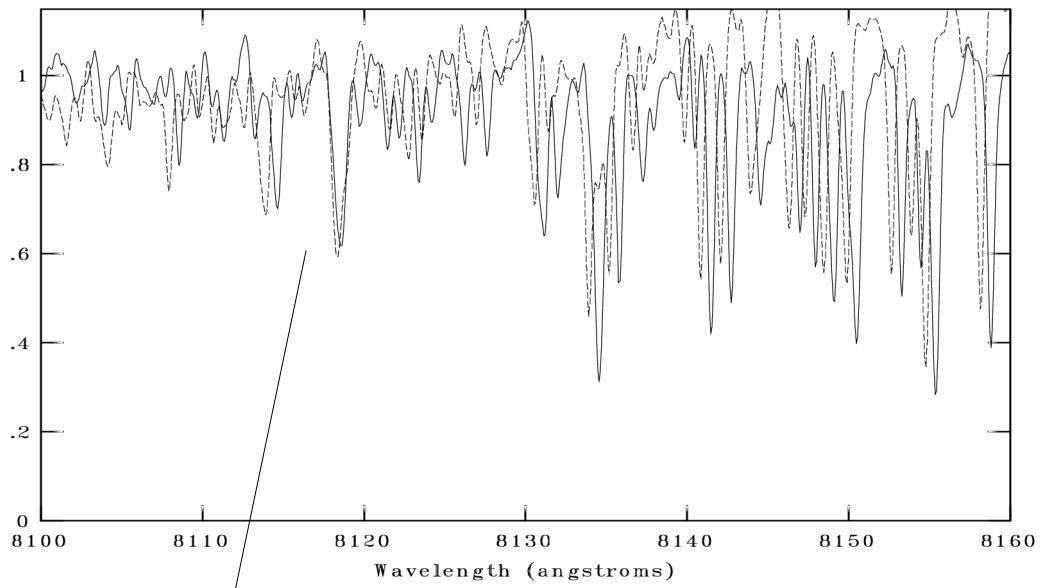


TiO
8 029.268 Å
VALID



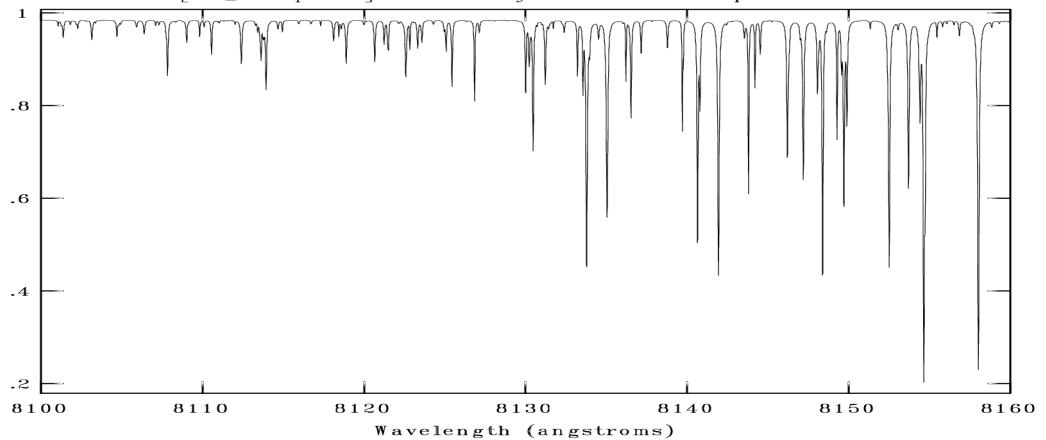


NOAO/IRAF V2.16 st4b417@hspc13 Wed 16:27:52 07-May-2014
[Sci_Mira_R_2013_08_09_03_37_07.sp_ech_small.fits]: Mira 12. ap:1 beam

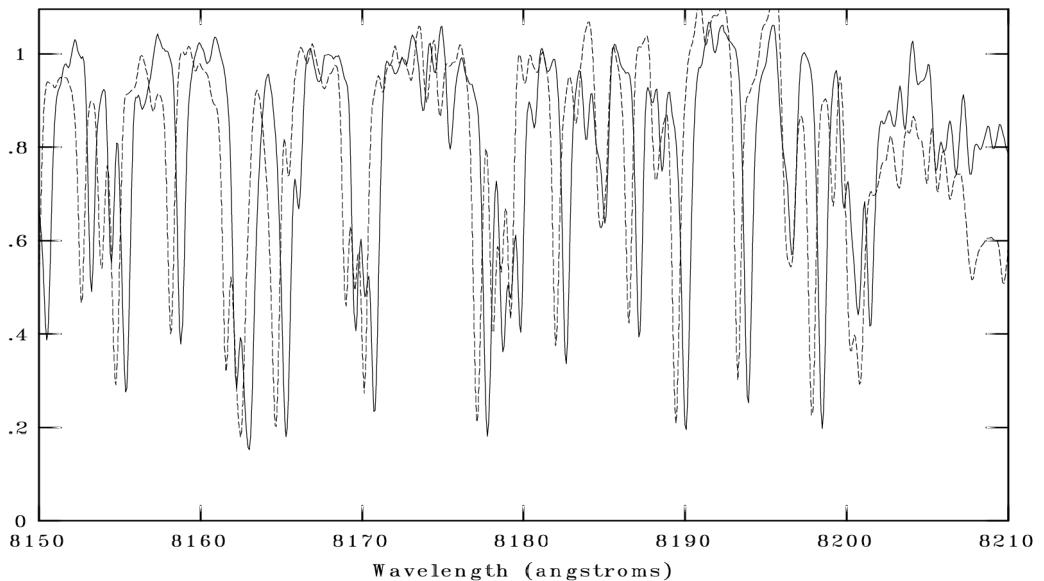


TiO
8 118.961 Å
VALID

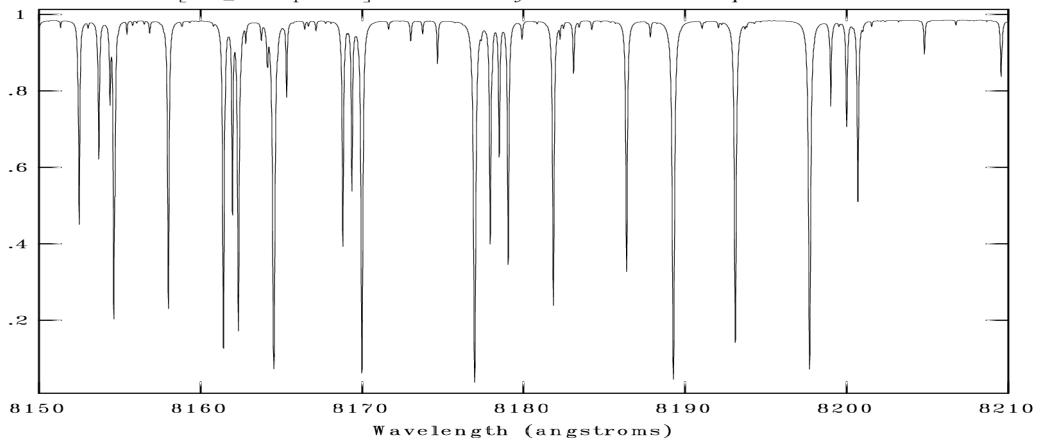
NOAO/IRAF V2.16 st4b417@hspc13 Wed 17:45:08 07-May-2014
[all_8000p.fits]: LBLRTM - synthetic INDEF ap:1 beam:1



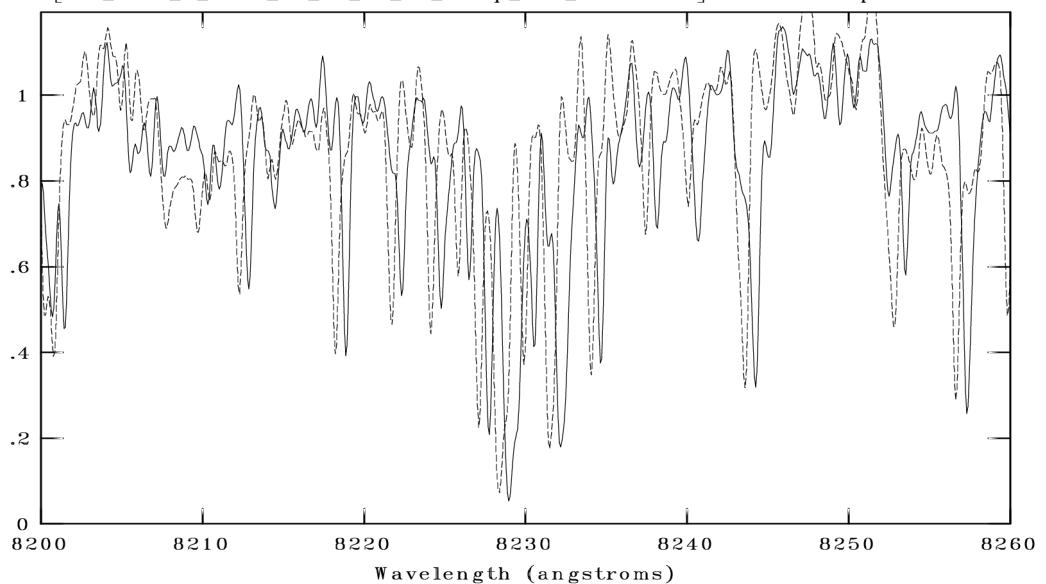
NOAO/IRAF V2.16 st4b417@hspc13 Wed 16:30:26 07-May-2014
[Sci_Mira_R_2013_08_09_03_37_07.sp_ech_small.fits]: Mira 12. ap:1 beam



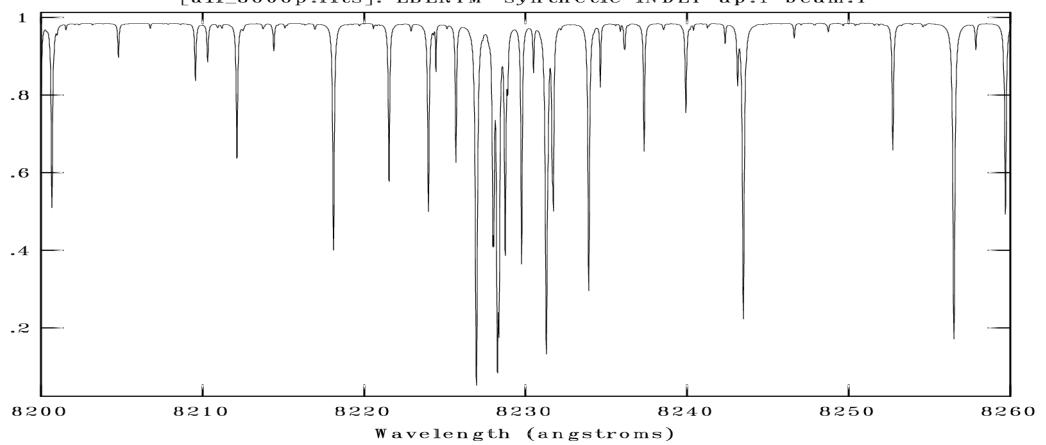
NOAO/IRAF V2.16 st4b417@hspc13 Wed 17:45:34 07-May-2014
[all_8000p.fits]: LBLRTM - synthetic INDEF ap:1 beam:1



NOAO/IRAF V2.16 st4b417@hspc13 Wed 16:32:11 07-May-2014
[Sci_Mira_R_2013_08_09_03_37_07.sp_ech_small.fits]: Mira 12. ap:1 beam



NOAO/IRAF V2.16 st4b417@hspc13 Wed 17:45:59 07-May-2014
[all_8000p.fits]: LBLRTM - synthetic INDEF ap:1 beam:1



Fe I
8 291.764 Å
NIST

